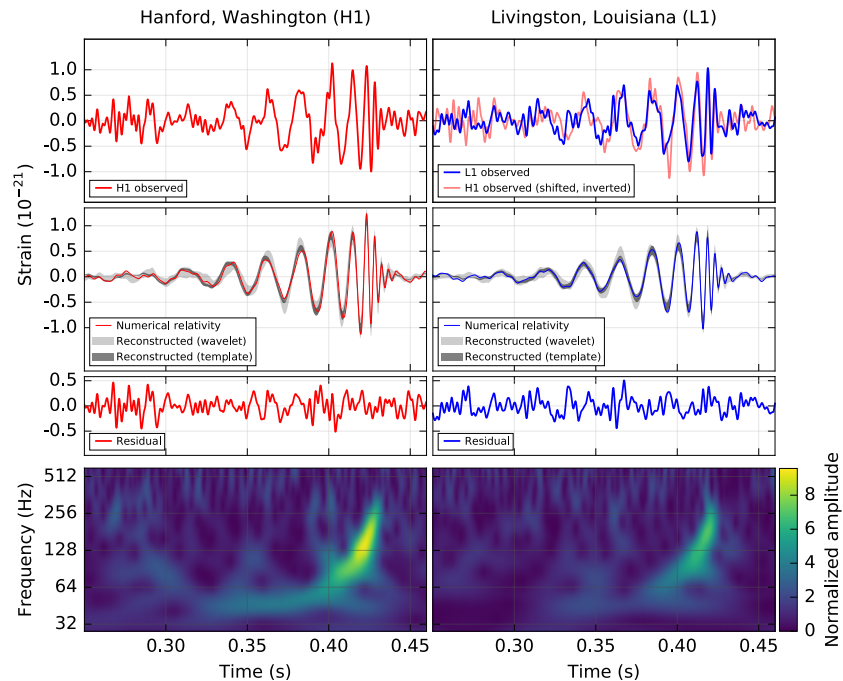




Methods for calibrating kilometer-scale interferometers

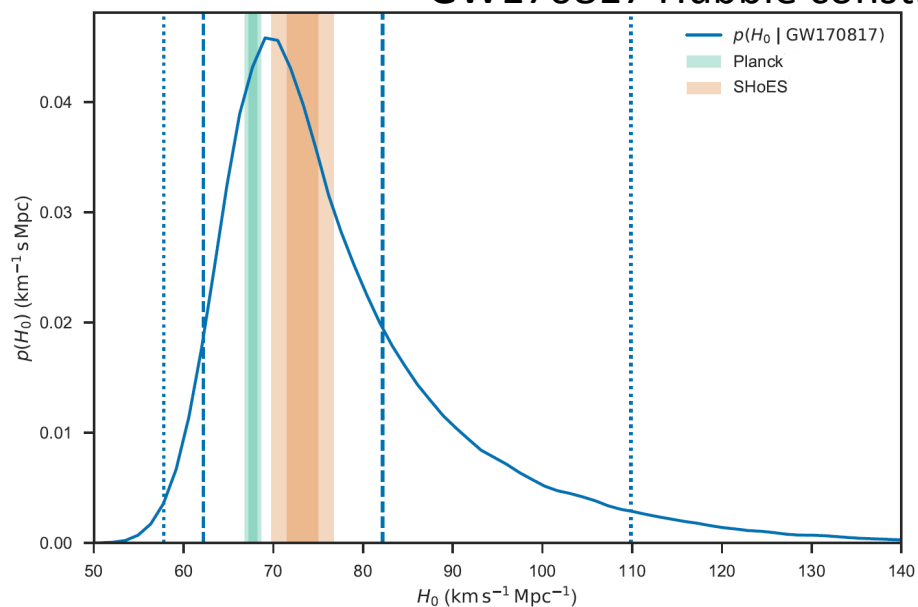
Evan Goetz
University of Michigan

Gravitational wave Metrology Workshop
NIST, Boulder, CO

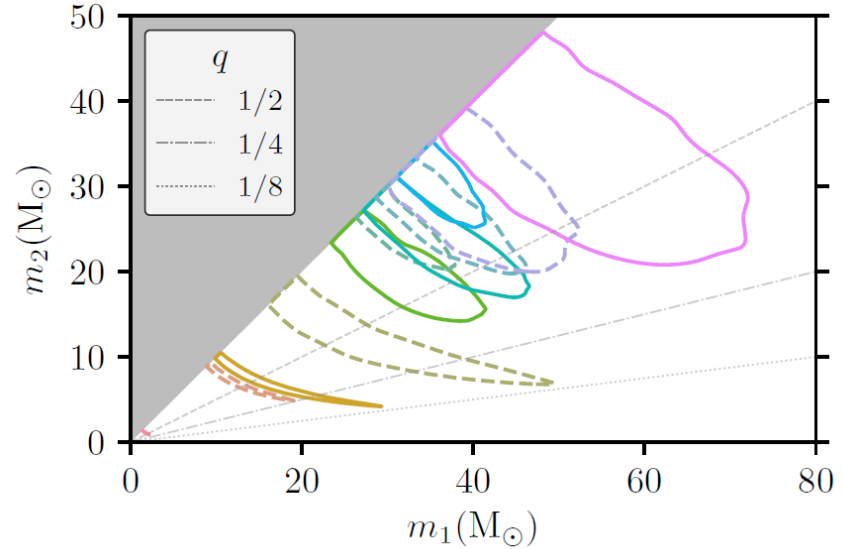


GW150914 strain

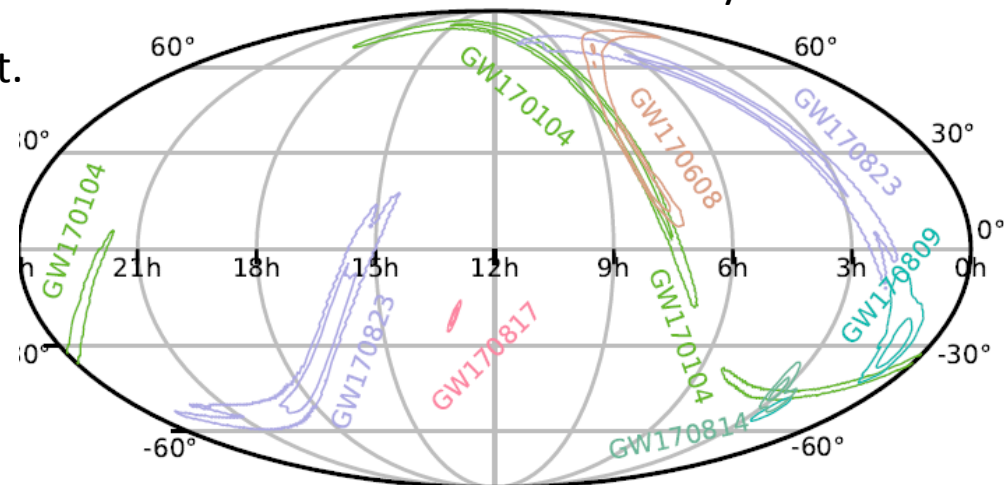
GW170817 Hubble const.



O2 events masses



O2 events sky locations

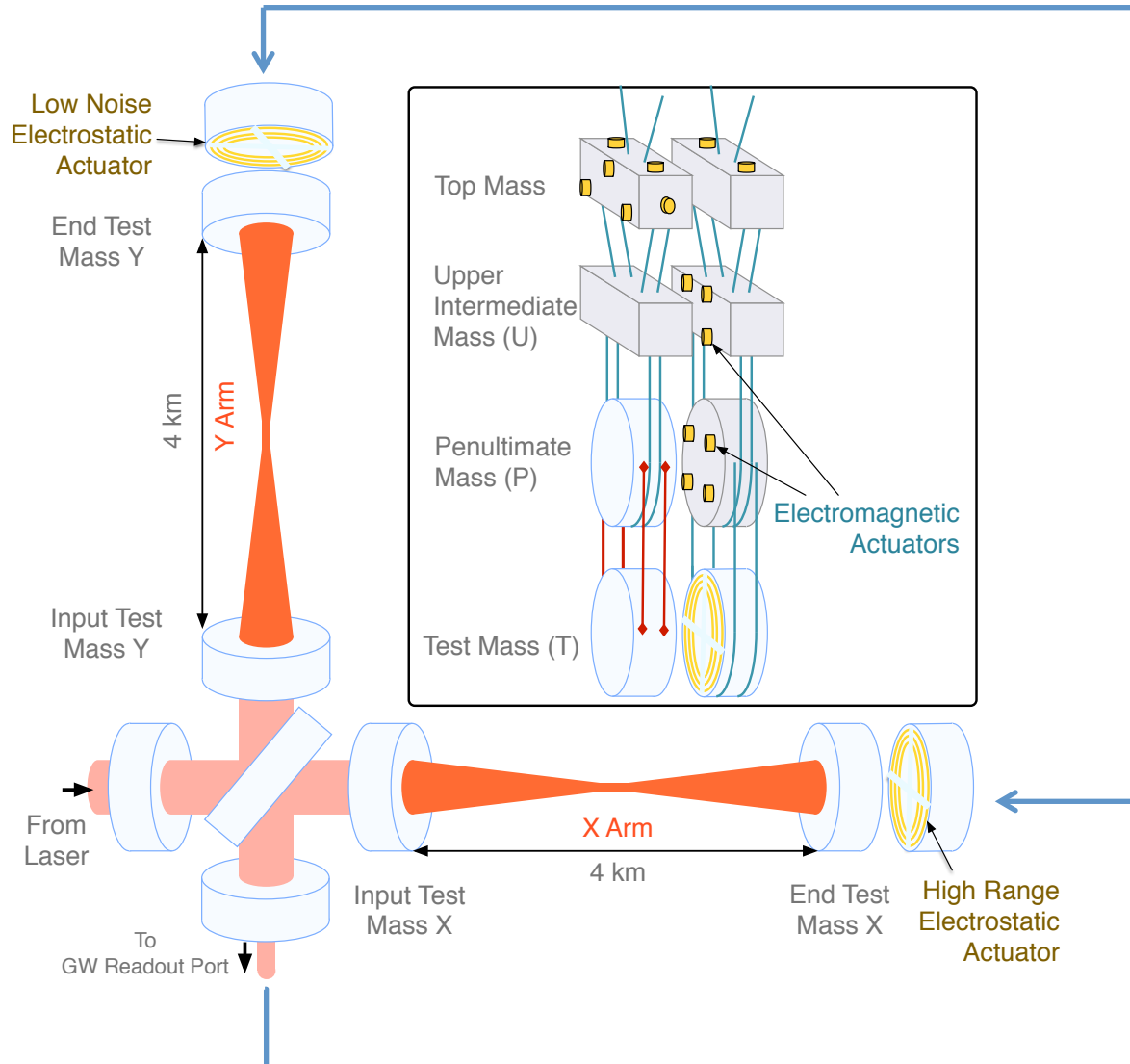


Abbott, B P, et al. PRL **116** 061102
 LVC, et al. Nature **551** 85-88
 arXiv:1811.12907 [astro-ph.HE]

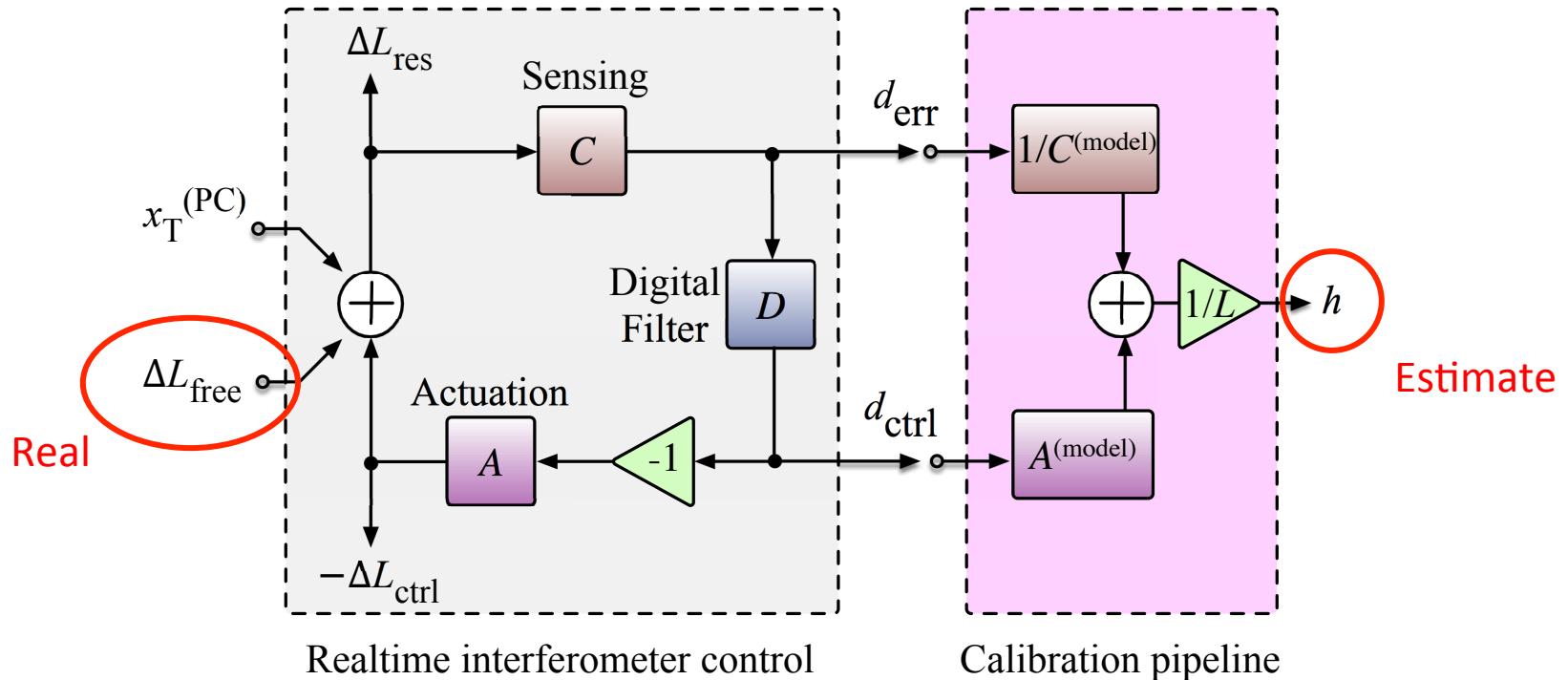
Outline

- What is GW interferometer calibration
- State-of-the-art calibration methods
 - Modeling and measurement
 - Hardware / techniques
 - Uncertainty estimation
 - Computing $h(t)$
- GW interferometer calibration challenges

Gravitational wave interferometers



Gravitational wave measurement



Abbott, B P, et al. PRD **95** 062003

$$d_{\text{err}} = \frac{\Delta L_{\text{free}} C}{1 + CDA}$$

$$d_{\text{ctrl}} = Dd_{\text{err}}$$

$$\Rightarrow \Delta L_{\text{free}} = C^{-1}d_{\text{err}} + Ad_{\text{ctrl}}$$

$$h(t) = \frac{\Delta L_{\text{free}}(t)}{L}$$

Measurement of loop components

$$\Delta L_{\text{free}} = C^{-1} d_{\text{err}} + A d_{\text{ctrl}}$$

- Measure free parameters in C and A
- Require techniques with small systematic and statistical uncertainty
- Techniques with different length fiducials are desirable

Measurement techniques

- Several methods developed over the last 20+ years
 - Laser wavelength
 - Photon radiation pressure
 - Frequency modulation
 - “Gravitational” calibration

Measurement of free parameters

Interferometer measurement

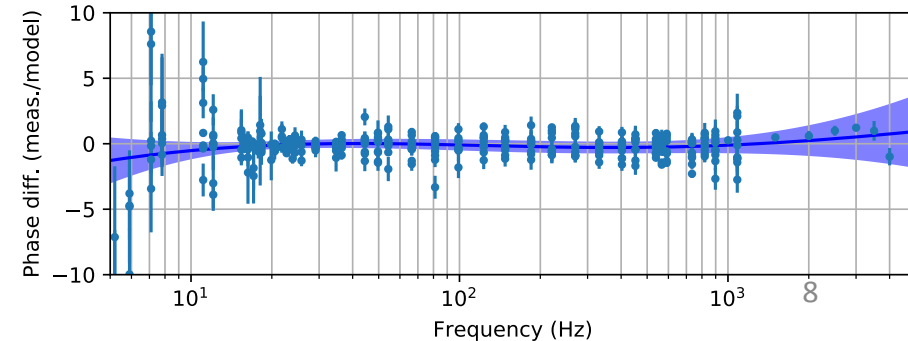
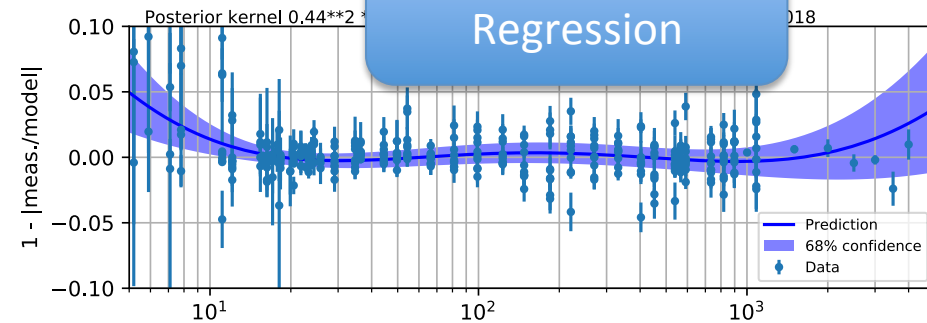
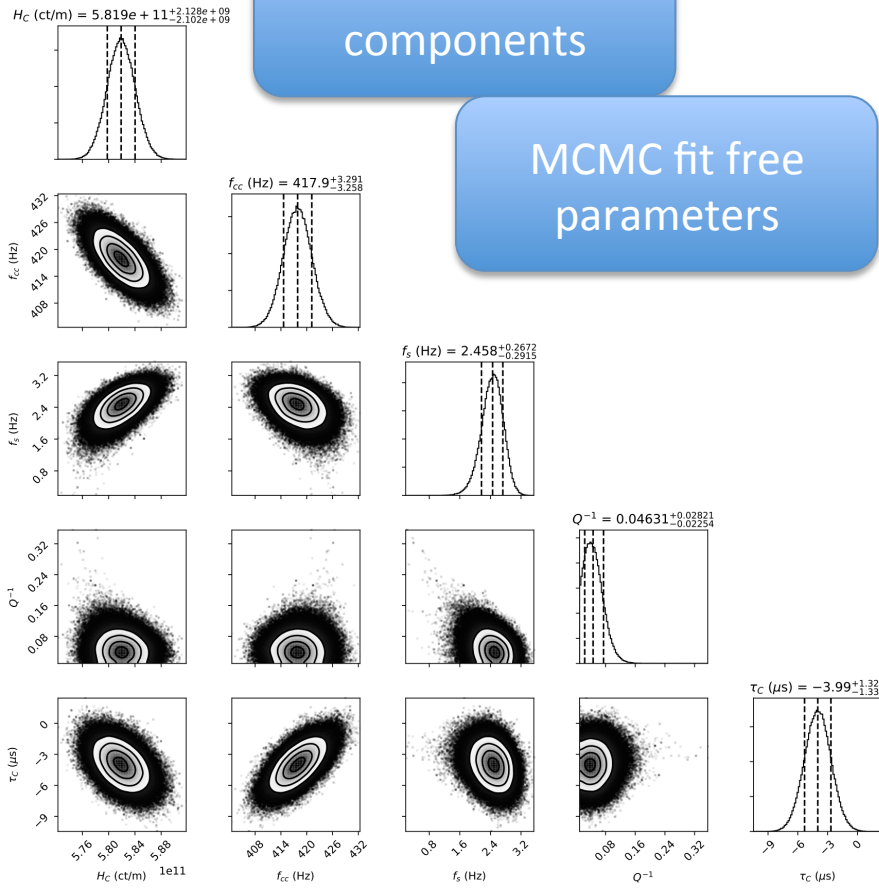
Multiple measurements

Divide known components

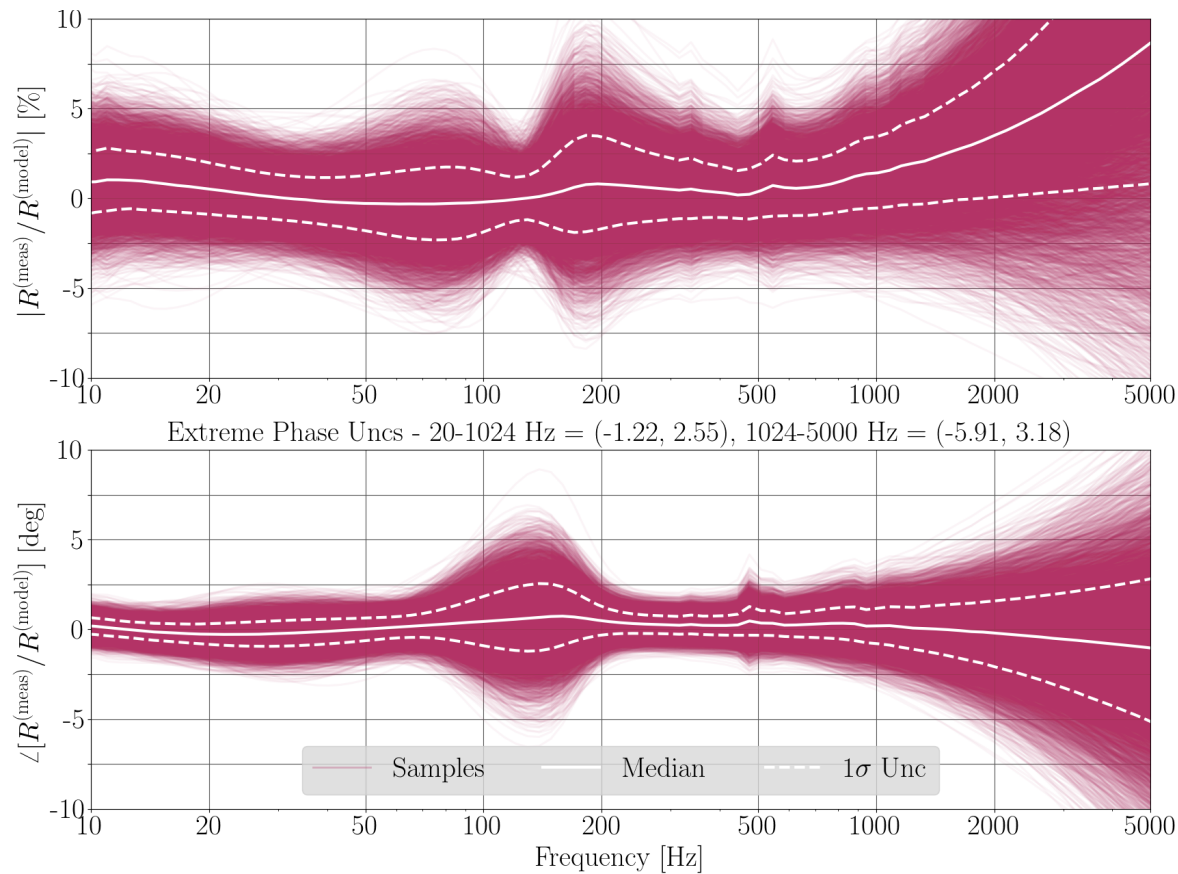
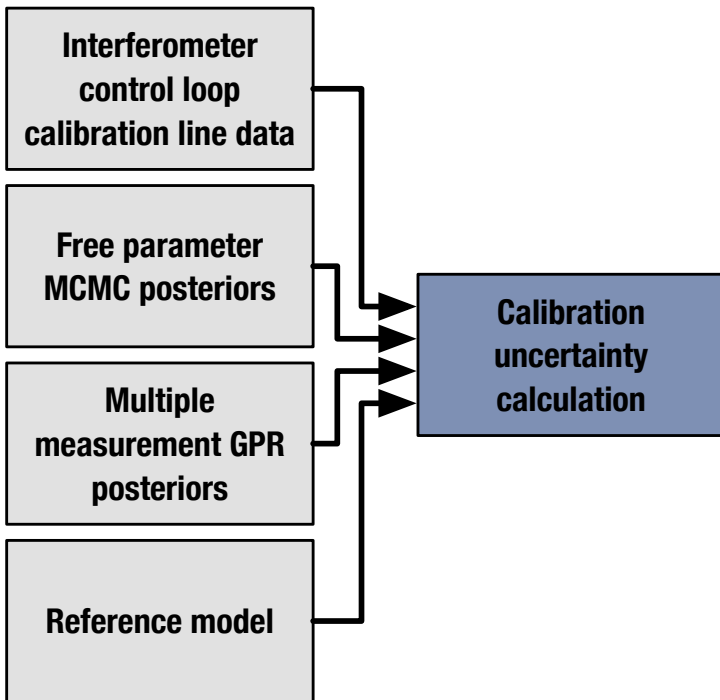
Remove known time-dependence

MCMC fit free parameters

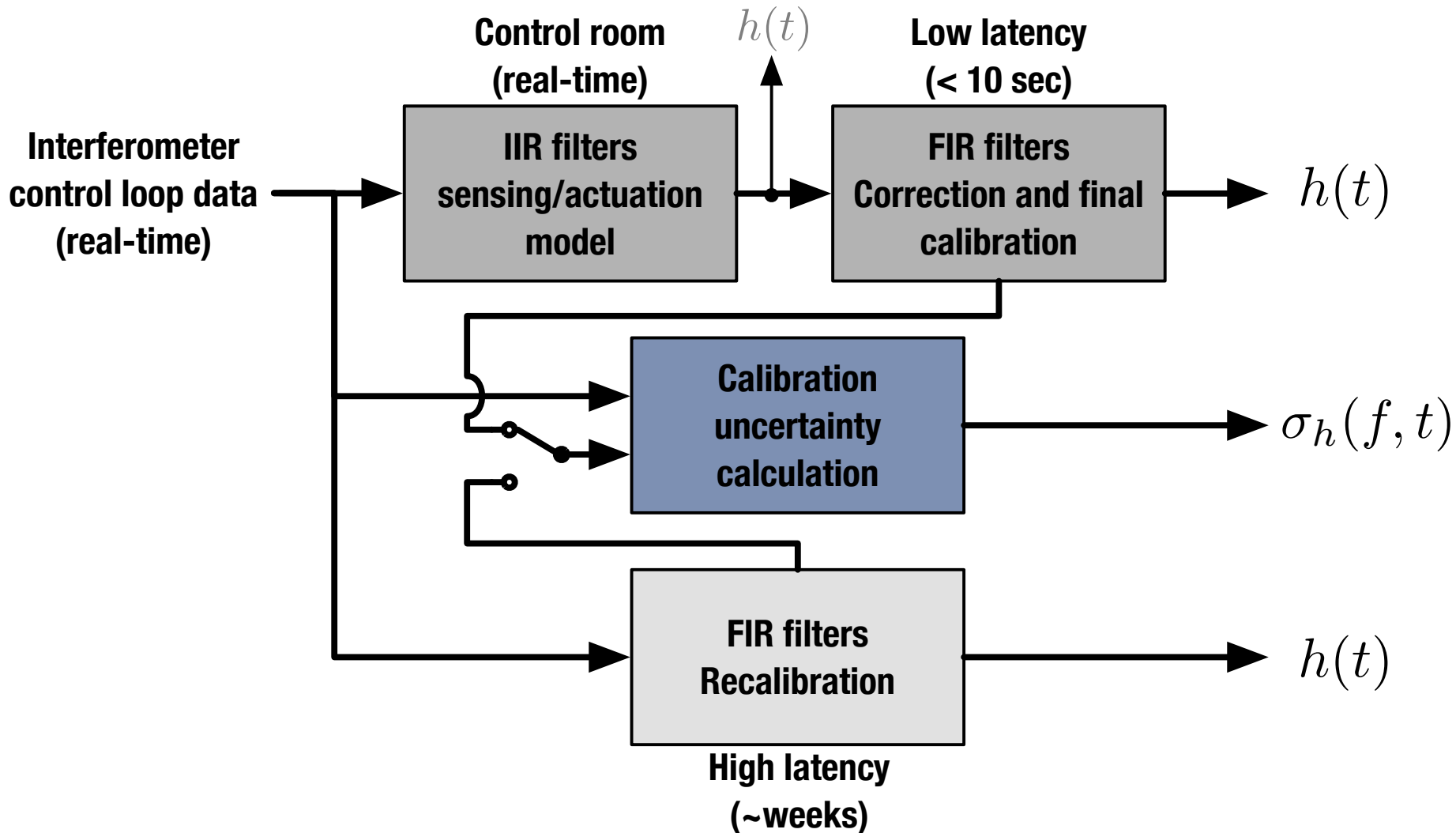
Gaussian Process Regression



Estimating GW measurement uncertainty



Putting everything together

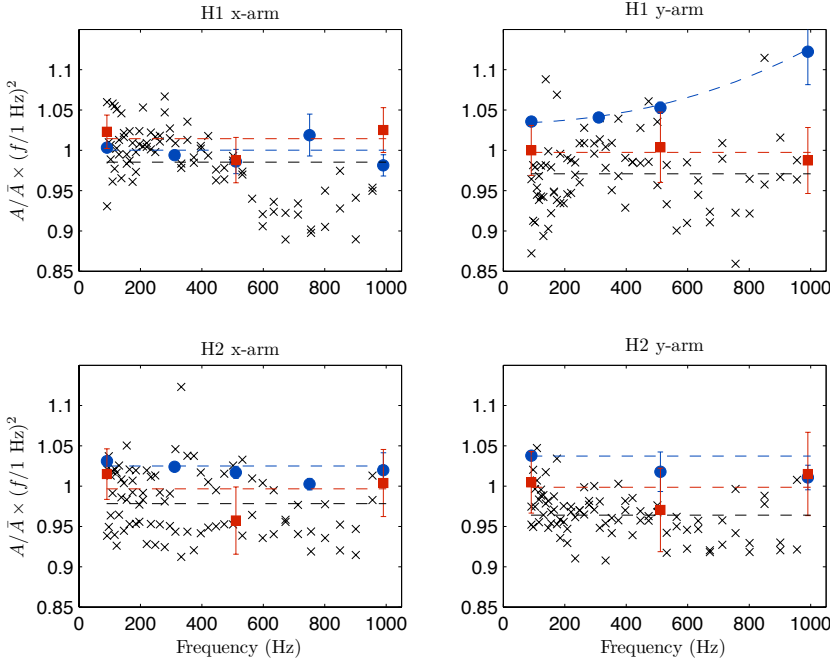


Challenges

- Time- and frequency-dependent systematic and statistical uncertainty must be estimated
- Record keeping and tracking
- Maintain calibration hardware / infrastructure
- Track a changing interferometer (purposefully or not)

Comparing fundamentally different techniques

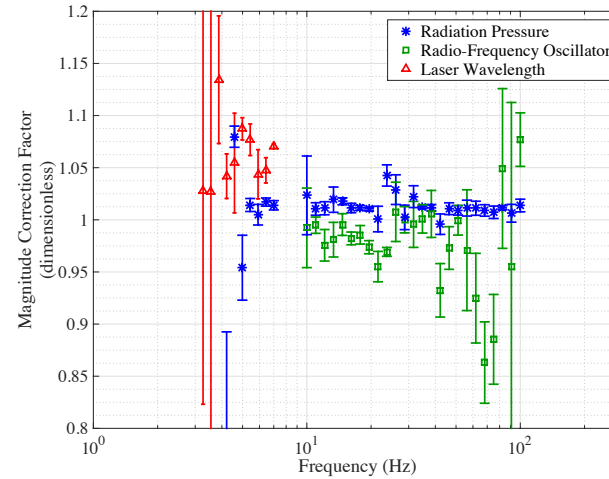
iLIGO comparison (S5, 2007)



Laser wavelength
 Frequency modulation
 Radiation pressure

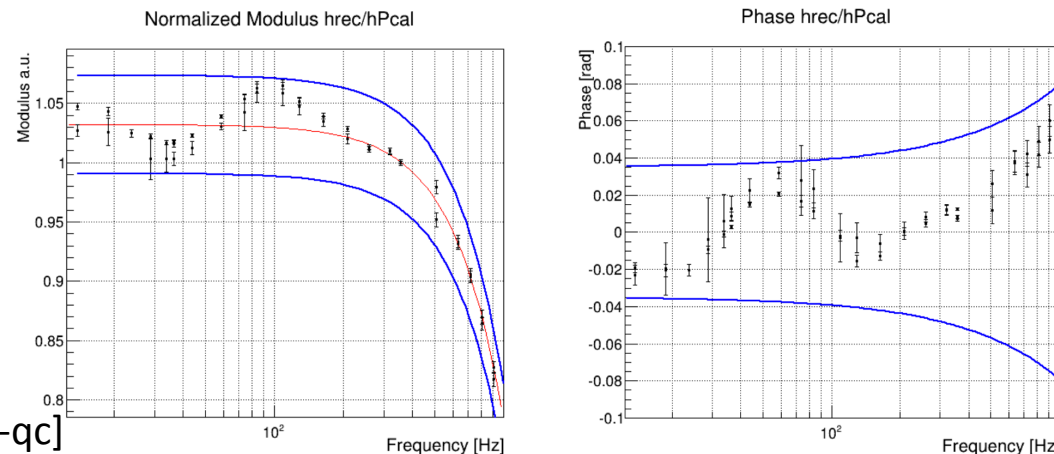
Goetz E, et al. CQG **27** 084024

aLIGO comparison (O1, 2015)



Abbott, B P, et al.
 PRD **95** 062003

Virgo comparison (O2, 2017)



Acernese F, et al.
 arXiv:1807.03275 [gr-qc]

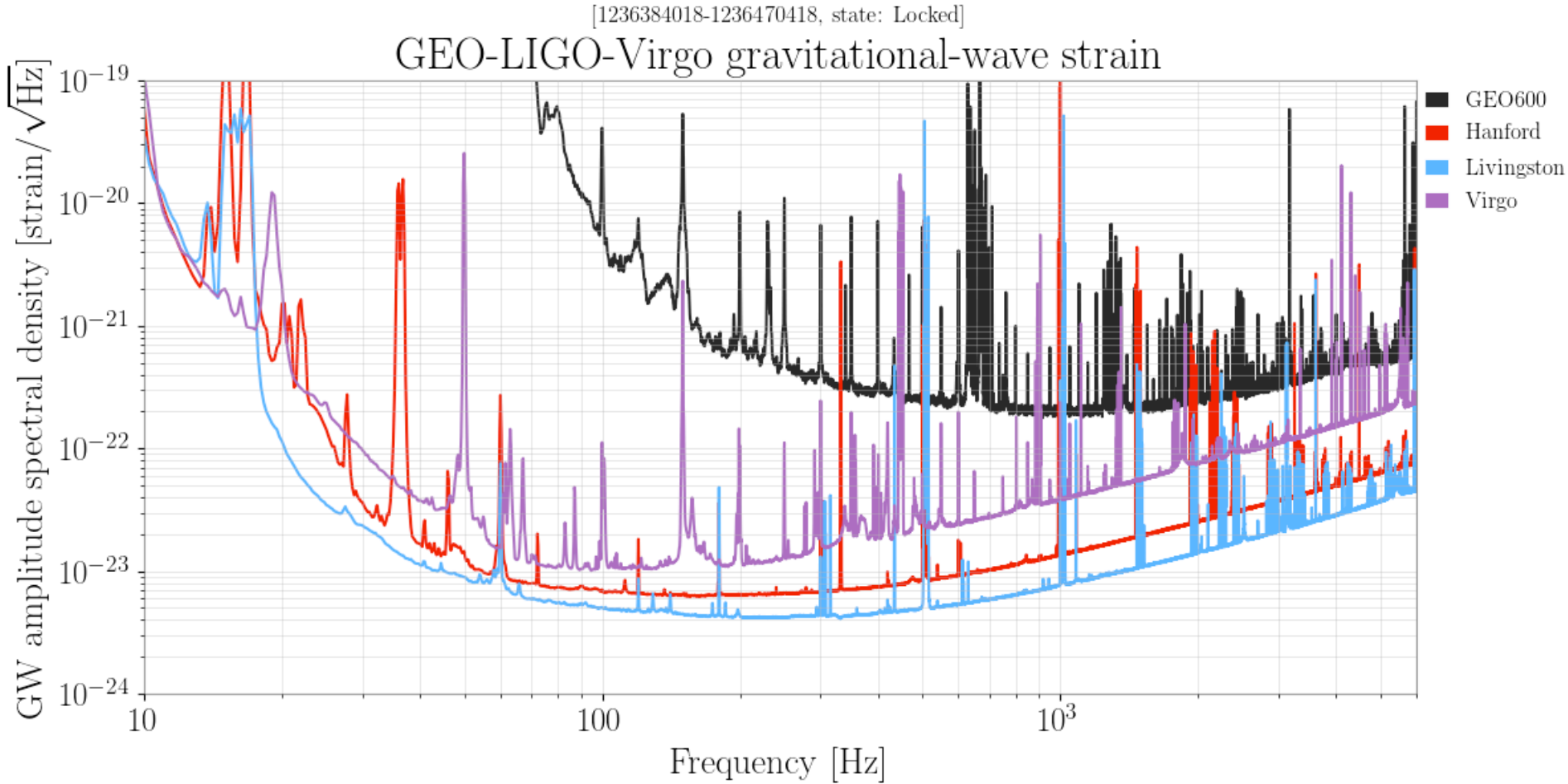
Conclusion and outlook

- More GW signals – the future is bright!
- Detector calibration will play increasingly important role
- Continuing characterization, maintenance, development
- Era of precision astrophysical/cosmological measurement with GWs is now

A white LEGO Technic structure is shown on a wooden surface. The structure consists of a long axle with a pulley at one end and a motor-like component at the other. The pulley is made of a small square block with a hole in the center, and the motor is a larger square block with a hole in the center. The axle is a long white Technic axle with several holes. The structure is built on a white Technic base plate with a grid of holes. The background is a wooden surface with a light brown color and a vertical grain pattern.

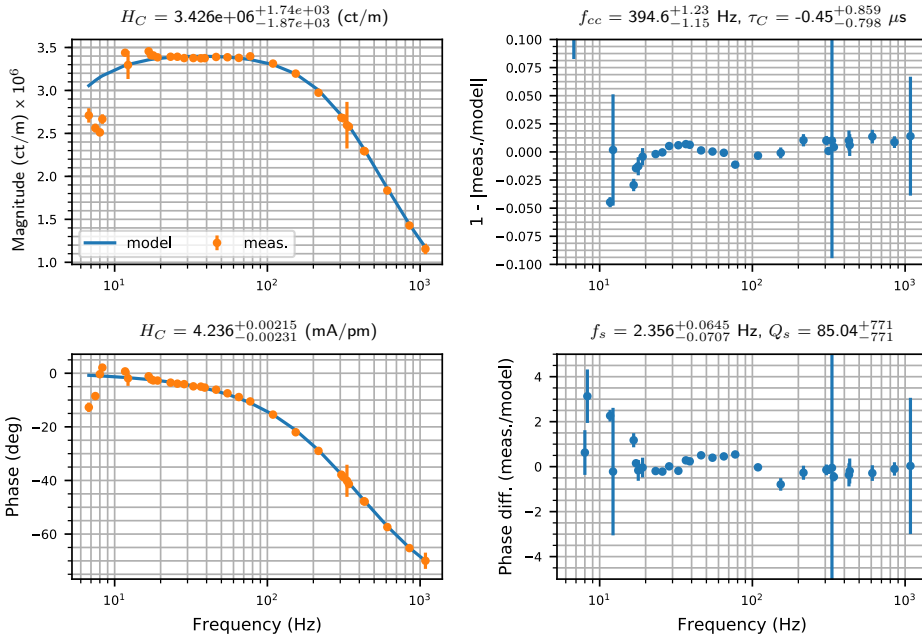
**THANK YOU
QUESTIONS?**

Calibrated interferometer sensitivity

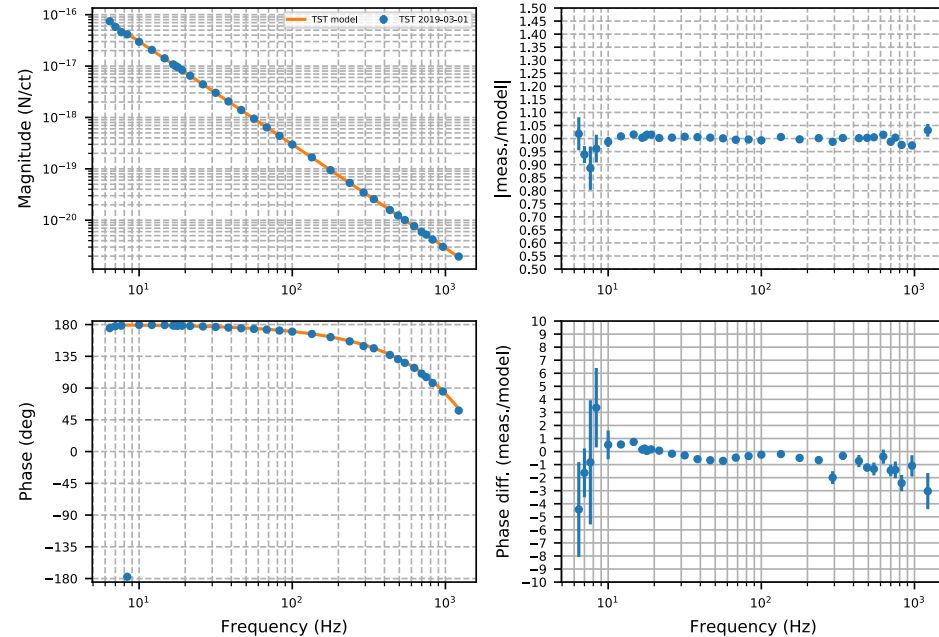


Sensing and actuation measurement examples

H1 sensing function measurement: 2019-03-07



2019-03-01 H1 TST Actuation Function: (PCAL/iStage SUS EXC) vs. Model

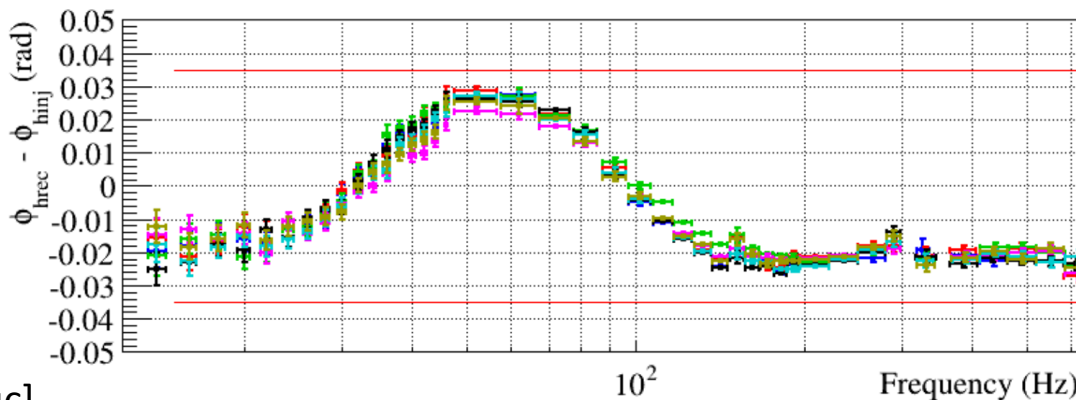
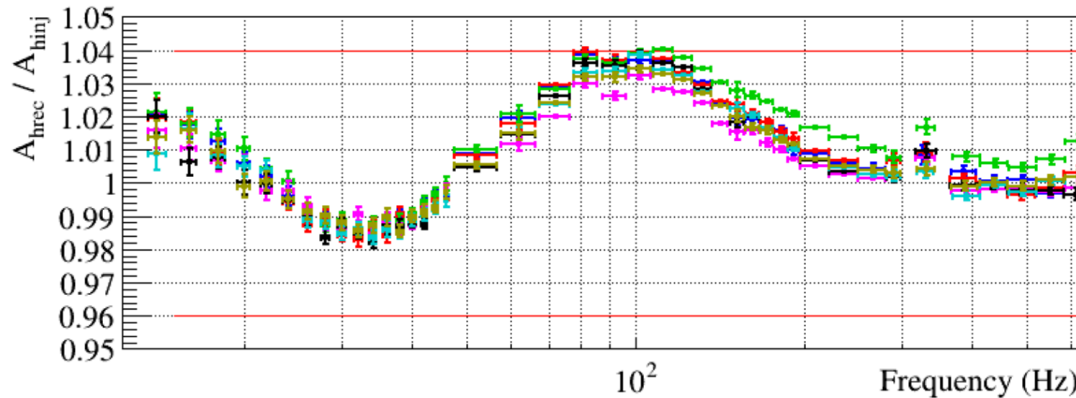


Measurement technique development

- Free-swinging Michelson (laser wavelength)
 - GEO600, iLIGO, Virgo, (aLIGO cross-check)
- Photon radiation pressure (laser power)
 - aLIGO, KAGRA, (GEO600, iLIGO, Virgo cross-checks)
- Frequency modulation (Fabry-Perot frequency-to-length transfer function)
 - (GEO600, iLIGO, aLIGO cross checks)
- Gravitational calibration (rotating masses, G)
 - (Under development: Virgo, aLIGO, KAGRA)

Virgo calibration

$h(t)$ version	Amplitude uncertainty (%)	Phase uncertainty (rad)	Timing bias
Online	+14/ - 8	$100 \times 10^{-3} + 2\pi f(20 \times 10^{-6})$	116 μ s
V1O2Repro1A	± 8	$50 \times 10^{-3} + 2\pi f(20 \times 10^{-6})$	0
V1O2Repro2A	± 5.1	$40 \times 10^{-3} + 2\pi f(20 \times 10^{-6})$	0



Many publications!

- Goetz, et al. CQG **26** 24
- Goetz, et al. CQG **27** 8
- Goetz, et al. CQG **27** 21
- Karki, et al. RSI **87** 114503
- Abadie et al NIM A **624** 1
- Tuyenbayev, et al CQG **34** 1
- Abbott, et al PRD **95** 062003
- Viets, et al. CQG **35** 9
- Leong, et al. CQG **29** 6
- Accadia, et al. JoP Conf. **228** 1
- Accadia, et al. CQG **28** 2
- Inoue, et al. PRD **98** 022005
- Adhikari, et al. CQG **20** 17
- Hewitson, et al. CQG **20** 17
- Siemens, et al. CQG **21** 20
- Landry, et al. CQG **22** 18
- Hewitson, et al. RSI **74** 4184
- Hewitson, et al. RSI **75** 4702
- Hewitson, et al. CQG **21** 20
- Mossavi, et al. PL A **353** 1
- Hewitson, et al. CQG **22** 20
- Clubley, et al. PL A **283** 1-2
- Acernese, et al. CQG **35** 20
- Yamamoto, et al. JPS Conf. **1** 013119

And more!