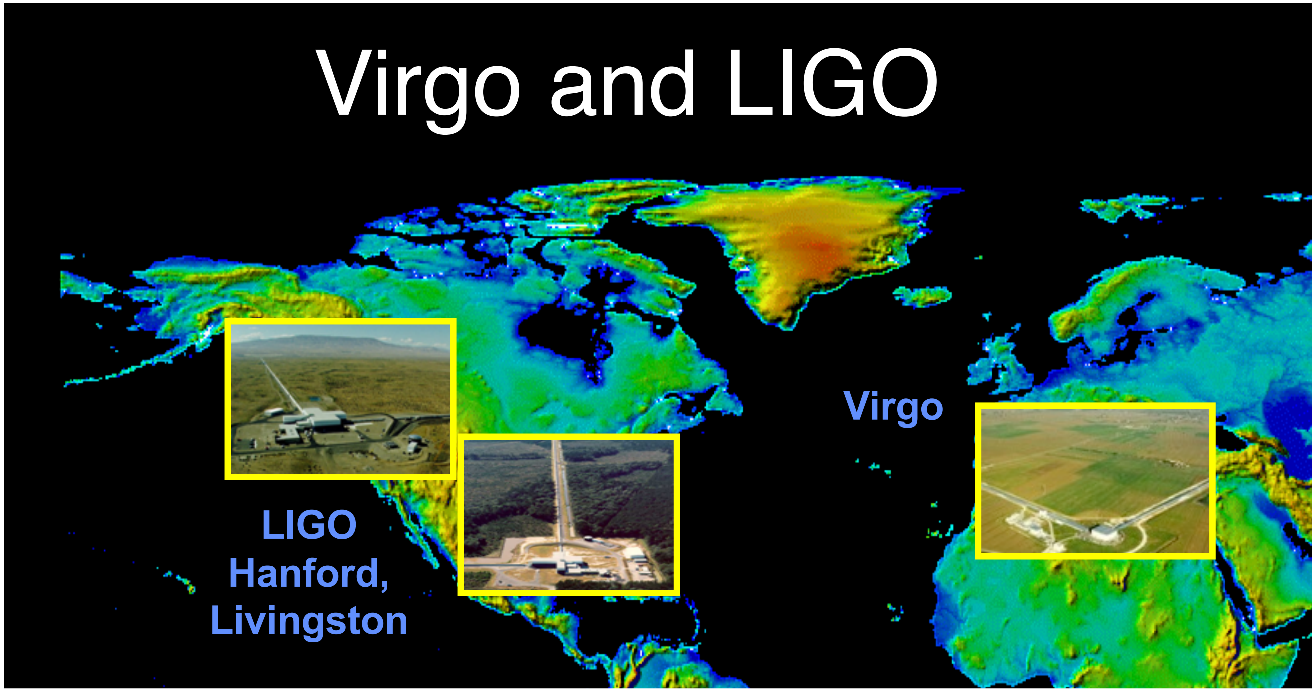
Status, Results, and Future Plans for LIGO and Virgo

YKIS2018a, Kyoto
19 February 2018

David Shoemaker
For the LIGO and Virgo Scientific Collaborations

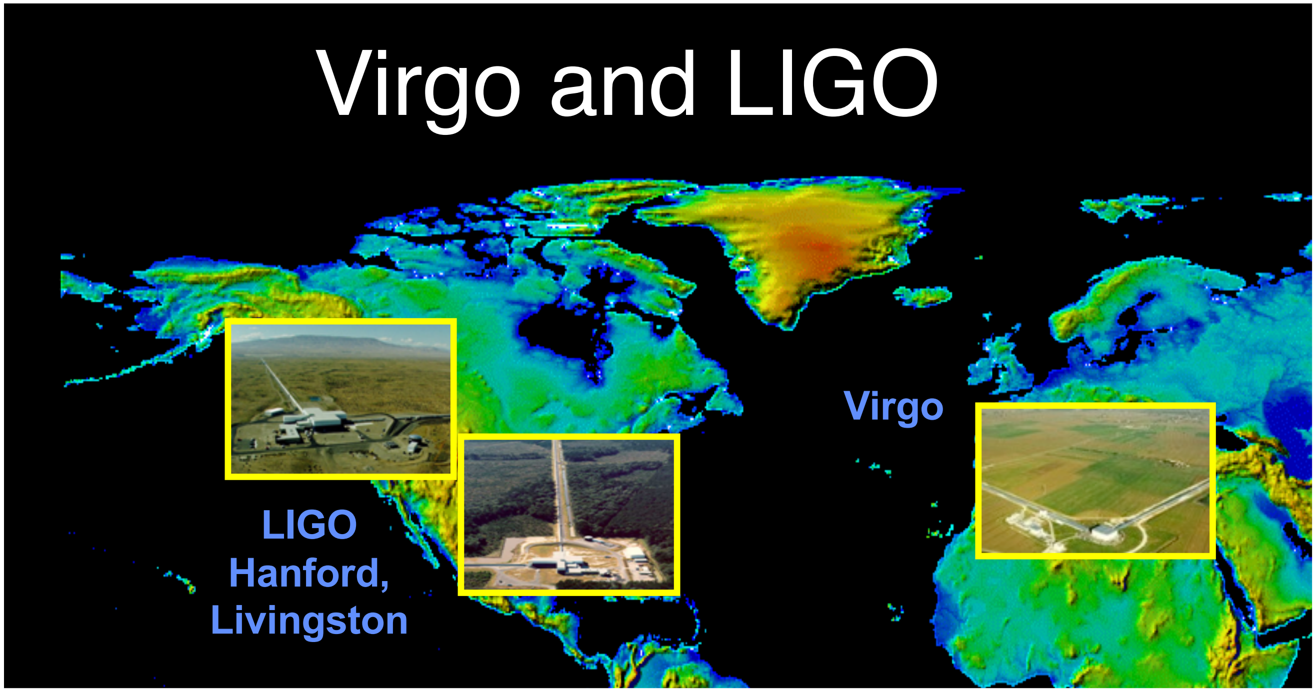
Credits
Measurement results: LIGO/Virgo Collaborations,
PRL 116, 061102 (2016); Phys. Rev. Lett. 119, 161101 (2017);
Phys. Rev. Lett. 119, 141101 (2017); hys. Rev. Lett. 118, 221101 (2017);
Phys. Rev. Lett. 116, 241103 (2016)
Simulations: SXS Collaboration; LIGO Laboratory
Localization: S. Fairhurst arXiv:1205.6611v1
Slides from (among others) L. Nuttal, P. Fritschel, L. Cadonati
Photographs: LIGO Laboratory; MIT; Caltech; Virgo

LIGO-G1800255-v1



Virgo and LIGO built new observatories in the 90’s

LIGO thanks the NSF for its vision and support!



Virgo and LIGO built new observatories in the 90’s

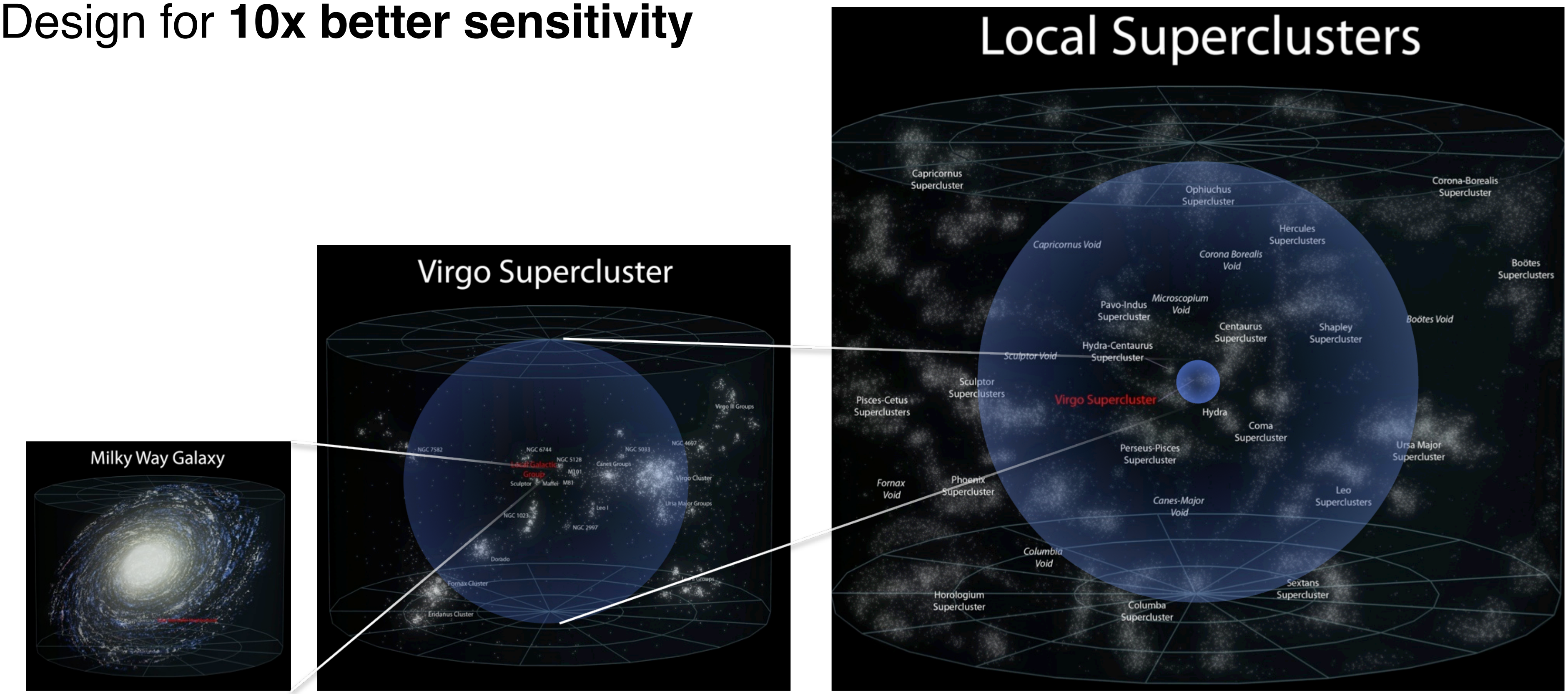
…and Observed with the initial detectors
2005-2011,
and saw **no signals**

(with some interesting non-detections)

Advanced Detectors: a *qualitative* difference

- Foreseen in original 1989 proposal
- While observing with initial detectors, parallel R&D led to better concepts
- Design for **10x better sensitivity**

- We measure amplitude, so signal falls as $1/r$
- **1000x more candidates**



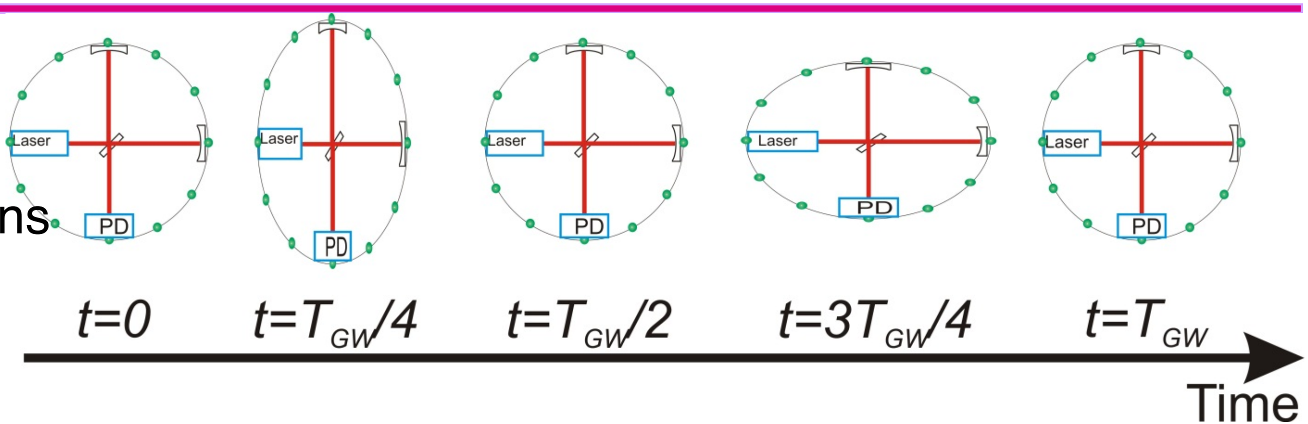
M. Evans

Initial Reach

Advanced Reach

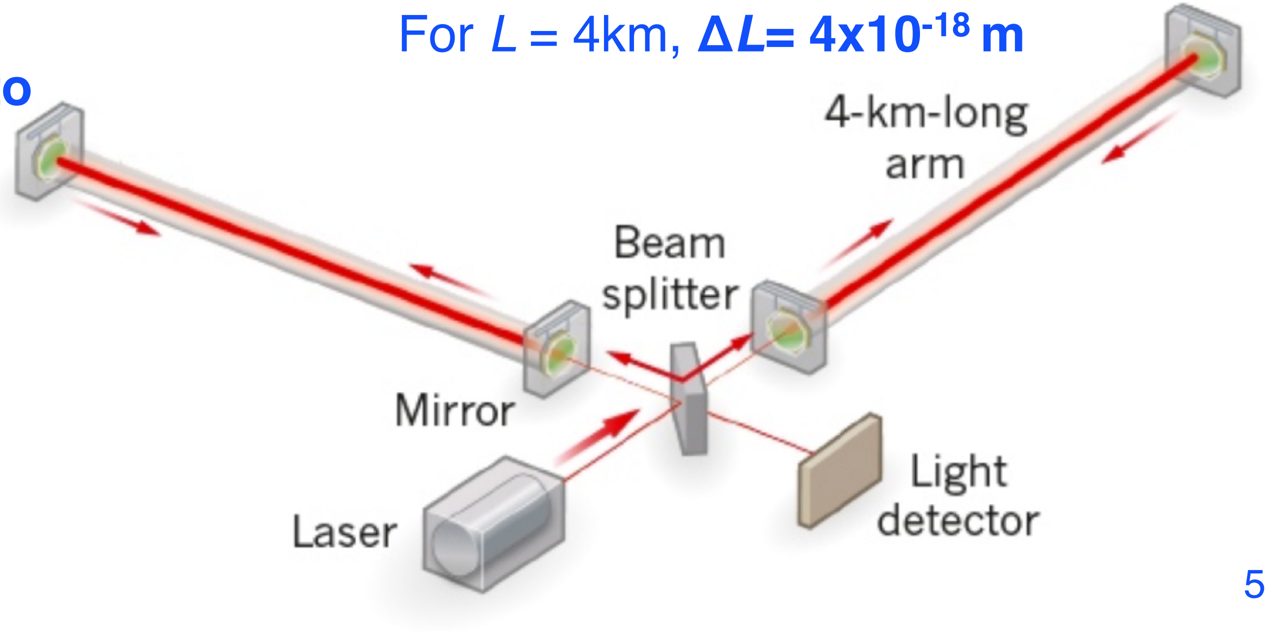
What is our measurement technique?

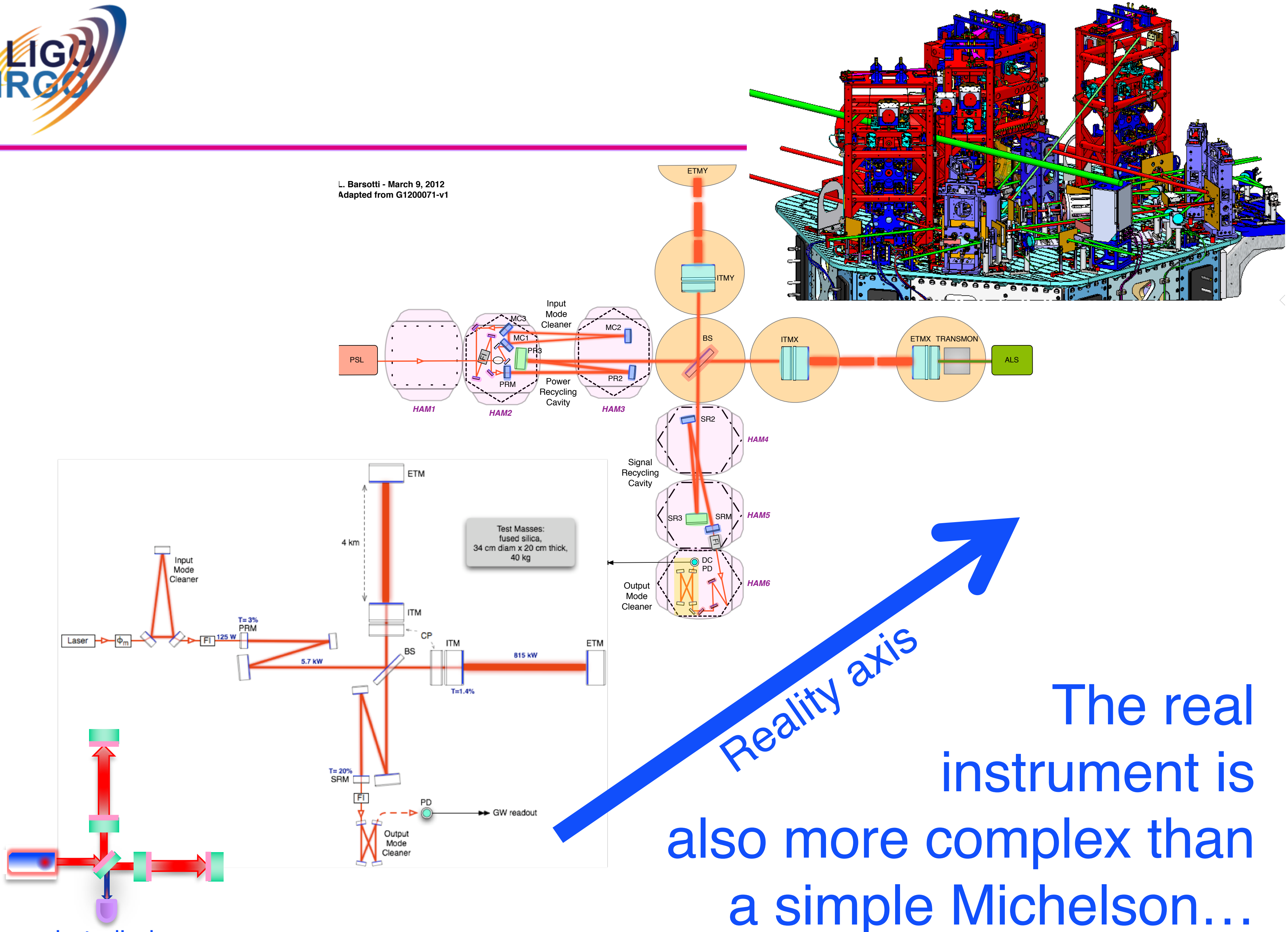
- Enhanced **Michelson interferometers**
  - » LIGO, Virgo use variations
- GWs modulate the distance between the end test mass and the beam splitter
- The interferometer acts as a transducer, turning GWs into photocurrent proportional to the strain amplitude
- **Arms are short compared to our GW wavelengths, so longer arms make bigger signals**
  → multi-km installations
- Arm length limited by taxpayer noise….



$$h \approx \frac{\Delta L}{L}$$

Magnitude of h at Earth:
Detectable signals h ~ 10^{-21}
(1 hair / Alpha Centauri)
For L = 1 m, $\Delta L = 10^{-21}$ m
For L = 4km, $\boldsymbol{\Delta L = 4\times10^{-18}}$ **m**



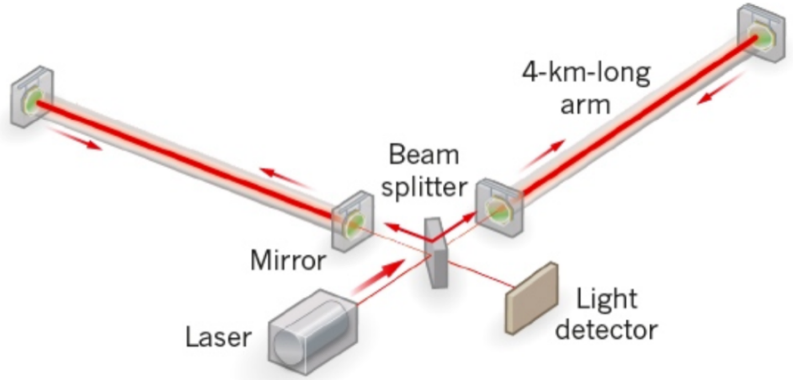


L. Barsotti - March 9, 2012
Adapted from G1200071-v1

Test Masses:
fused silica,
34 cm diam x 20 cm thick,
40 kg

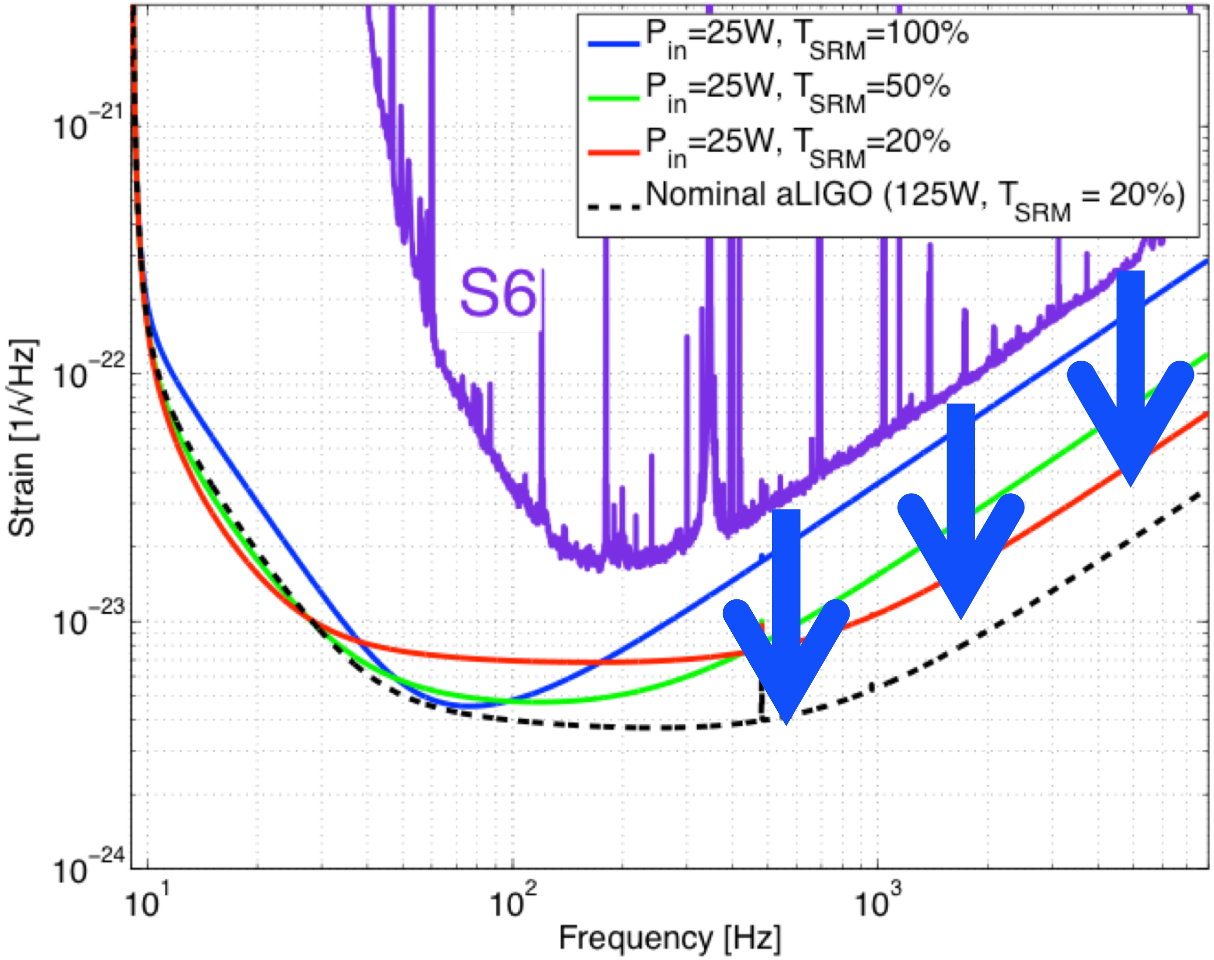
The real instrument is also more complex than a simple Michelson…

Measuring ΔL= 4x10^{-18} m
Readout



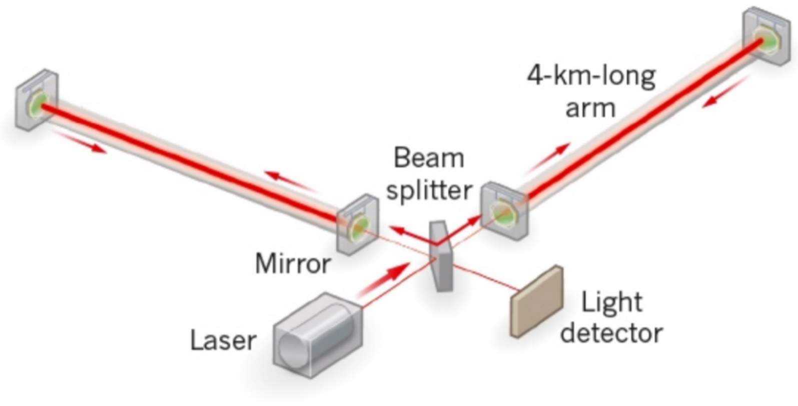
- **Shot noise** – ability to resolve a fringe shift due to a GW (counting statistics)
- *Zum gegenwärtigen Stand des Strahlungsproblems*, A. Einstein, 1909
- Fringe Resolution at high frequencies improves as as (laser power)$^{1/2}$

$$h_{\mathrm{sn}}(f) = \frac{1}{L}\sqrt{\frac{\hbar c \lambda}{2\pi P}}$$



Measuring Δ*L*= 4x10^{-18} m

Readout

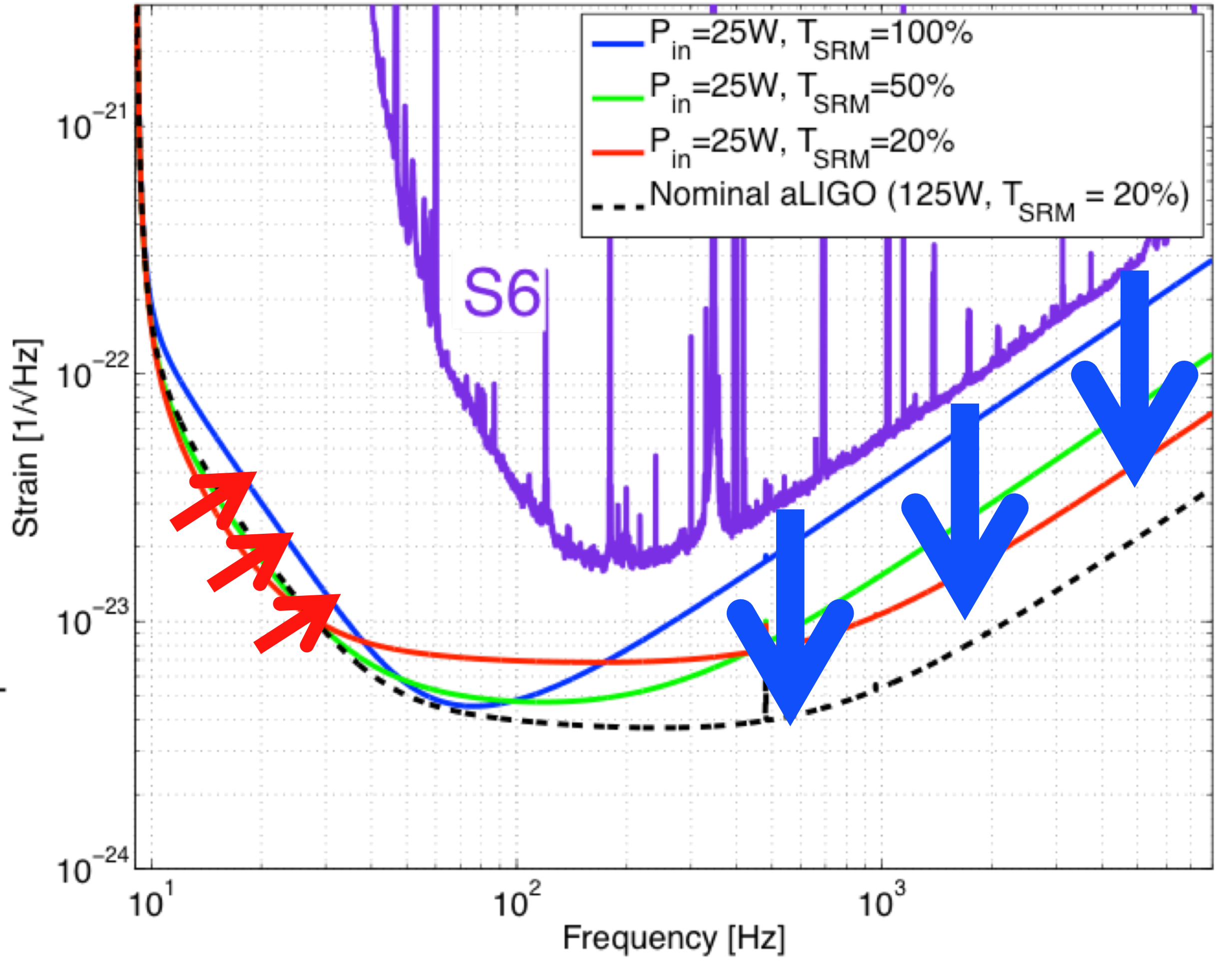


- Shot noise – ability to resolve a fringe shift due to a GW (counting statistics)

$$h_{\rm sn}(f) = \frac{1}{L}\sqrt{\frac{\hbar c \lambda}{2\pi P}}$$

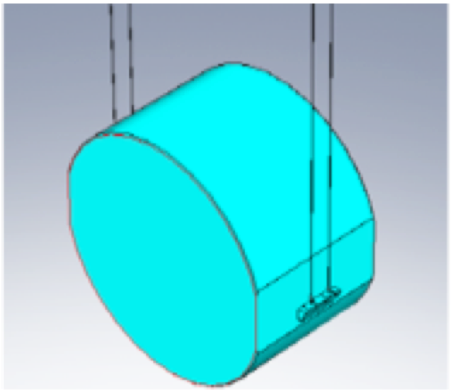
- **Radiation Pressure noise** – Point of diminishing returns when buffeting of test mass by photons increases low-frequency noise – use heavy test masses!

$$h_{\rm rp}(f) = \frac{1}{m f^2 L}\sqrt{\frac{\hbar P}{2\pi^3 c \lambda}}$$



Measuring $\Delta L = 4\times10^{-18}$ m

Internal motion

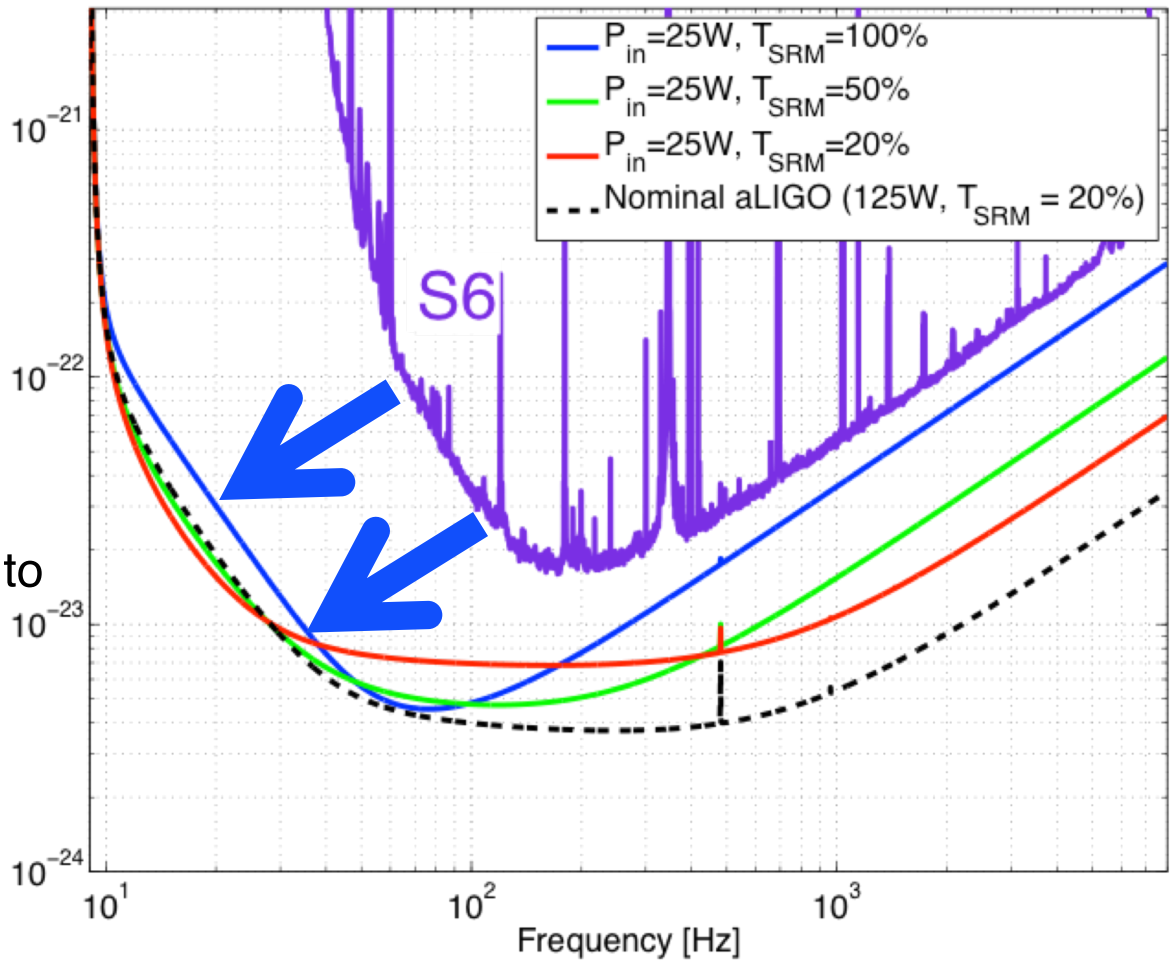


- **Thermal noise** – kT of energy per mechanical mode
- *Über die von der molekularkinetischen Theorie der Wärmegeforderte Bewegung von in ruhenden Flüssigkeiten suspendierten Teilchen,* A. Einstein, 1905
- Simple Harmonic Oscillator:

$$x_{rms} = \sqrt{\left\langle (\delta x)^2 \right\rangle} = \sqrt{k_B T / k_{spring}}$$

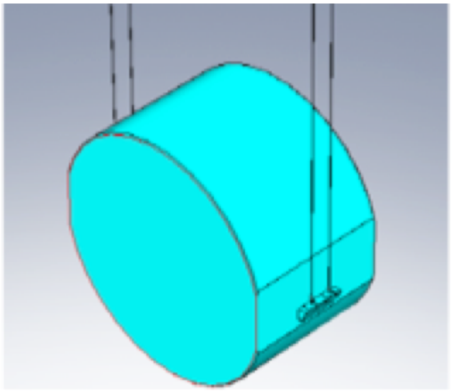
- Distributed in frequency according to real part of impedance $\Re(Z(f))$

$$\tilde{x}(f) = \frac{1}{\pi f}\sqrt{\frac{k_B T}{\Re(Z(f))}}$$



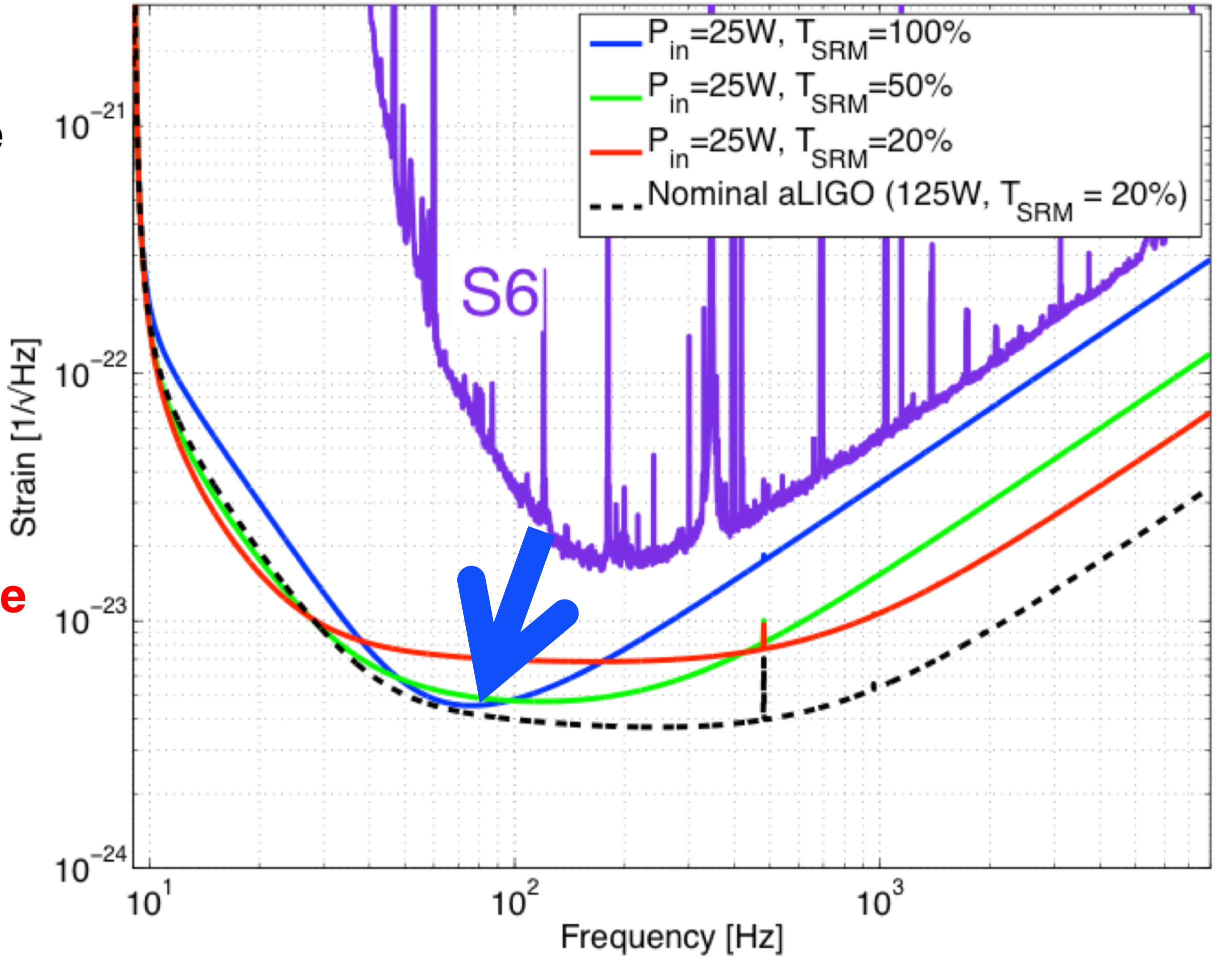
Measuring Δ*L*= 4x10^{-18} m

Internal motion



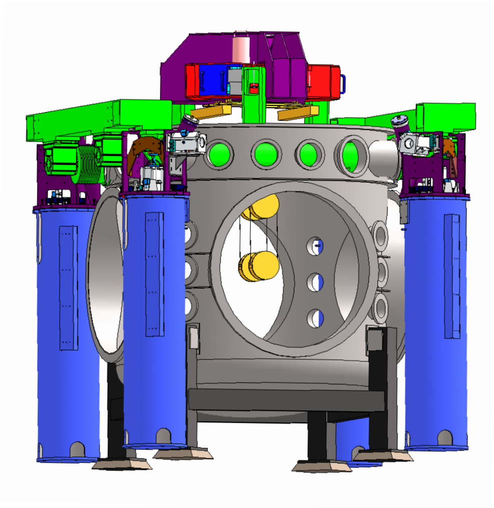
- In Advanced LIGO, the dielectric optical coating has a rather large loss tangent
  - Some 10^{-4}
- And: the coating is the surface that is sensed by the laser

- **This is the dominant limit in the critical 50-200 Hz band**

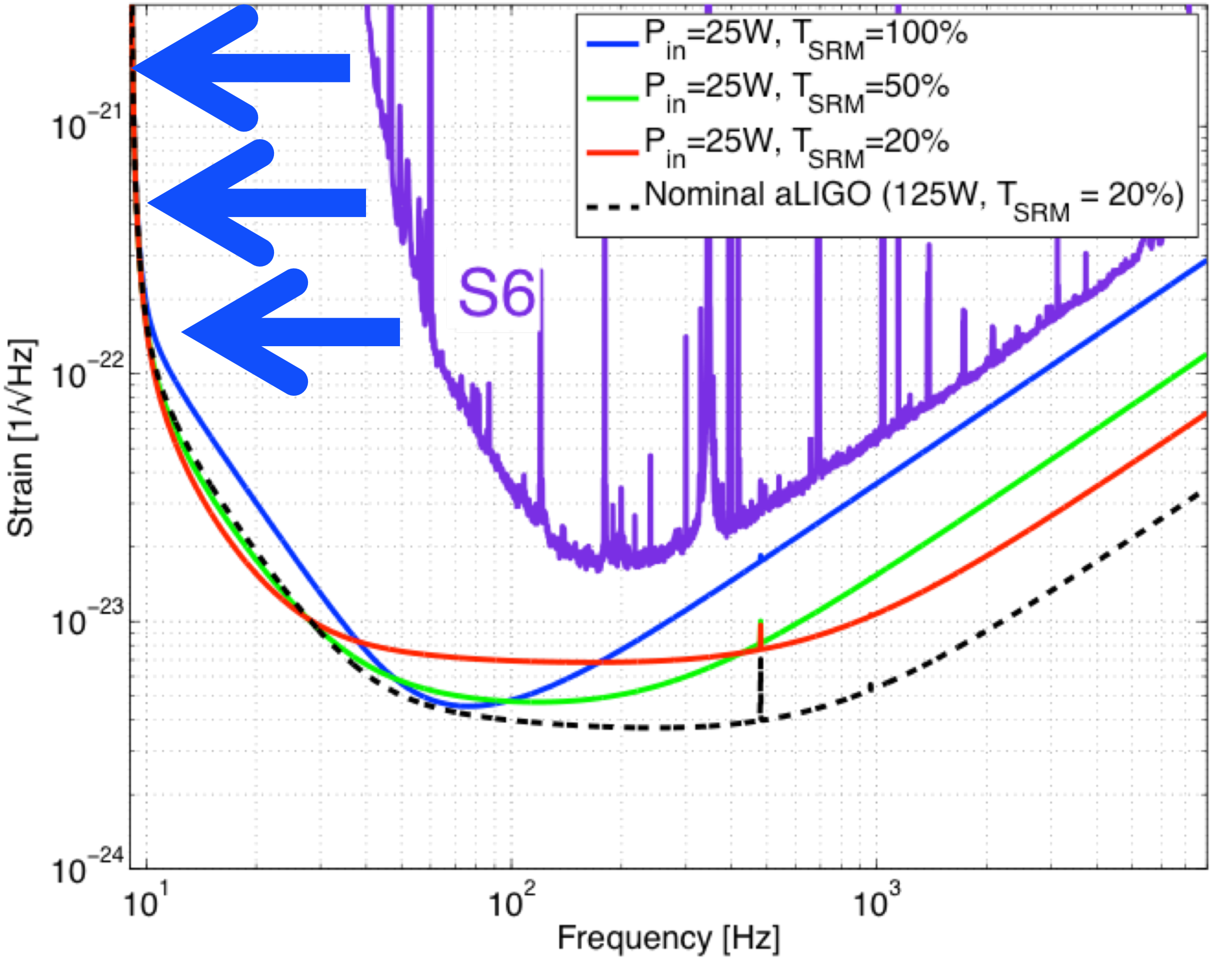


Measuring ΔL= 4x10^{-18} m

Forces on test mass



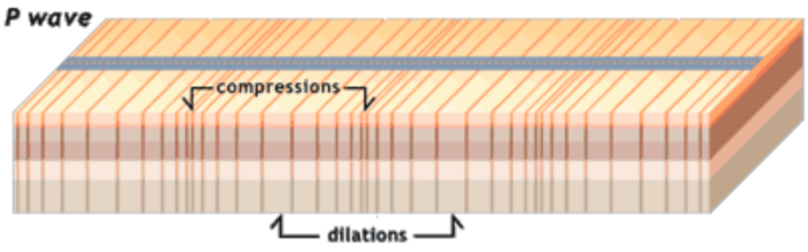
- **Seismic noise** – must prevent masking of GWs, enable practical control systems
- aLIGO uses **active servo-controlled platforms, multiple pendulums**
- 3 layers, each of 6 degrees-of-freedom



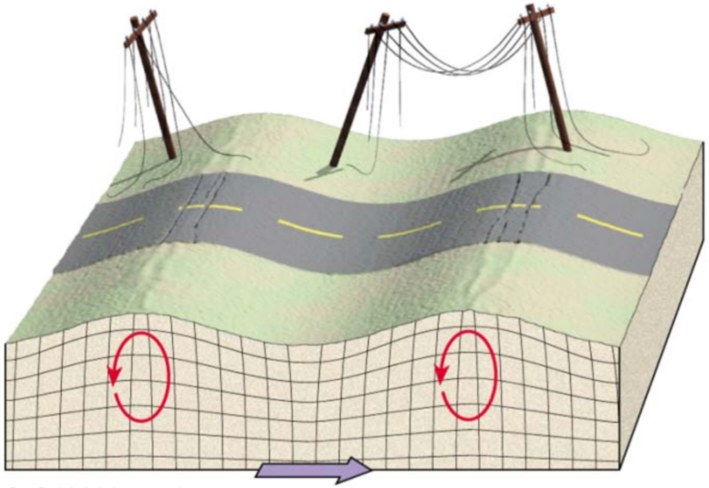
Measuring Δ*L*= 4x10^{-18} m
Forces on test mass

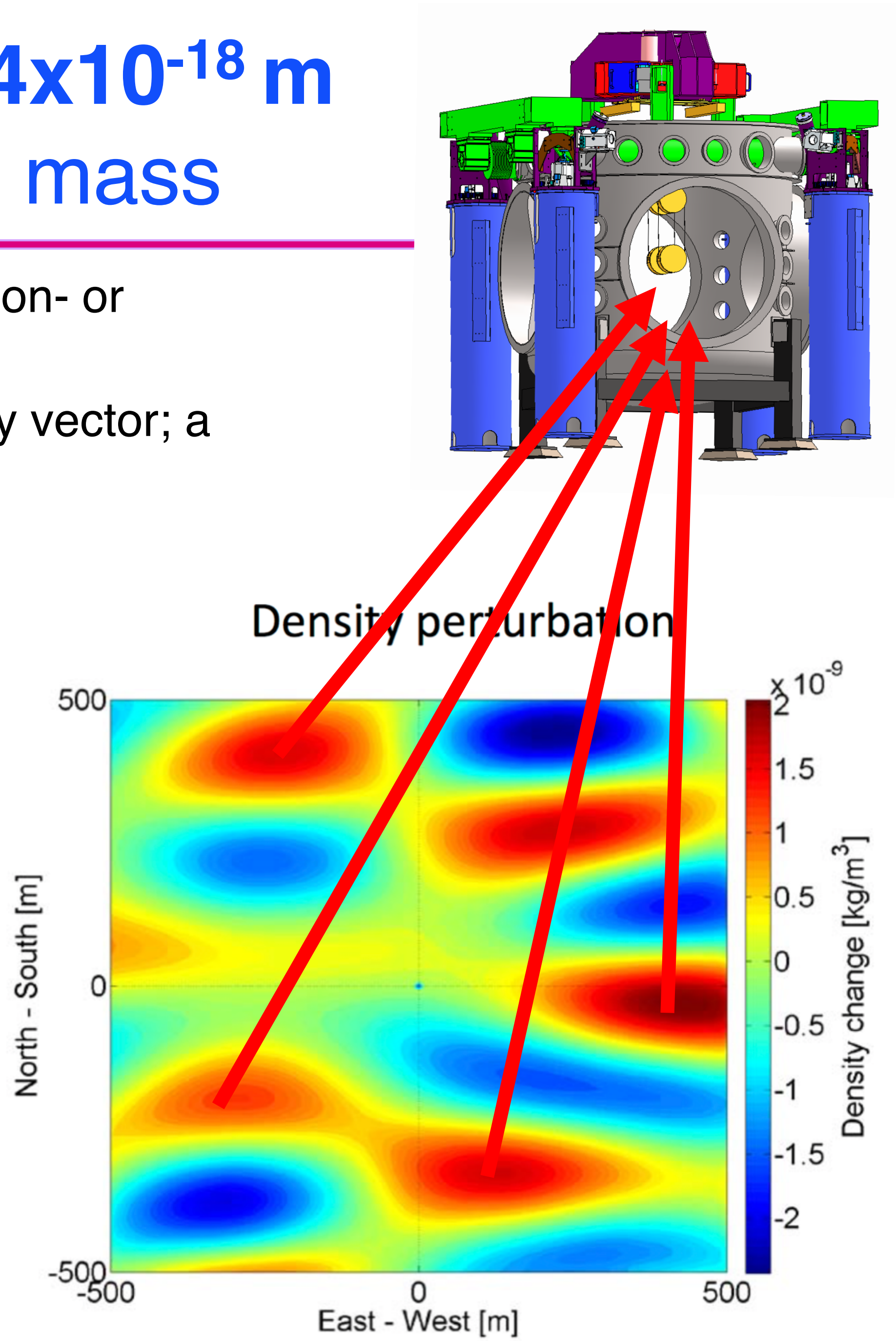
- Ultimate limit on the lowest frequency detectors on- or under-ground:
- Newtownian background – wandering net gravity vector; a limit in the 10-20 Hz band

Body waves



Rayleigh waves



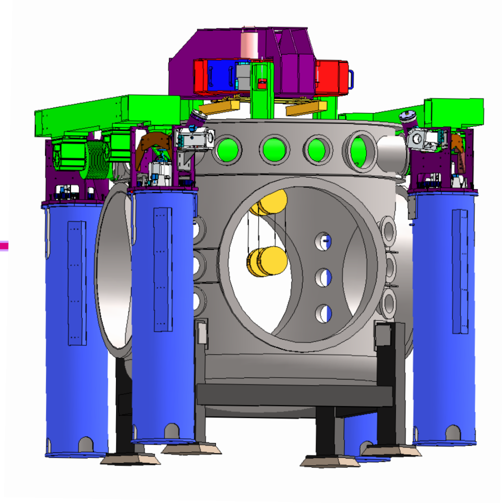


Density perturbations cause gravity perturbations.

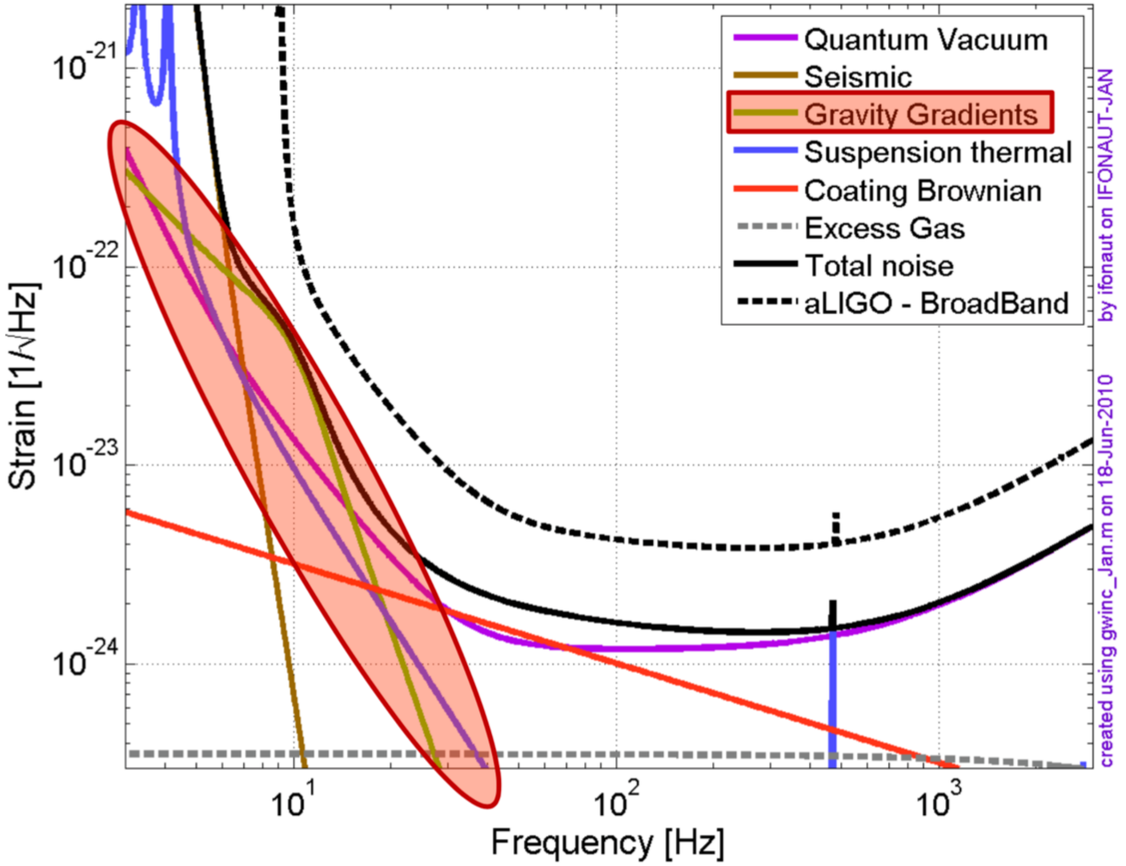
Images: J. Harms

Measuring ΔL= 4x10^{-18} m

Forces on test mass

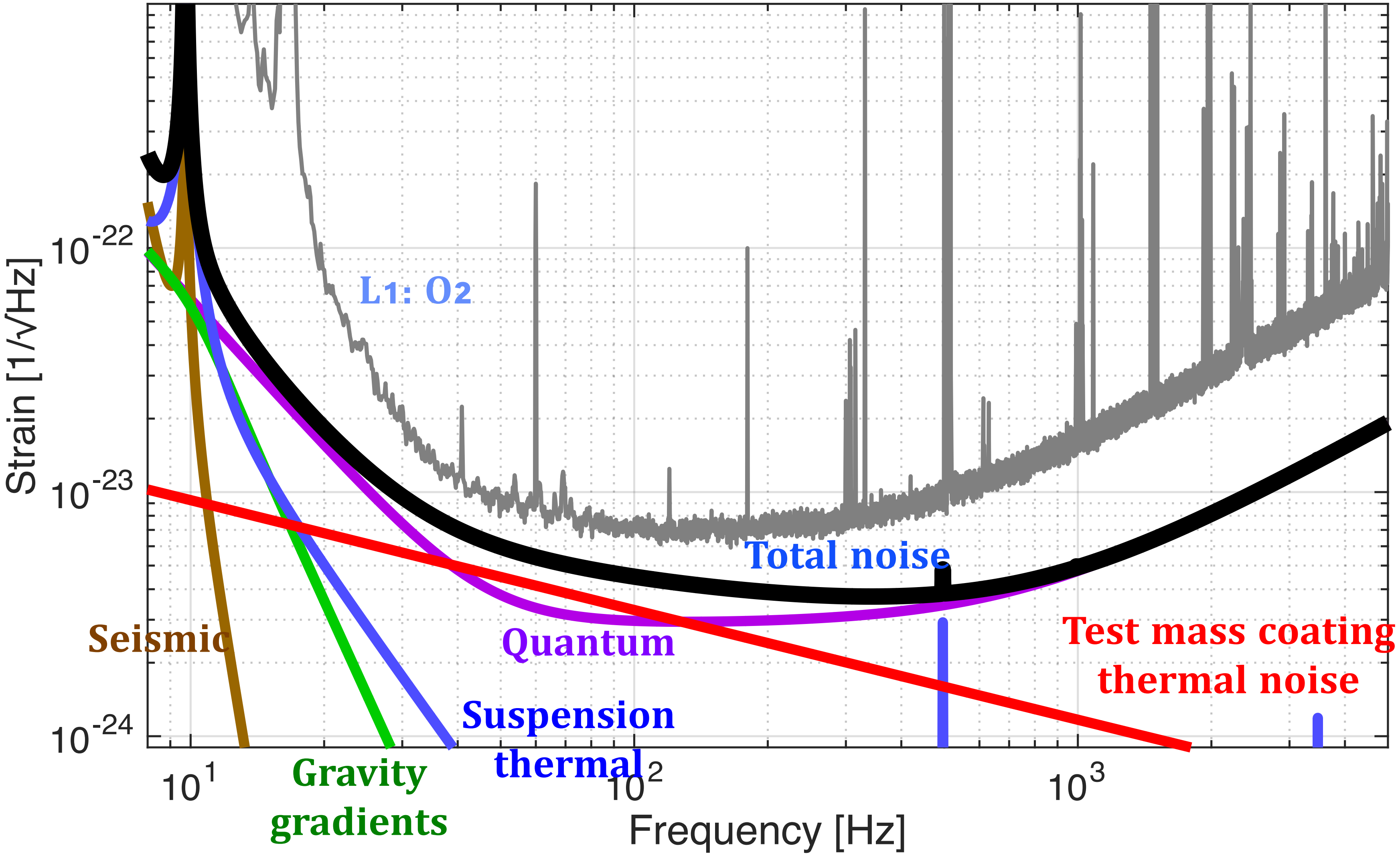


- Advanced LIGO (and Virgo) expect to be limited by this noise source –
  - After all technical noise sources beaten down
  - At low optical power (no radiation pressure noise)
  - In the 10-30 Hz range
- **We would *love* to be limited only by this noise source!**
- Want to go a bit lower? Go underground.
- Want to go much lower? Go to space.

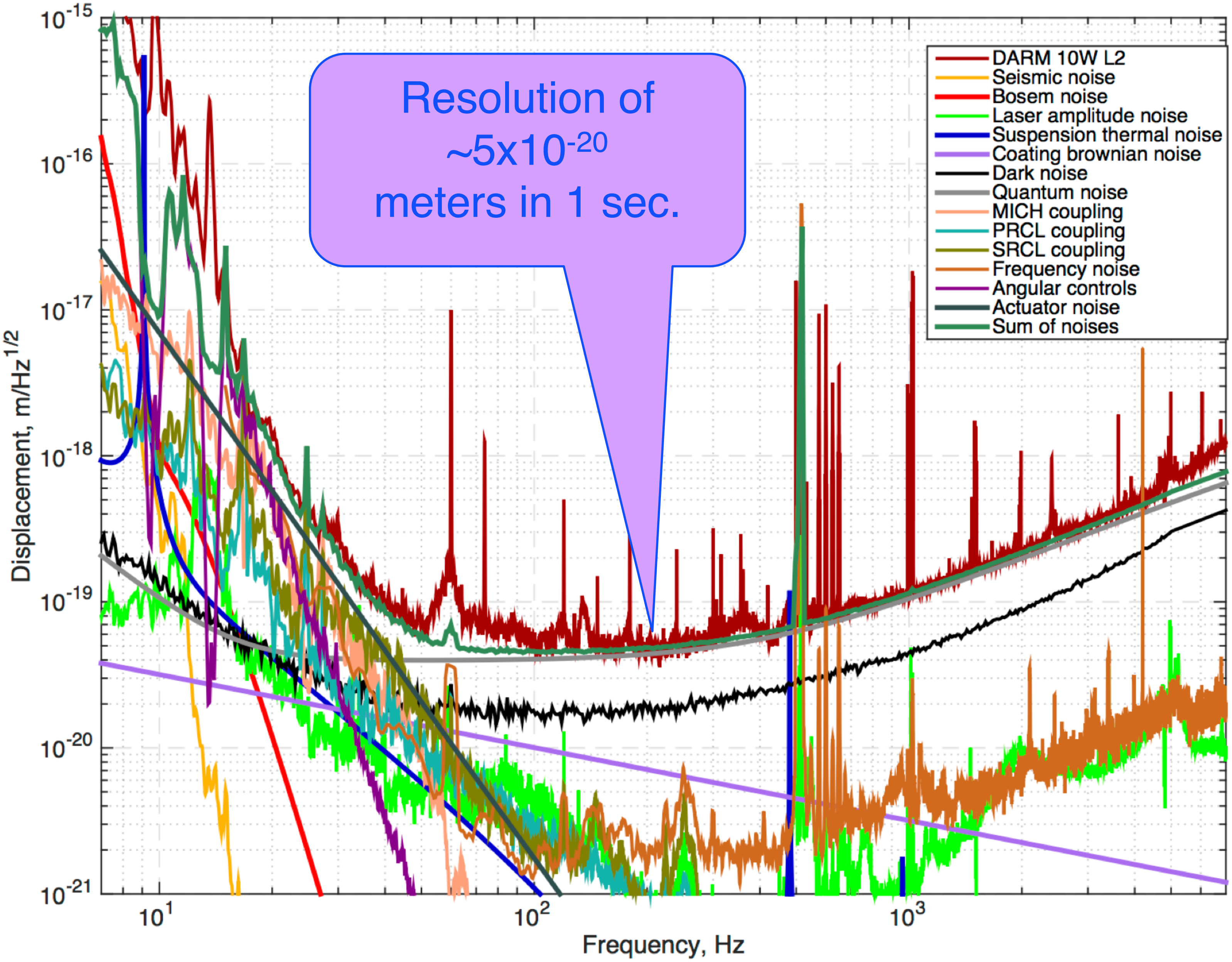


Images: J. Harms

Adv LIGO Target Design
Sensitivity, basic noise sources

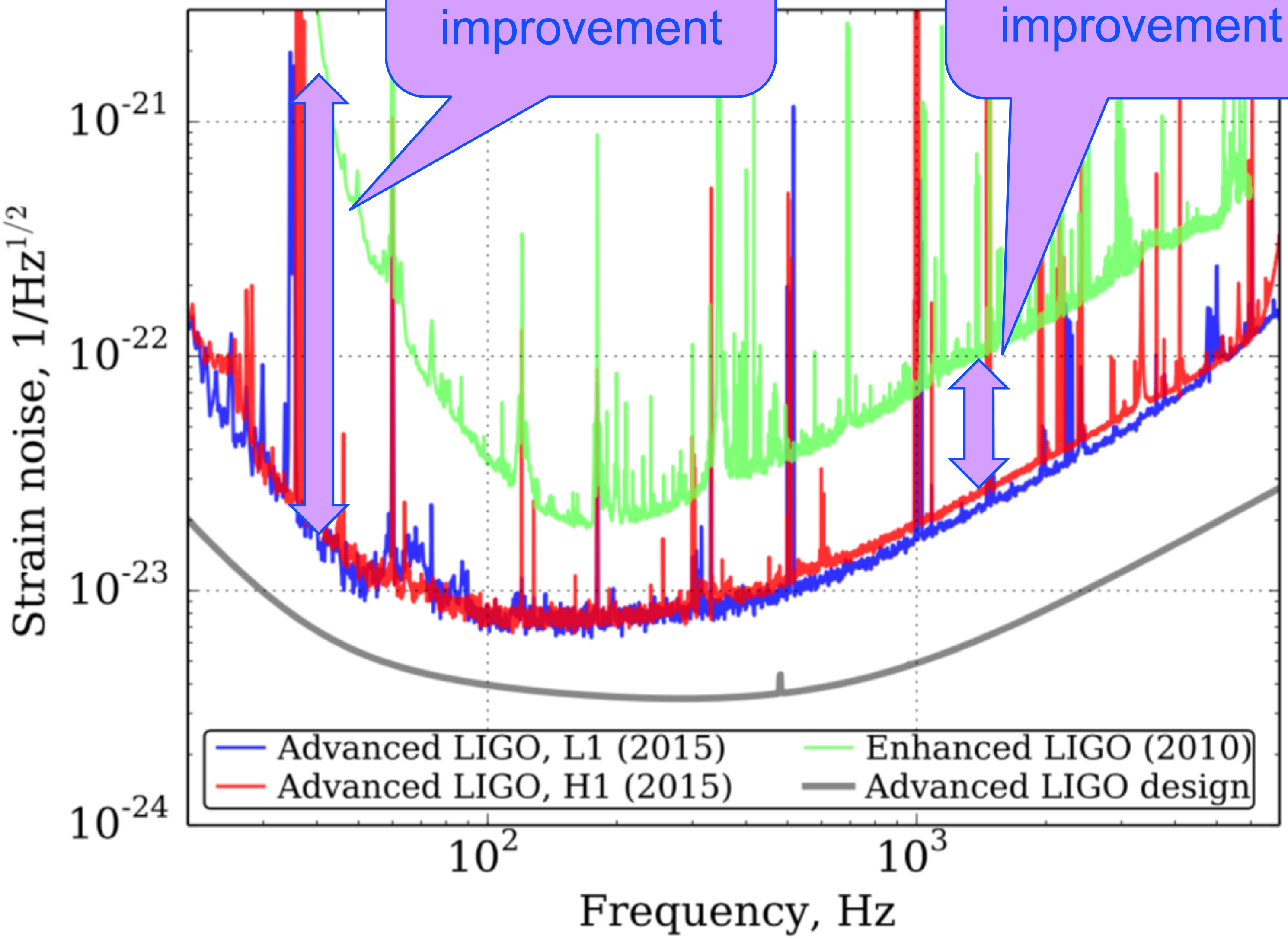


Then there are the technical noise sources….



Resolution of ~5x10^{-20} meters in 1 sec.

Sensitivity for first Observing runs



At ~40 Hz, Factor ~100 improvement

Broadband, Factor ~3 improvement

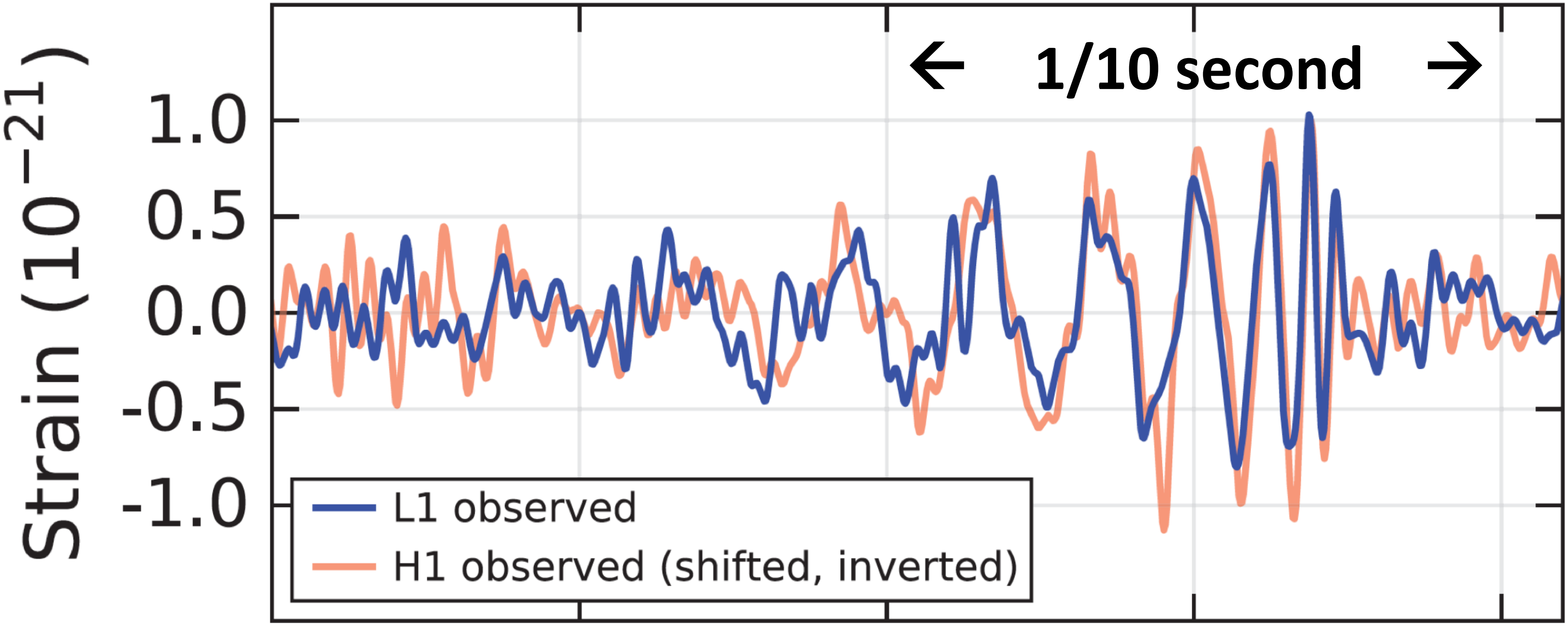
Initial LIGO

O1 aLIGO

Design aLIGO

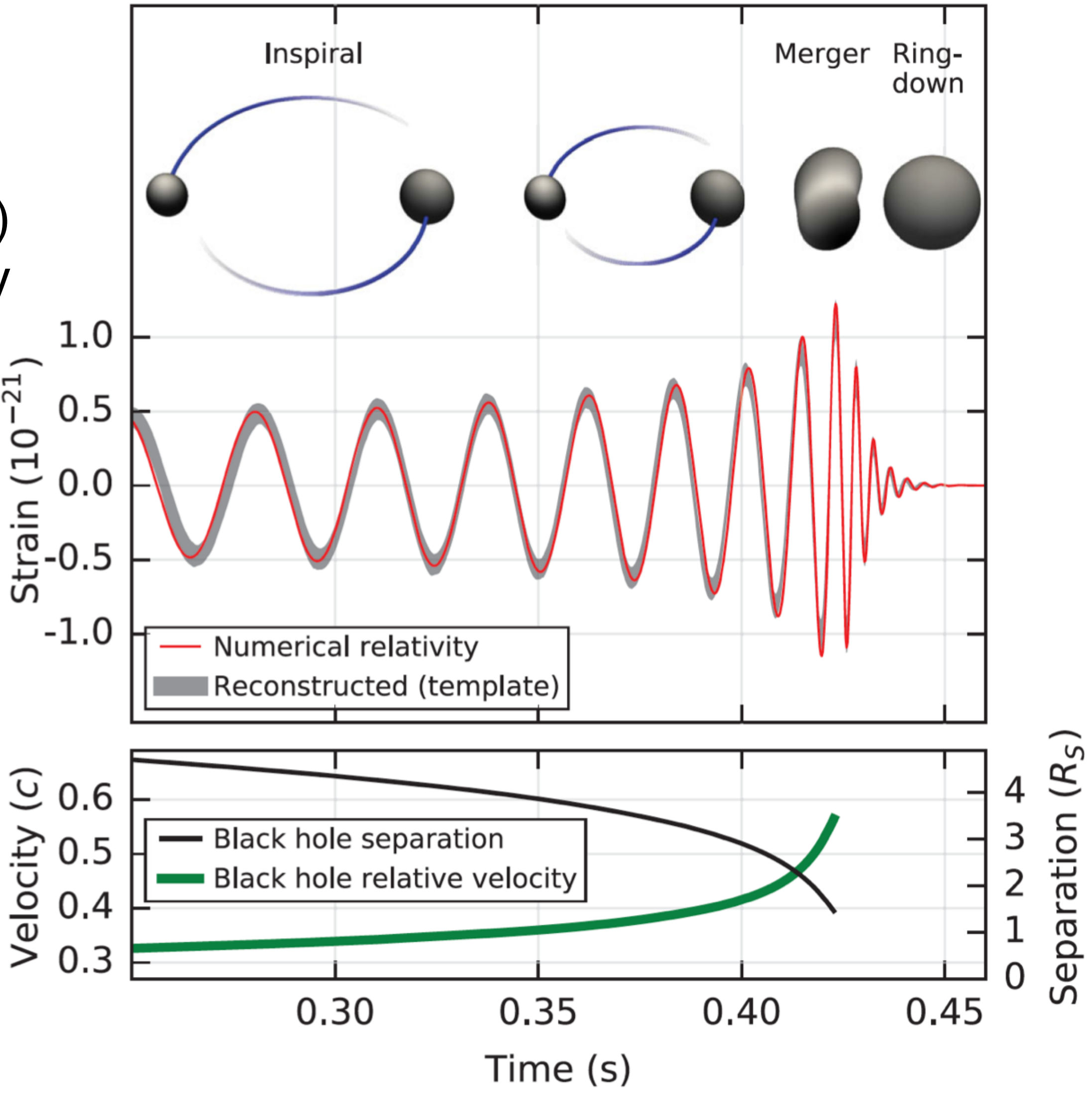
The first signal

On September 14, 2015 at 09:50:45 UTC the two detectors of the Laser Interferometer Gravitational-Wave Observatory observed a transient gravitational-wave signal



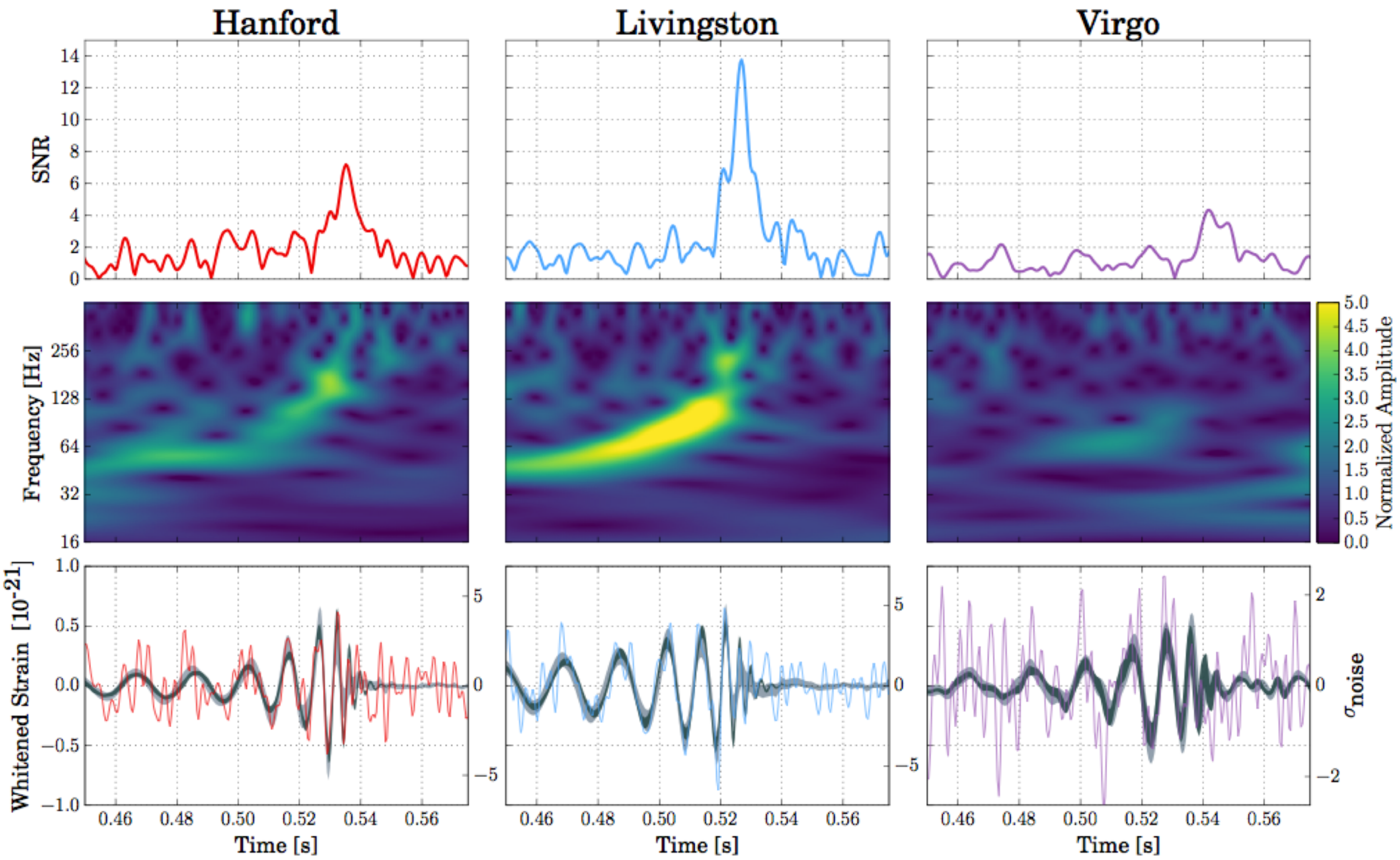
We measure *h(t)* – think ‘strip chart recorder’

- The output of the detector is the (signed) strain as a function of time
- Earlier measurements of the pulsar period decay (Taylor/Hulse/Weisberg) measured energy loss from the binary system – a beautiful experiment
  - radiation of gravitational waves confirmed to *remarkable* precision for 0th post-Newtonian
- **LIGO can actually measure the change in distance between our own test masses, due to a passing space-time ripple**
  - Instantaneous amplitude rather than time-averaged power
  - Much richer information!



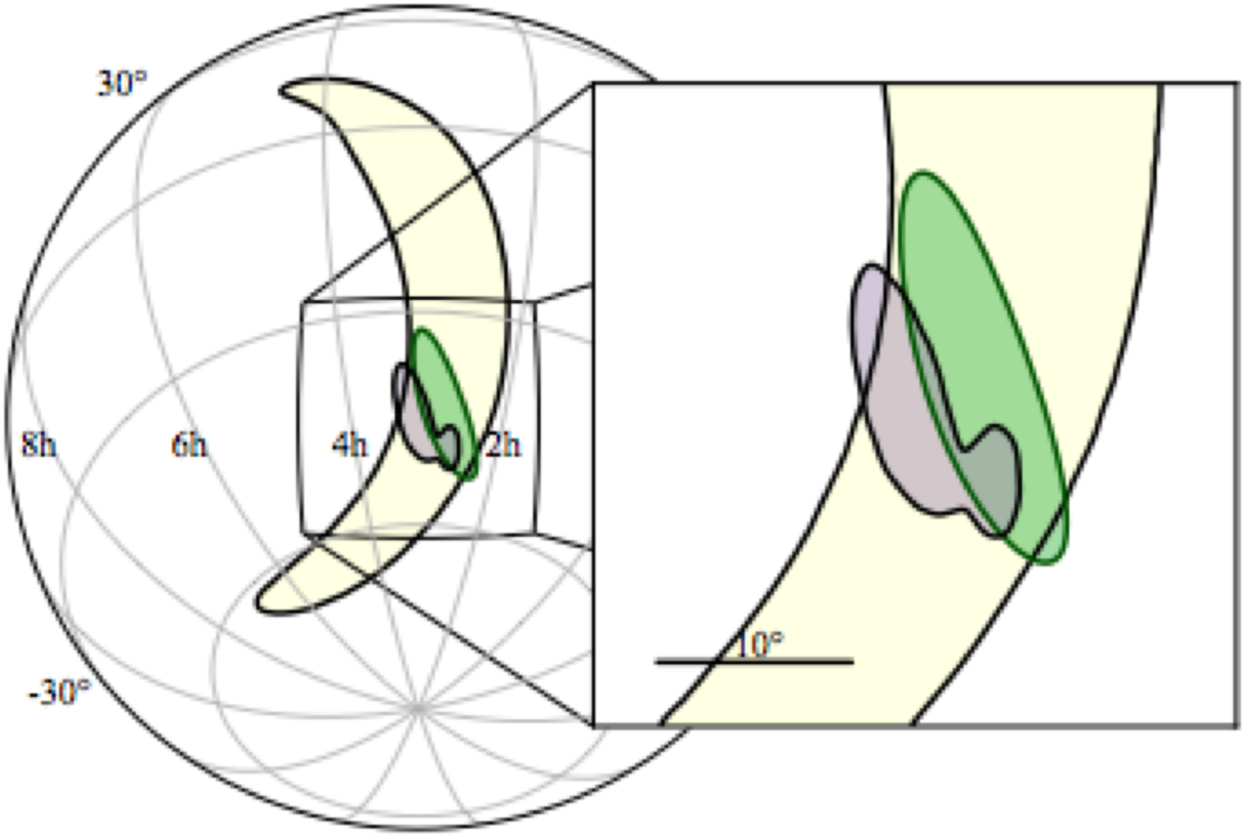
GW170814

The first GW signal observed by LIGO-Hanford, LIGO-Livingston and Virgo

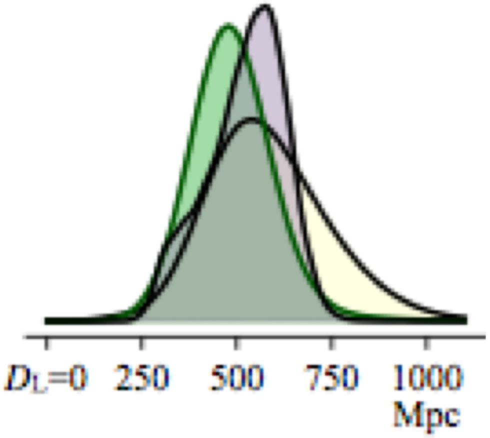


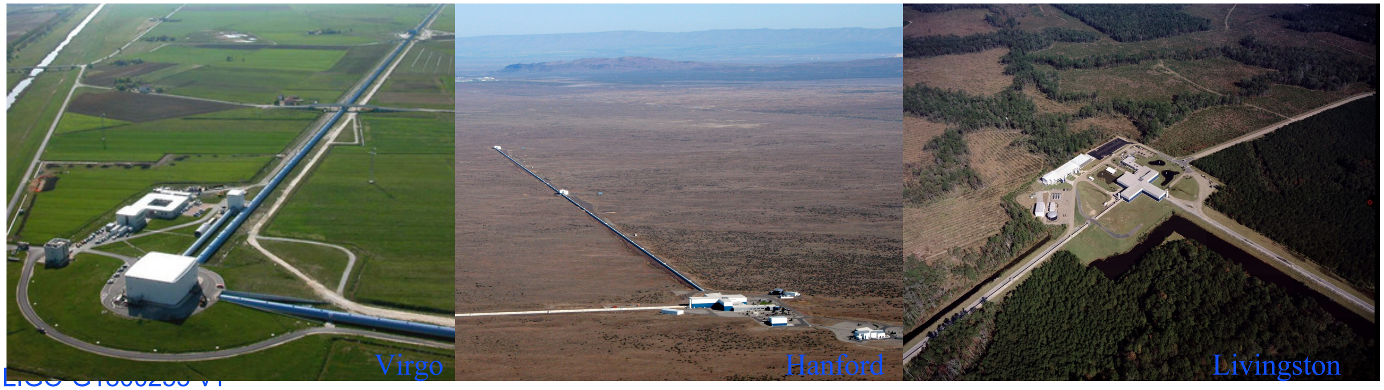
B. P. Abbott et al., *A Three-Detector Observation of Gravitational Waves from a Binary Black Hole Coalescence*, 2017, Phys. Rev. Lett., 119, 141101

GW170814



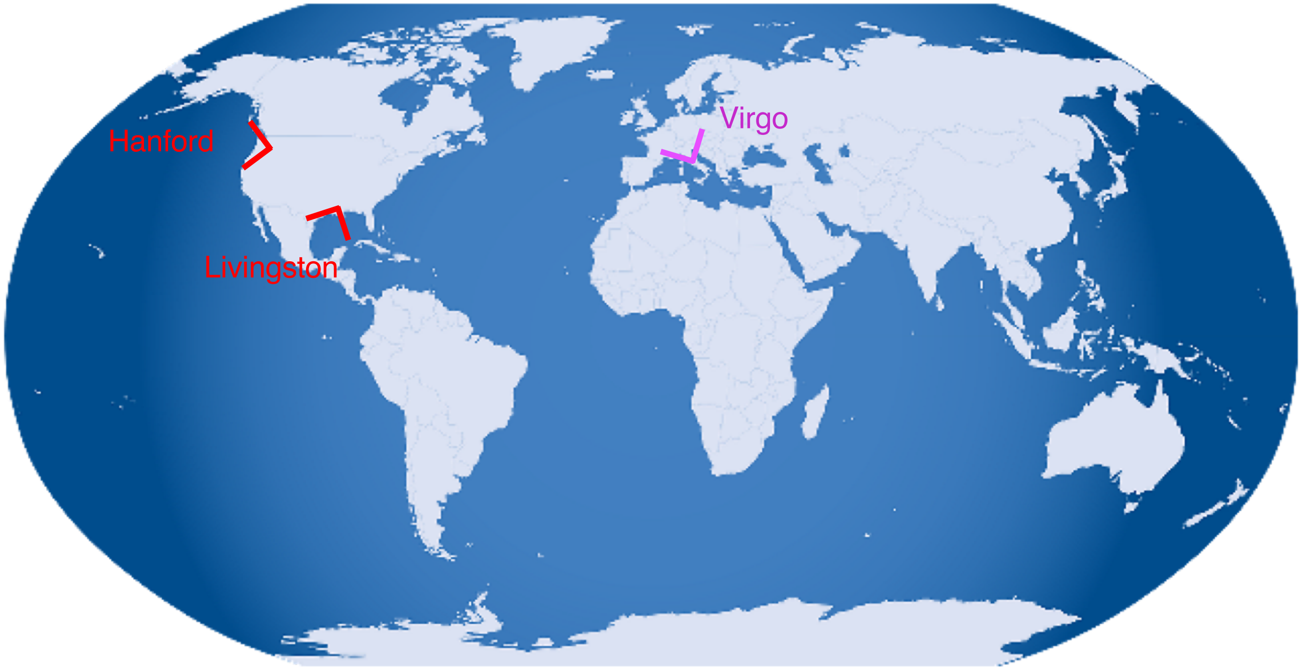
Sky localization improves ~20x





Uncertainty in volume reduced ~34x

GW170814

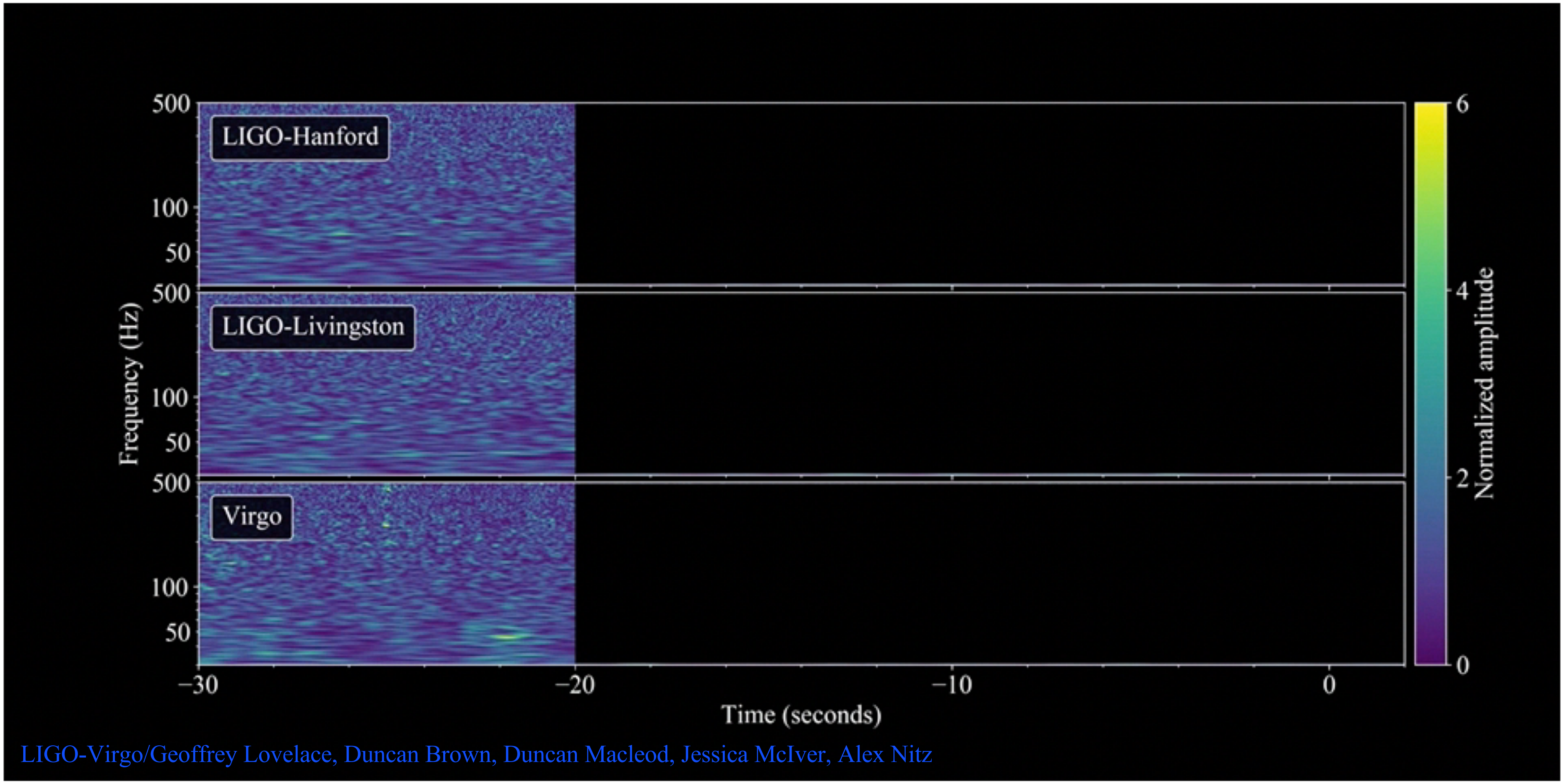


LIGO-Hanford and Livingston have similar orientations -> little information about GW polarizations

Virgo is not aligned with LIGO – giving polarization information

For GW170817, purely tensor polarization is strongly favored over purely scalar or vector polarizations – consistent with General Relativity

GW170817



LIGO-Virgo/Geoffrey Lovelace, Duncan Brown, Duncan Macleod, Jessica McIver, Alex Nitz

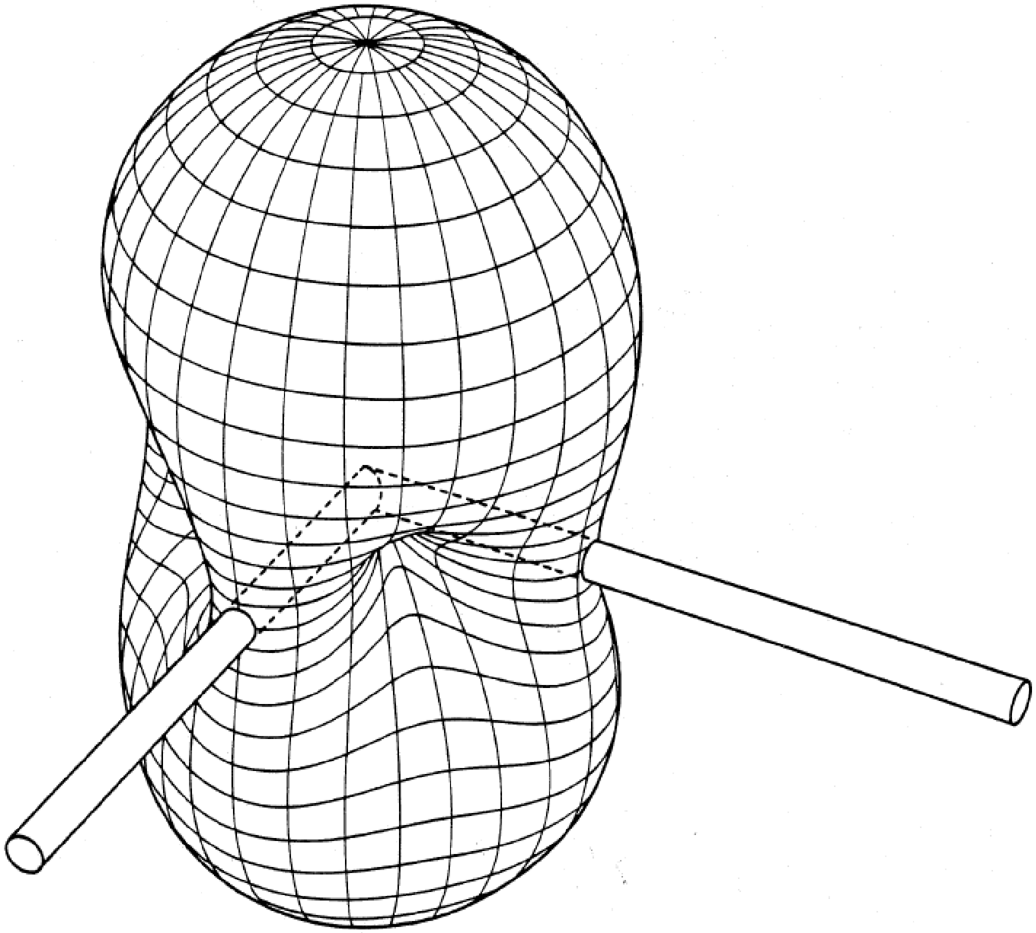
https://doi.org/10.1103/PhysRevLett.119.161101

Antenna pattern for a single detector

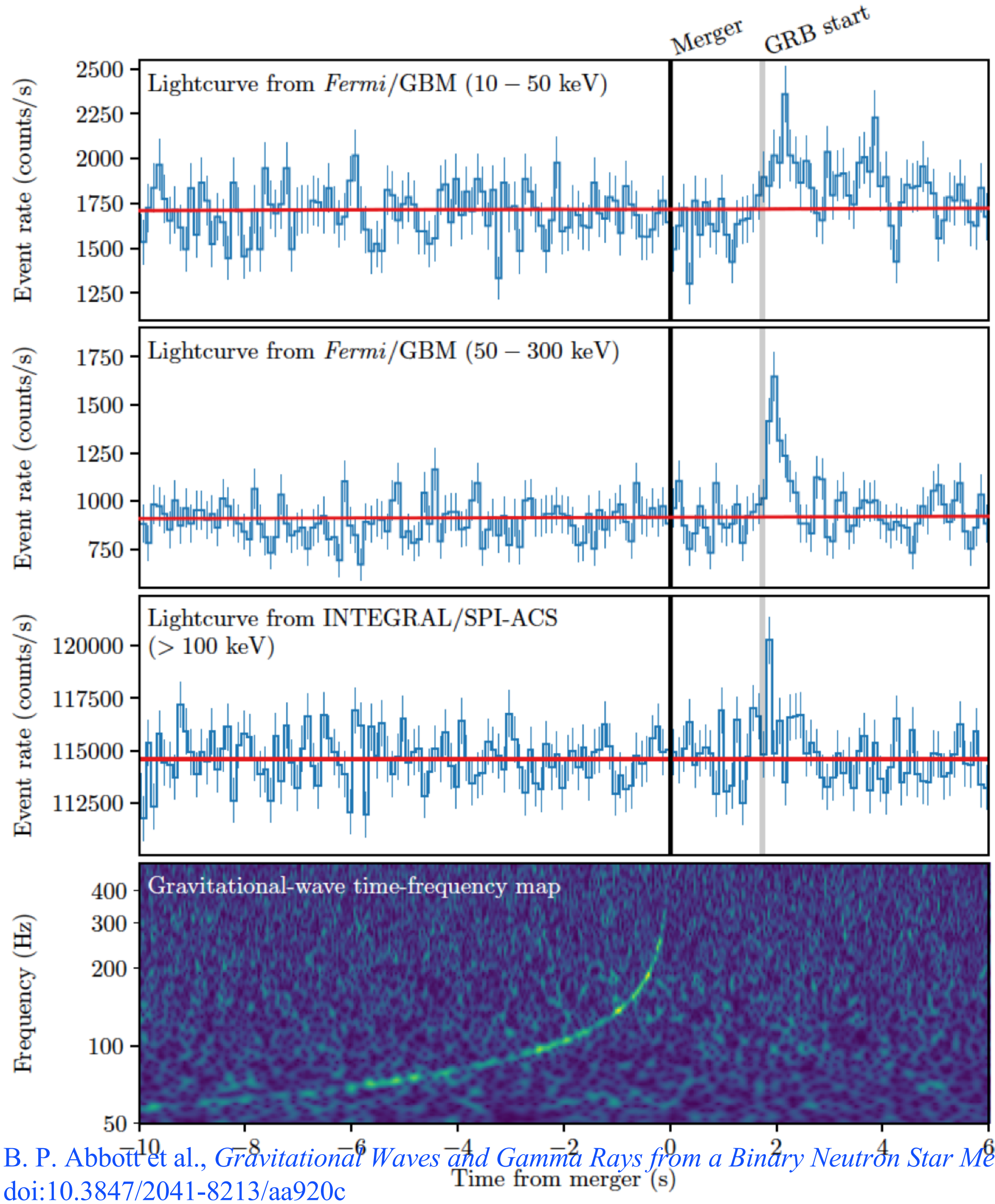
- Maximal for overhead or underfoot source
- 1/2 for signals along one arm
- …and zero at 45 degrees

- GW170817 fell on Virgo close to 45 degrees!

- Did no harm for localization. (GW170814 proved the detector was working, happily)



GRB 170817A



GRB 170817A occurs (1.74 ± 0.05) seconds after GW170817

It was autonomously detected in-orbit by Fermi-GBM (GCN was issued 14s after GRB) and in the routine untargeted search for short transients by INTEGRAL SPI-ACS

Probability that GW170817 and GRB 170817A occurred this close in time and with location agreement by chance is $5.0\text{x}10^{-8}$ (Gaussian equivalent significance of 5.3σ)

BNS mergers are progenitors of (at least some) SGRBs

B. P. Abbott et al., *Gravitational Waves and Gamma Rays from a Binary Neutron Star Merger: GW170817 and GRB 170817A*, 2017, ApJL in press.
doi:10.3847/2041-8213/aa920c

Multimessenger Observations

Approximate timeline:

GW170817 - August 17, 2017 12:41:04 UTC = t_0

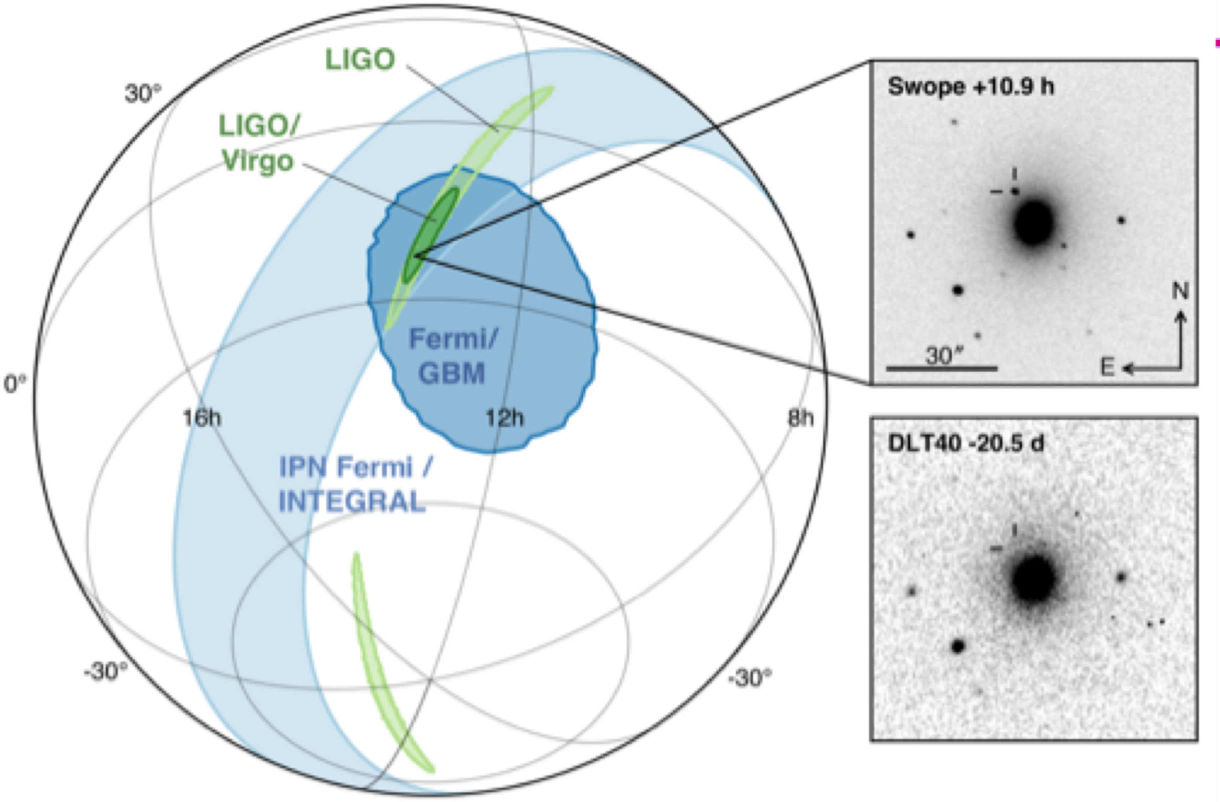
GRB 170817A
t_0 + 2 sec

LIGO signal found
t_0 +6 minutes

LIGO-Virgo GCN reporting BNS signal associated with the time of the GRB
t_0 +41 minutes

SkyMap from LIGO-Virgo
t_0 + 4 hours

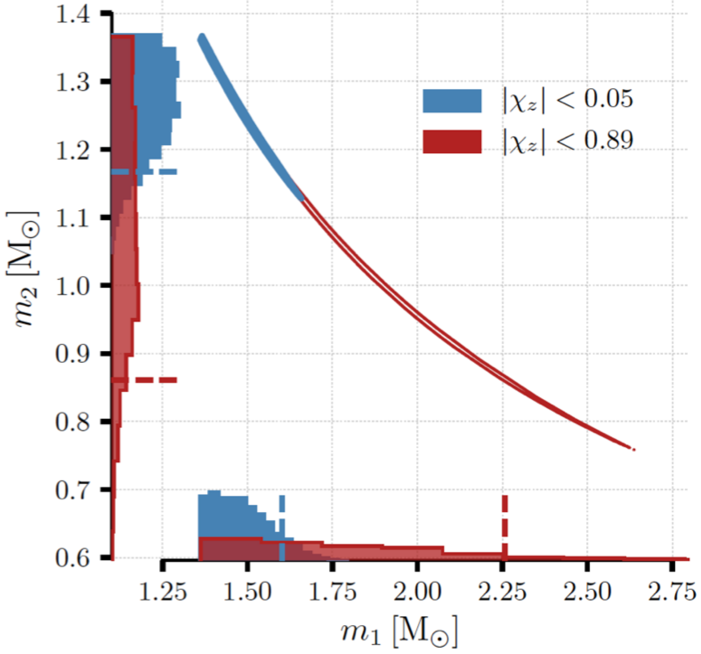
Optical counterpart found
t_0 + 11 hours



- The localisation region became observable to telescopes in Chile 10 hours after the event time (wait for nightfall!)
- Approximately 70 ground- and space- based observatories followed-up on this event

B. P. Abbott, *Multi-messenger Observations of a Binary Neutron Star Merger*, 2017, ApJL in press. doi:10.3847/2041-8213/aa91c9

BNS properties



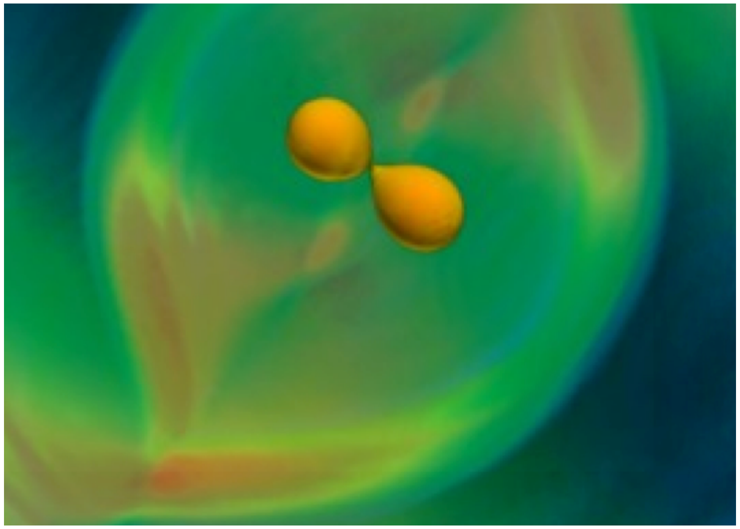
Sum of NS masses tightly constrained, individual masses less so

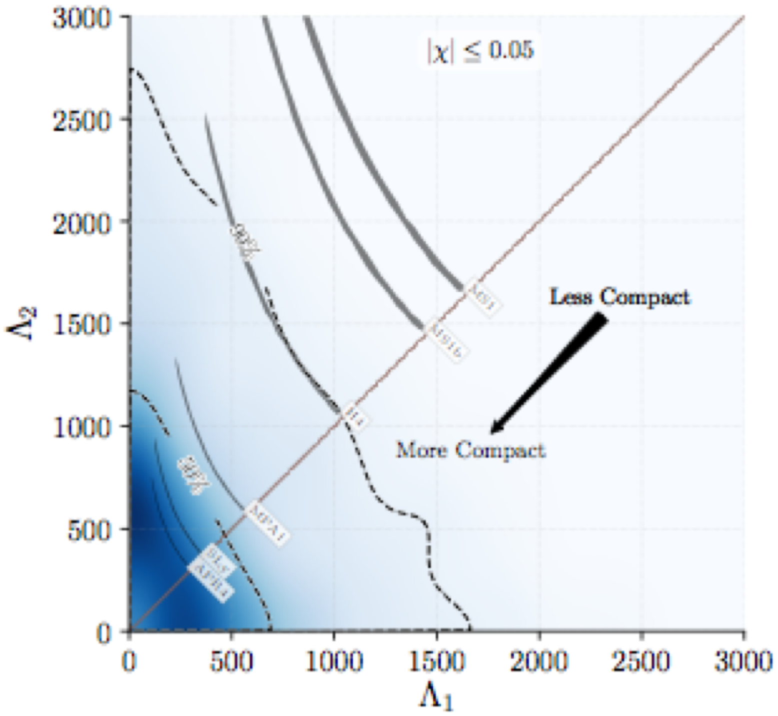
- |χ| ≤ 0.89 limit imposed by available rapid waveform models
- |χ| ≤ 0.05 limit consistent with the observed population of BNS

	Low-spin priors ($|\chi| \leq 0.05$)
Primary mass m_1	$1.36 - 1.60\ M_\odot$
Secondary mass m_2	$1.17 - 1.36\ M_\odot$
Chirp mass $\mathcal{M}$	$1.188^{+0.004}_{-0.002}\ M_\odot$
Mass ratio m_2/m_1	$0.7 - 1.0$
Total mass m_{tot}	$2.74^{+0.04}_{-0.01}\ M_\odot$
Radiated energy E_{rad}	$> 0.025\ M_\odot c^2$

B. P. Abbott et al., Phys. Rev. Lett., 119, 161101

Neutron star equation-of-state





- Tidal disruption is encoded in the BNS gravitational waveform
- For this event, mostly masked by high-frequency noise in detector
- Some constraints possible

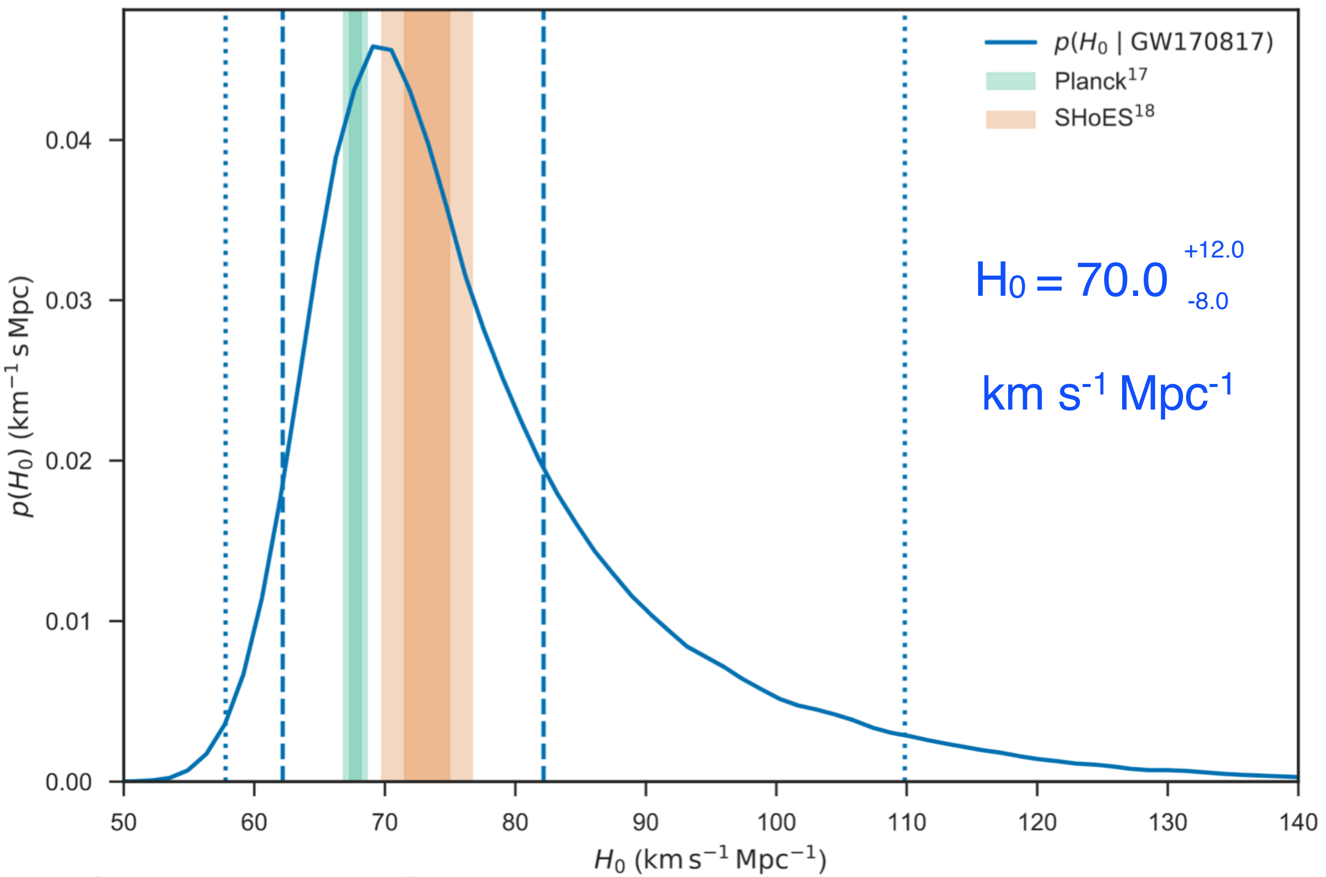
tidal deformability parameter $\Lambda \sim k_2 (R/m)^5$

k_2 - second Love number

R, m = radius, mass of the neutron star

GW170817: Observation of Gravitational Waves from a Binary Neutron Star Inspiral
Phys. Rev. Lett., 119:161101, 2017

GWs as standard sirens: Hubble Constant



$H_0 = 70.0^{+12.0}_{-8.0}$

km s^{-1} Mpc^{-1}

v_H - local “Hubble flow” velocity of the source

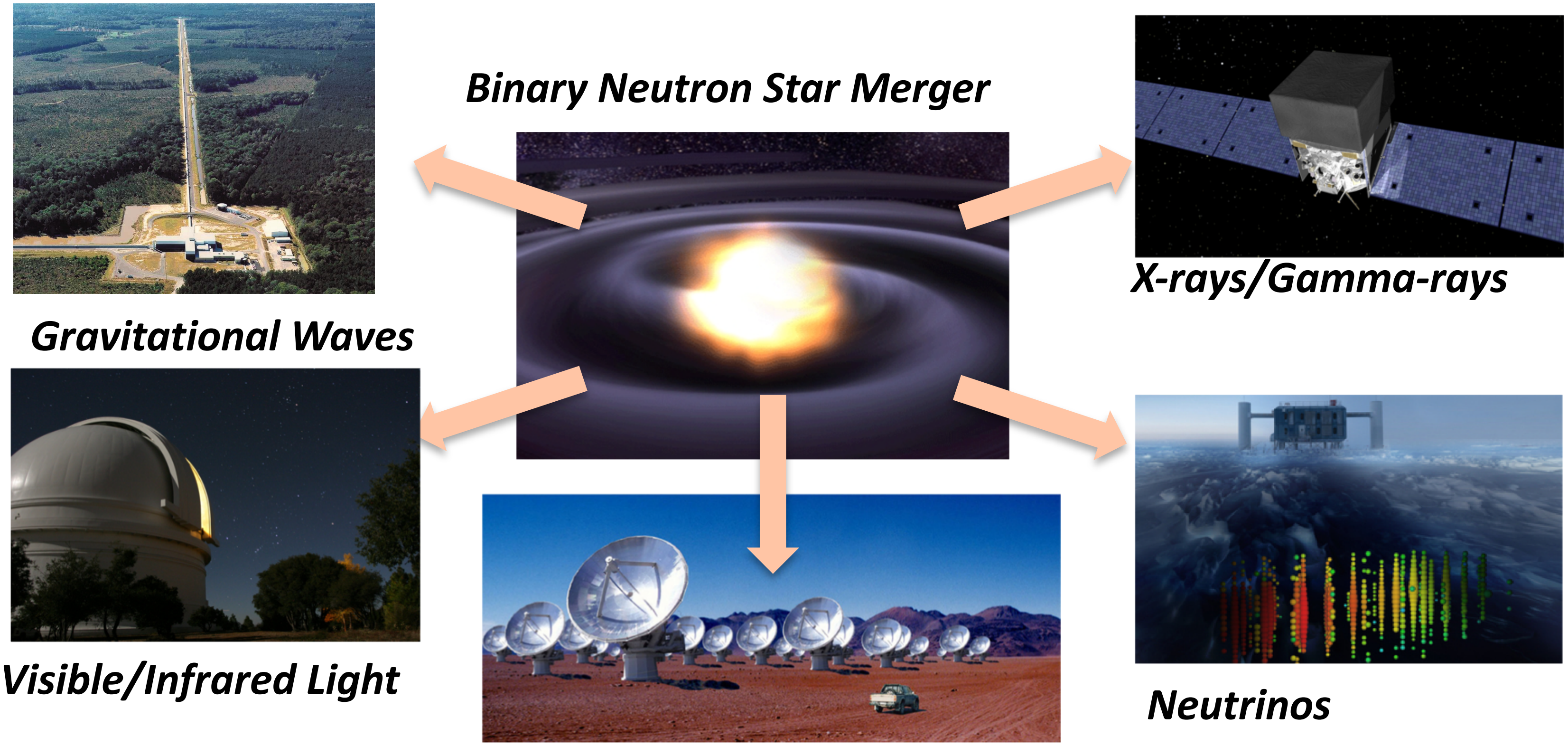
Use optical identification of the host galaxy NGC 4993

$$\boxed{v_H = H_0\, d}$$

d - distance to the source
Use the GW distance estimate

B. P. Abbott et al., *A gravitational-wave standard siren measurement of the Hubble constant*, 2017, Nature. doi:10.1038/nature24471

Multi-messenger Astronomy



Binary Neutron Star Merger

Gravitational Waves

X-rays/Gamma-rays

Visible/Infrared Light

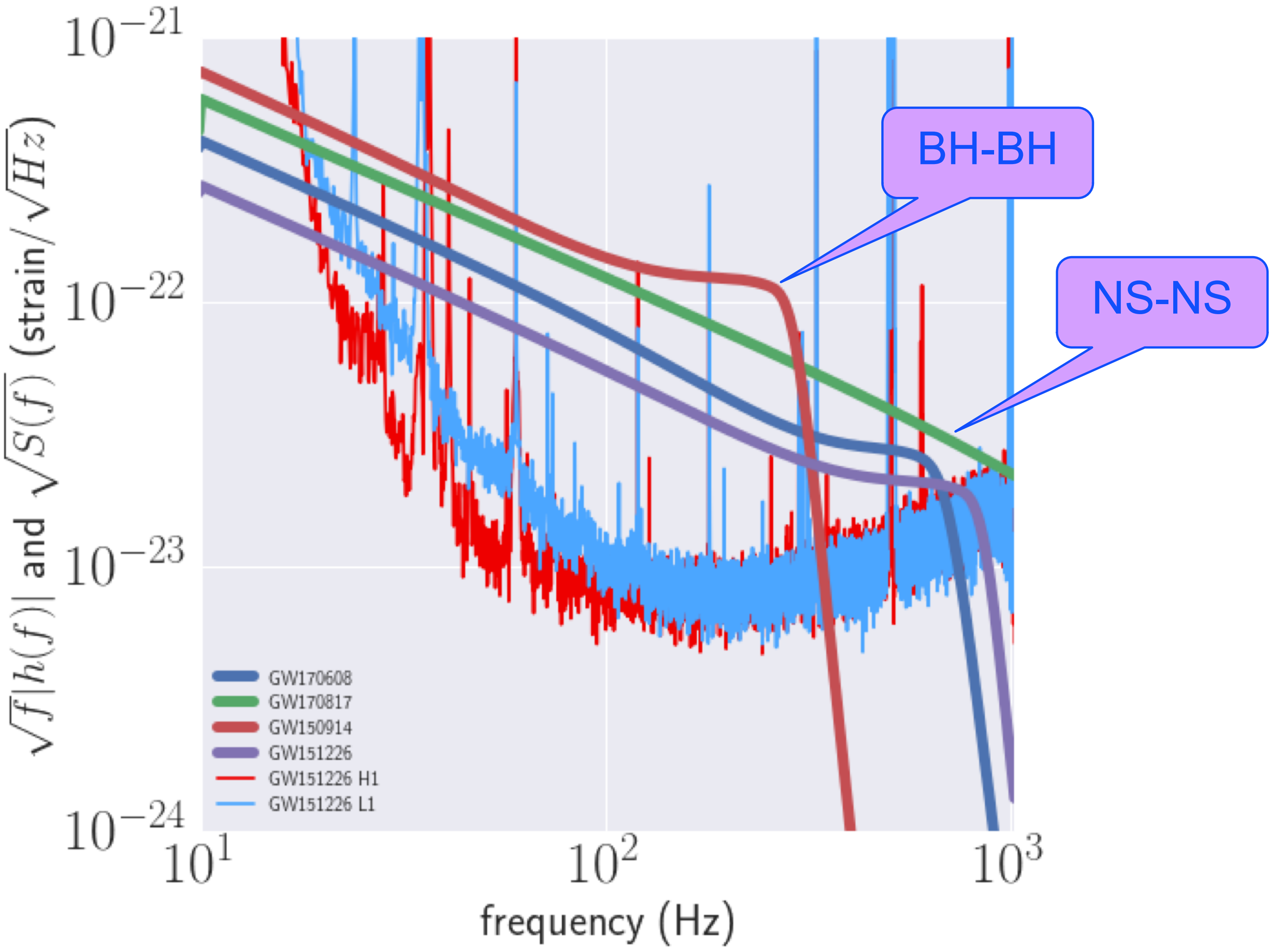
Neutrinos

Radio Waves

LIGO and Virgo signed agreements with 95 groups for EM/neutrino followup of GW events

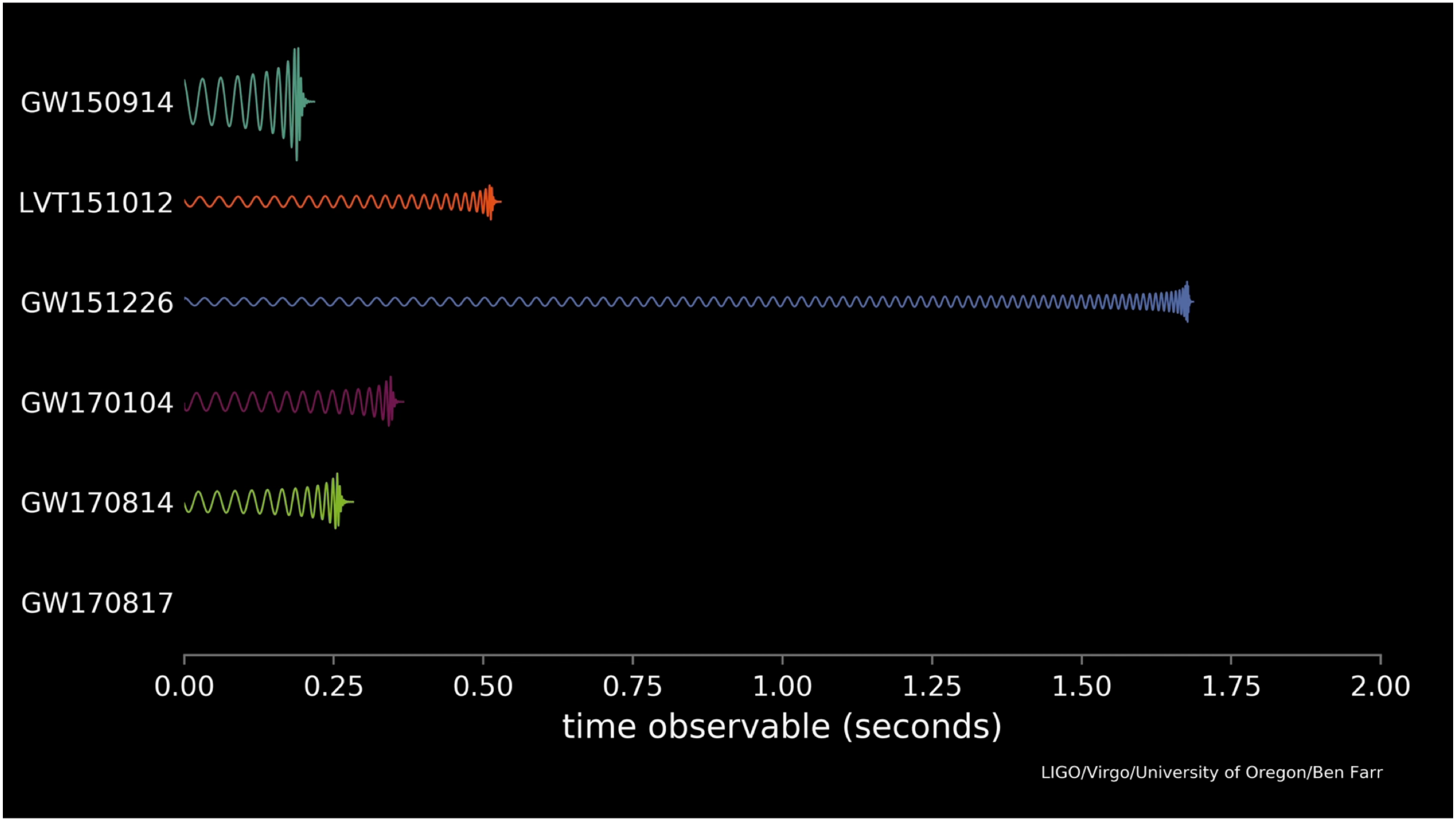
- ~200 EM instruments - satellites and ground based telescopes covering the full spectrum from radio to very high-energy gamma-rays
- Worldwide astronomical institutions, agencies and large/small teams of astronomers

Events in the detector response context



BH-BH

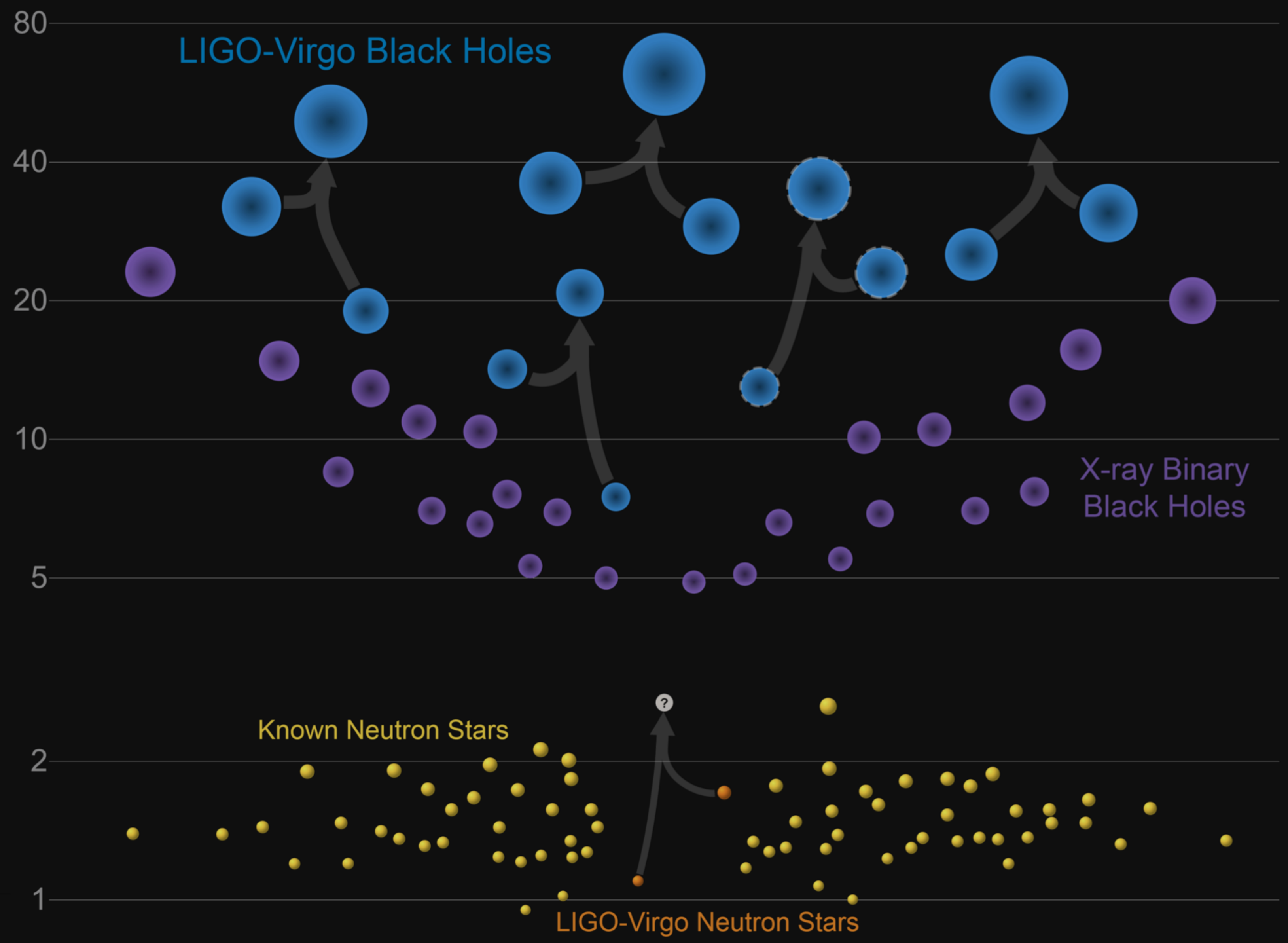
NS-NS



LIGO/Virgo/University of Oregon/Ben Farr

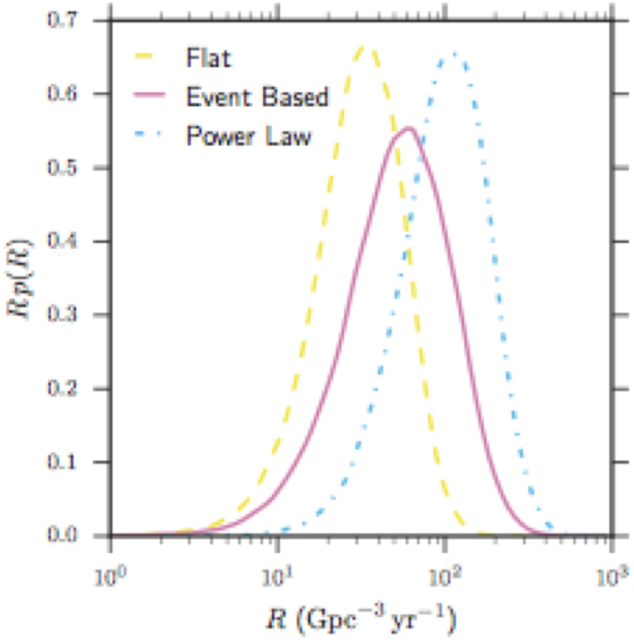
Masses in the Stellar Graveyard

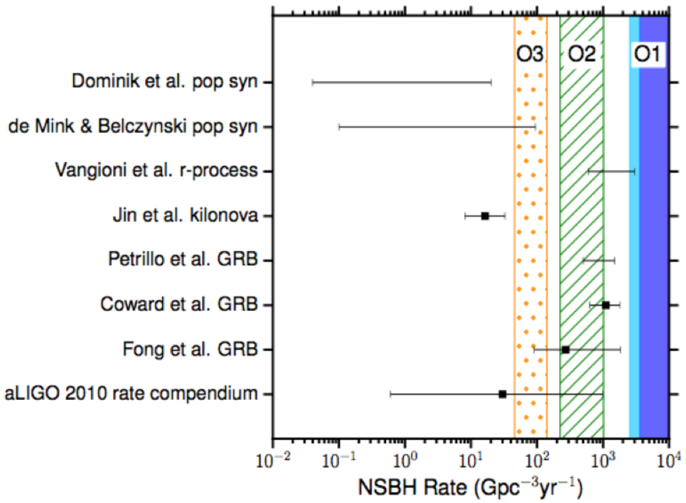
in Solar Masses



LIGO-Virgo/Frank Elavsky/Northwestern

Rates of compact object mergers





Binary Black Hole Merger Rate

- Based on O1 BBH mergers: 9-240 Gpc^{-3} yr^{-1}
- Addition of GW170104, BBH merger rate: 12-213 Gpc^{-3} yr^{-1}
- Observation of GW170814 consistent with this population

Binary Neutron Star Merger Rate

- Based on O1 non-detections: < 12,600 Gpc^{-3} yr^{-1}
- Based on GW170817: 320-4740 Gpc^{-3} yr^{-1}

Neutron Star - Black Hole Merger Rate

- Based on O1 non-detections (black hole mass at least 5 M⊙): < 3,600 Gpc^{-3} yr^{-1}

B. P. Abbott et al., *Binary Black Hole Mergers in the First Advanced LIGO Observing Run, 2016,* Phys. Rev. X 6, 041015

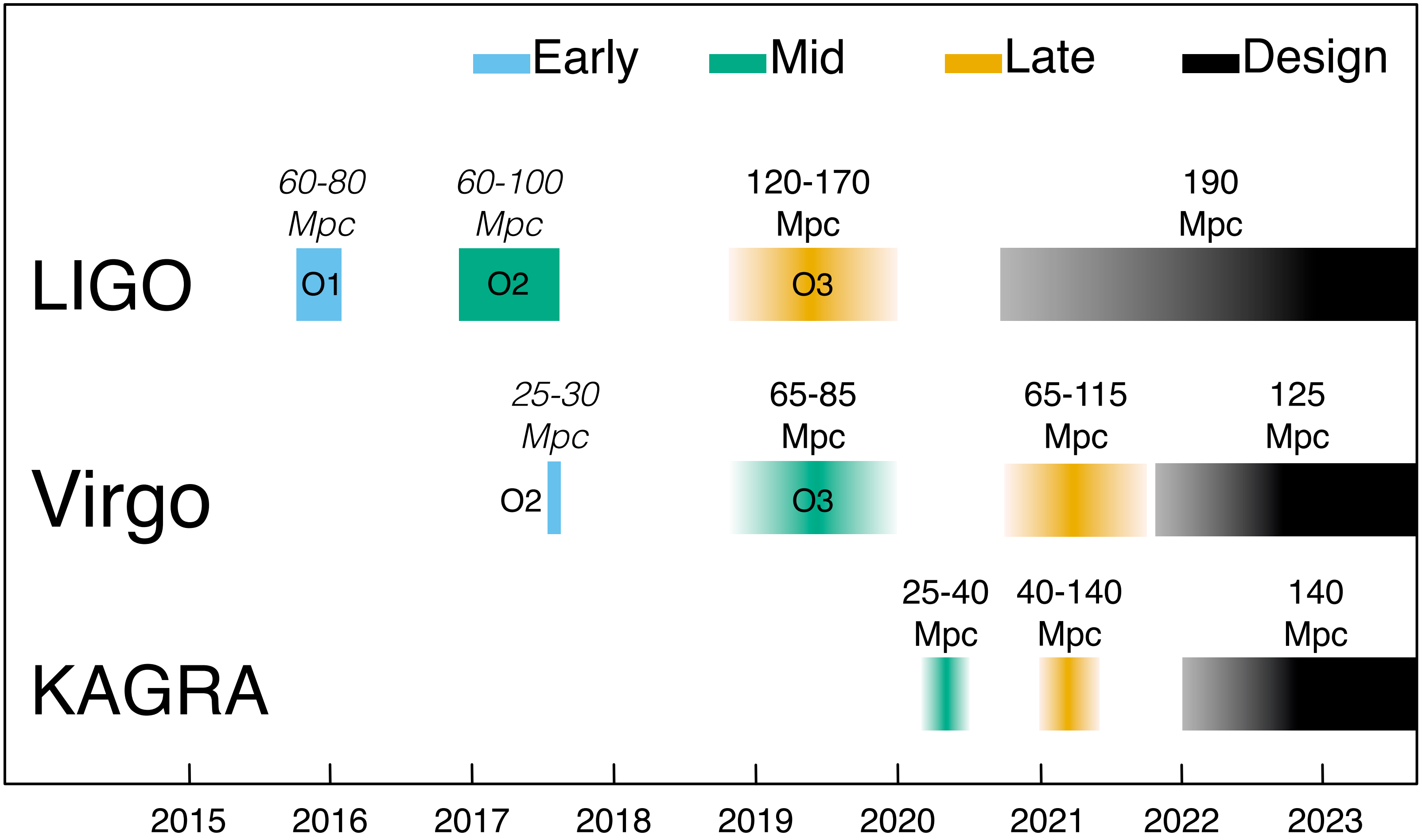
B. P. Abbott et al., *GW170401: Observation of a 50-Solar-Mass Binary Black Hole Coalescence at Redshift 0.2*, 2017, Phys. Rev. Lett., 118, 221101

B. P. Abbott et al., *Upper limits on the rates of binary neutron star and neutron-star--black-hole mergers from Advanced LIGO's first observing run*, 2016, ApJL, L21

Work continues on analyzing O2….
keep tuned!

Plausible Observing Timeline

Binary Neutron Star Range



B. P. Abbott et al., *Prospects for Observing and Localizing Gravitational-Wave Transients with Advanced LIGO, Advanced Virgo and KAGRA*, 2016, Living Rev. Relativity 19

LIGO Scientific Collaboration and Virgo Collaboration

~1500 members, ~120 institutions, 21 countries

LIGO VIRGO

LIGO Scientific Collaboration

LSC

Caltech

American University Washington, DC

Andrews University

Washington State University

California State University Fullerton

UTRGV

Southern University and Agricultural & Mechanical College

Massachusetts Institute of Technology 1861

Leland Stanford Junior University 1891

The University Of Sheffield.

University of Maryland 1856

indigo

Texas Tech University.

PennState

清華大学 Tsinghua University

Montclair State University

Max Planck Institute for Gravitational Physics Albert Einstein Institute

International Institute of Physics Federal University of Rio Grande do Norte

Cal State LA

Bellevue College

UIB Universitat de les Illes Balears

University of Cambridge

Université de Montréal

Polytechnique Montréal

Monash University

University of Southampton

The University of Adelaide Australia

NCSA

W Bothell

Michigan

Universitas Bruxellensis

ACU

University of Wisconsin UWMilwaukee

Trinity University

Australian National University

University of the West of Scotland UWS

University of Glasgow

Cornell University Founded A.D. 1865

Columbia University in the City of New York

Lomonosov Moscow State University

Syracuse University Founded AD 1870

OzGrav

The University of Western Australia

R·I·T

Sonoma State University

Tsing Hua University

Cardiff University Prifysgol Caerdydd

NASA Marshall Space Flight Center

The University of Chicago

Whitman College

Montana State University

University of Washington

Università degli Studi del Sannio

University of Minnesota

Carleton College Northfield, Minn. A.D. 1866

Charles Sturt University

LSU Louisiana State University

INPE

University of Birmingham

King's College London

University of Zurich

University of Strathclyde

ACIGA

Universität Hamburg der Forschung | der Lehre | der Bildung

東京大学 The University of Tokyo

Georgia Institute of Technology

Korean Gravitational-Wave Group

Northwestern

UF University of Florida

Science & Technology Facilities Council Rutherford Appleton Laboratory

The University of Melbourne

Monash University

GEO 600

Embry-Riddle Aeronautical University

Leibniz Universität Hannover

Swinburne University of Technology

CITA ICAT Canadian Institute for Theoretical Astrophysics L'institut Canadien d'astrophysique théorique

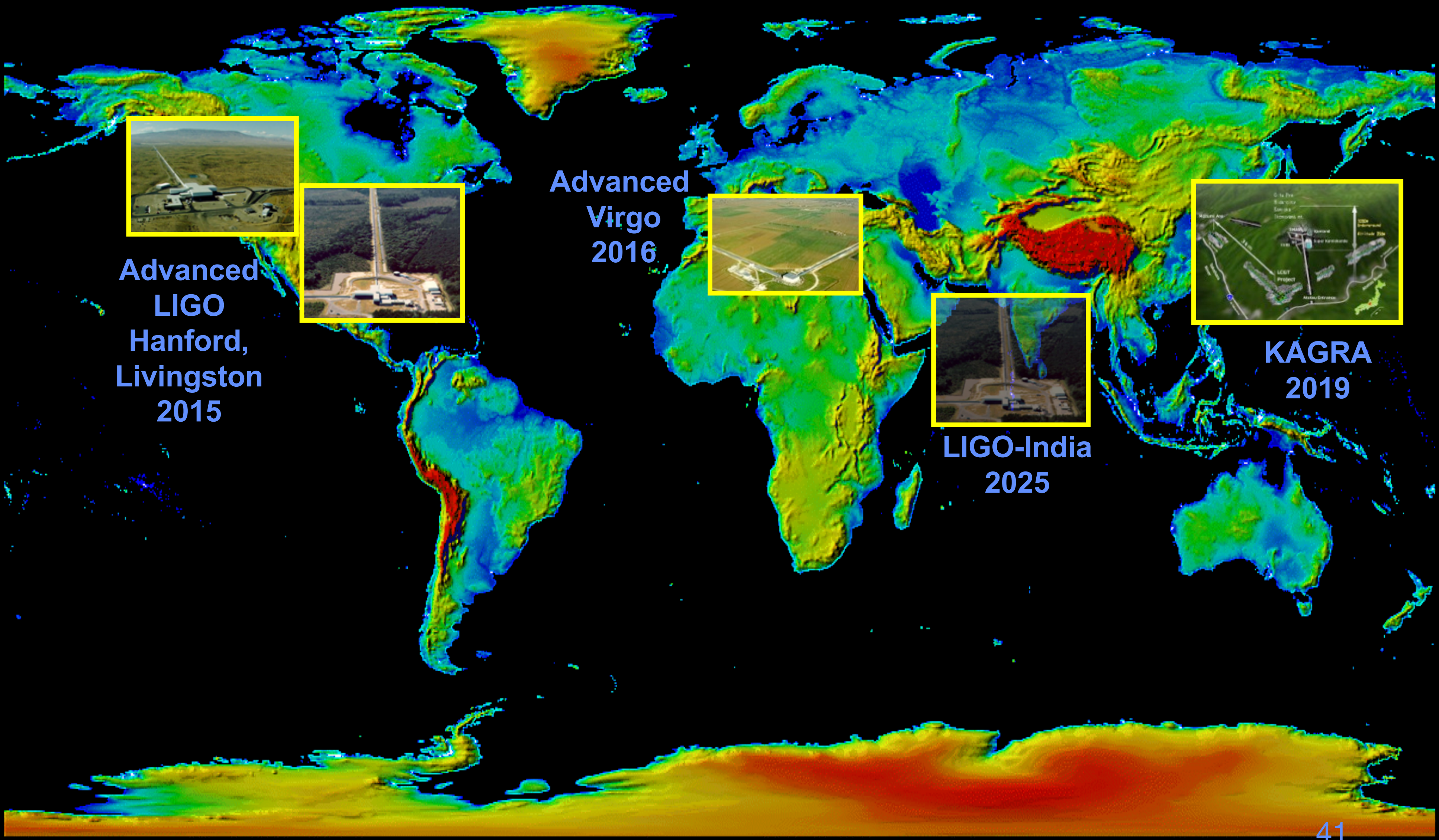
DISCE

The University of Mississippi 1848

NASA Goddard Space Flight Center

What does the future hold?

The advanced GW detector network



Advanced LIGO Hanford, Livingston 2015

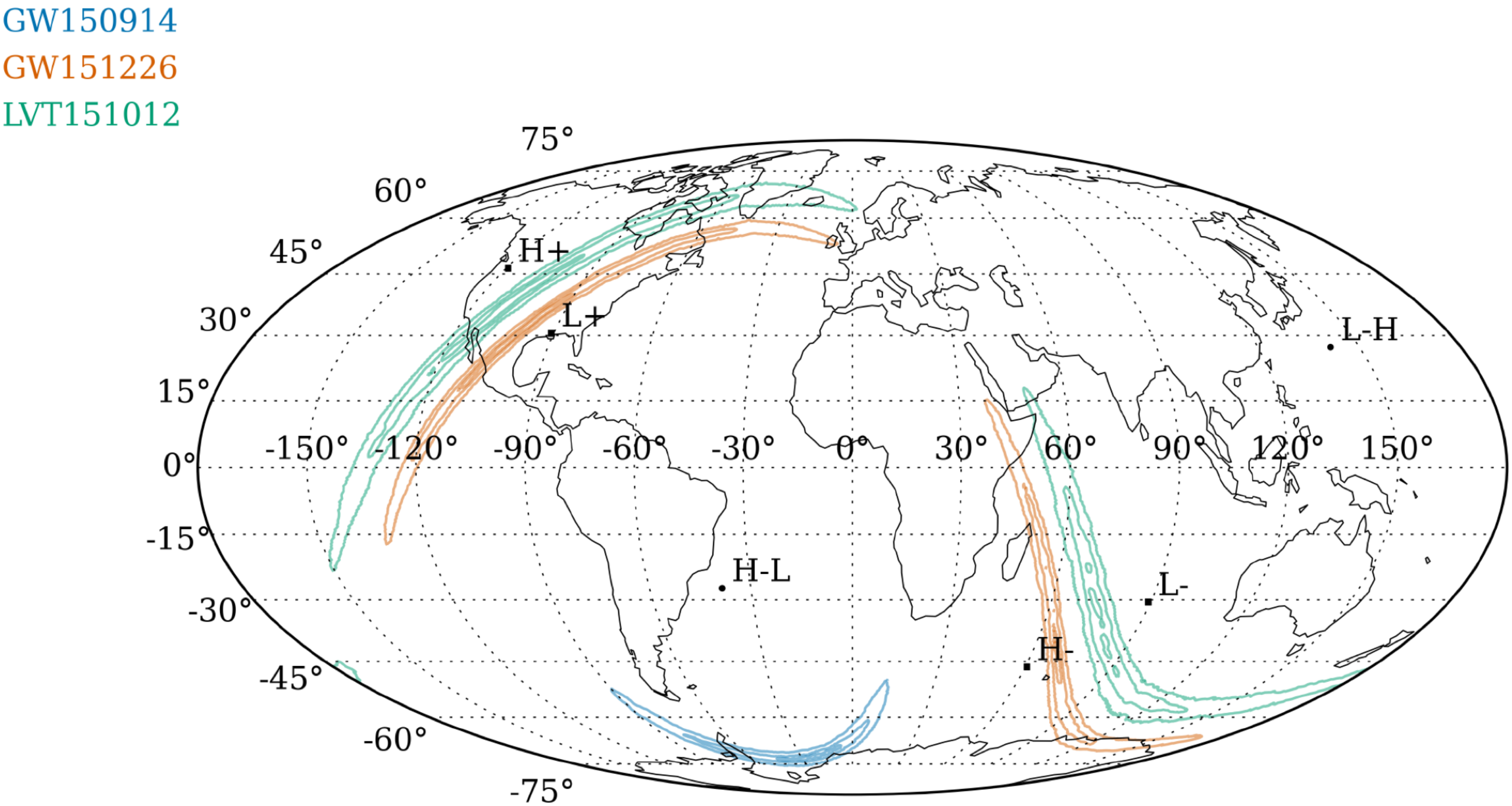
Advanced Virgo 2016

LIGO-India 2025

KAGRA 2019

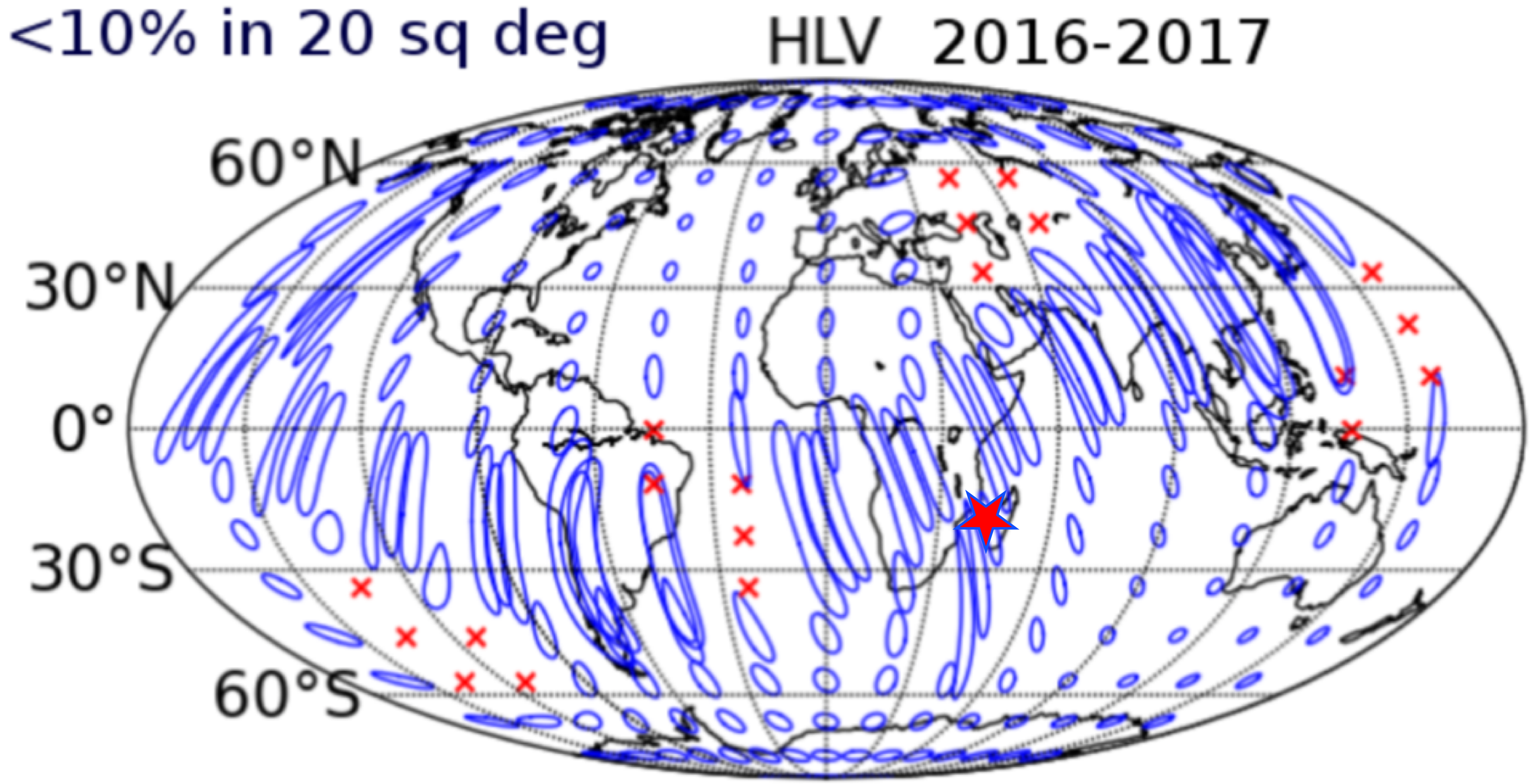
First Detection Sensitivity/configuration:

2 detectors, 1/3 goal sensitivity -- saw ~3+ signals in ~6 months of observation



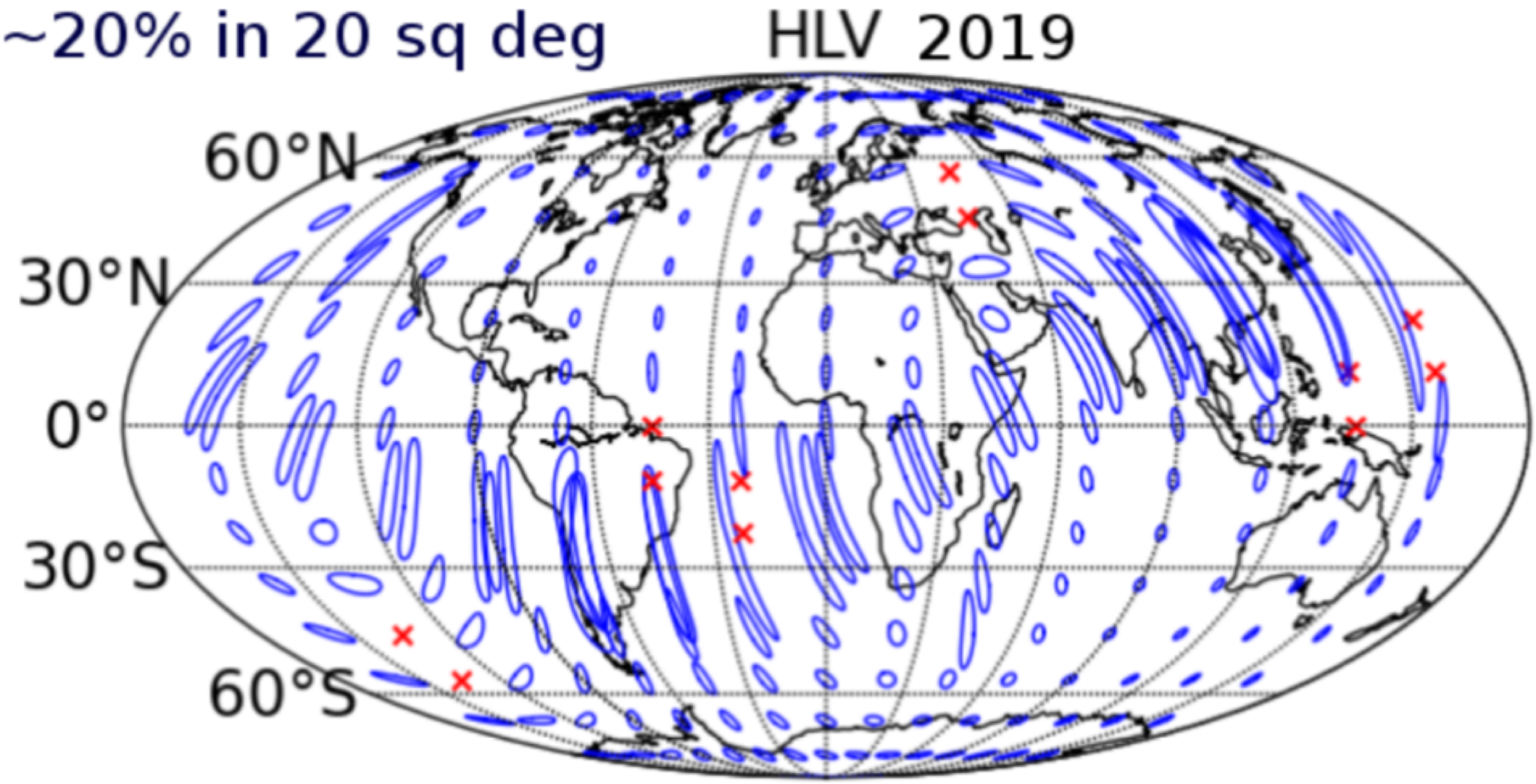
2017 Sensitivity/configuration:

3 detectors (adding Virgo), see
~1 signal per month of observation
GW170817 marked with ★



2018-19 Sensitivity/configuration:

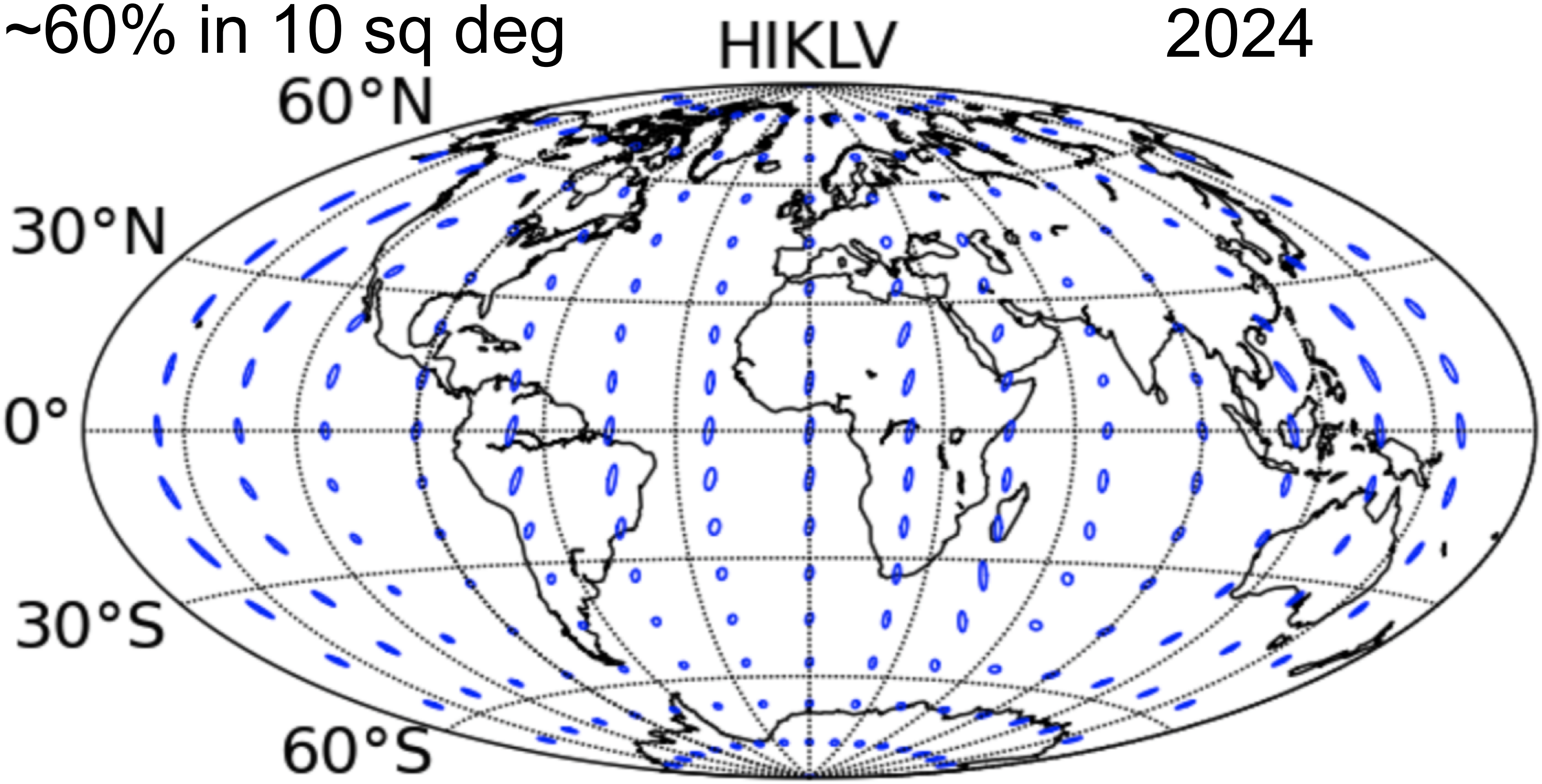
3 detectors, perhaps
~1-2 signals *per week*



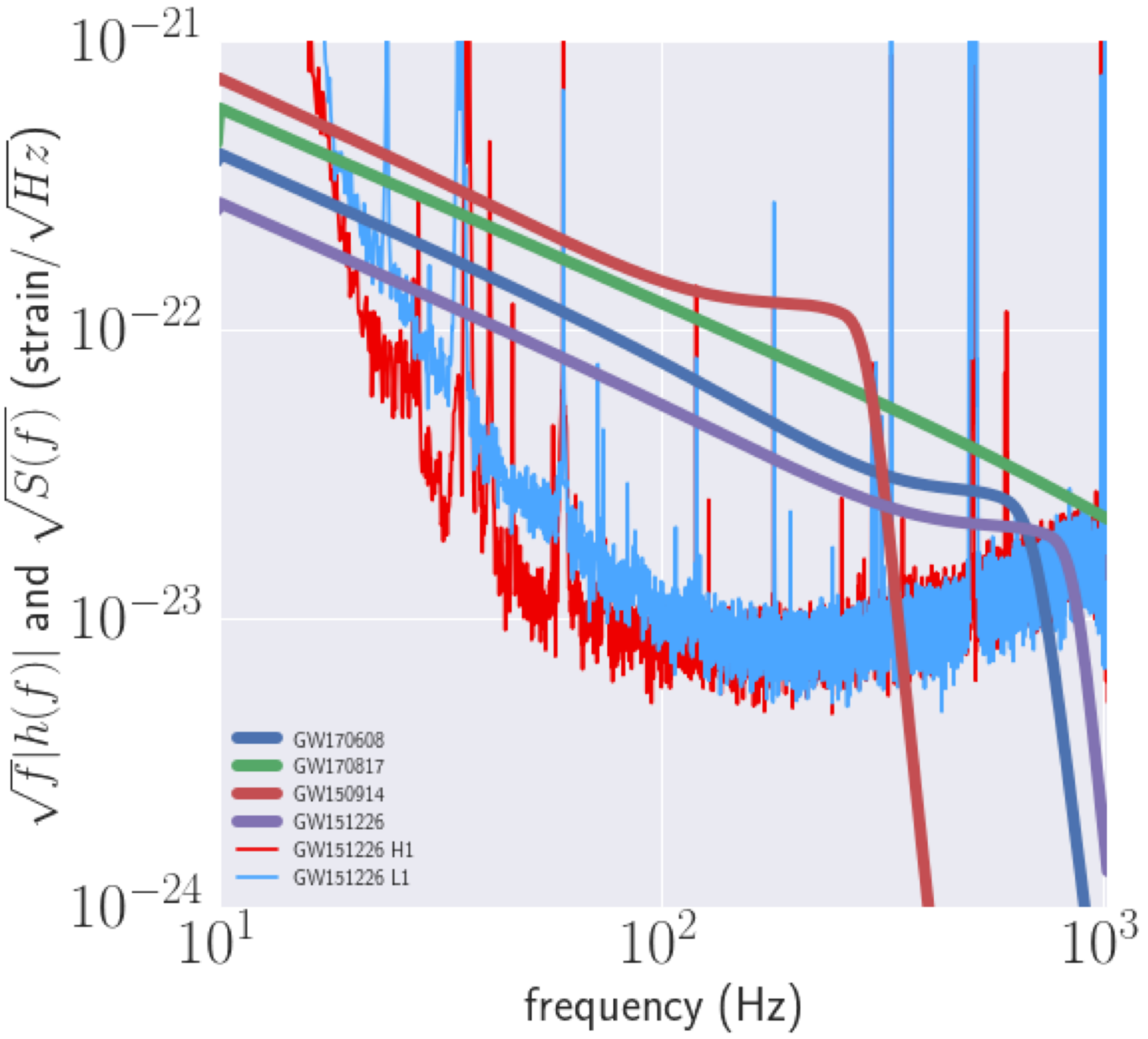
2024 Sensitivity/configuration:

5 detectors (add India and Japan)
far improved source localization

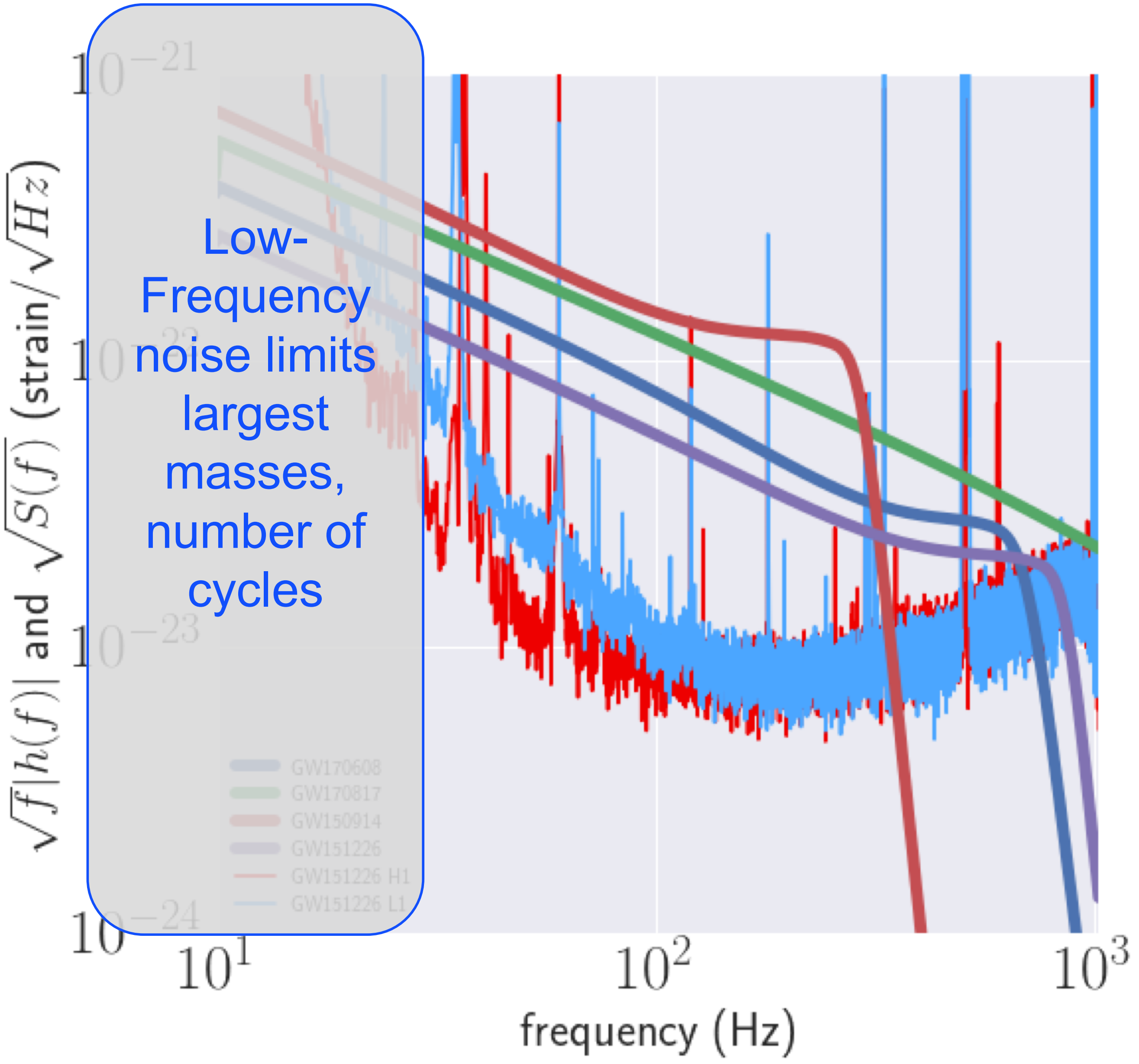
~60% in 10 sq deg



Evolution of detectors

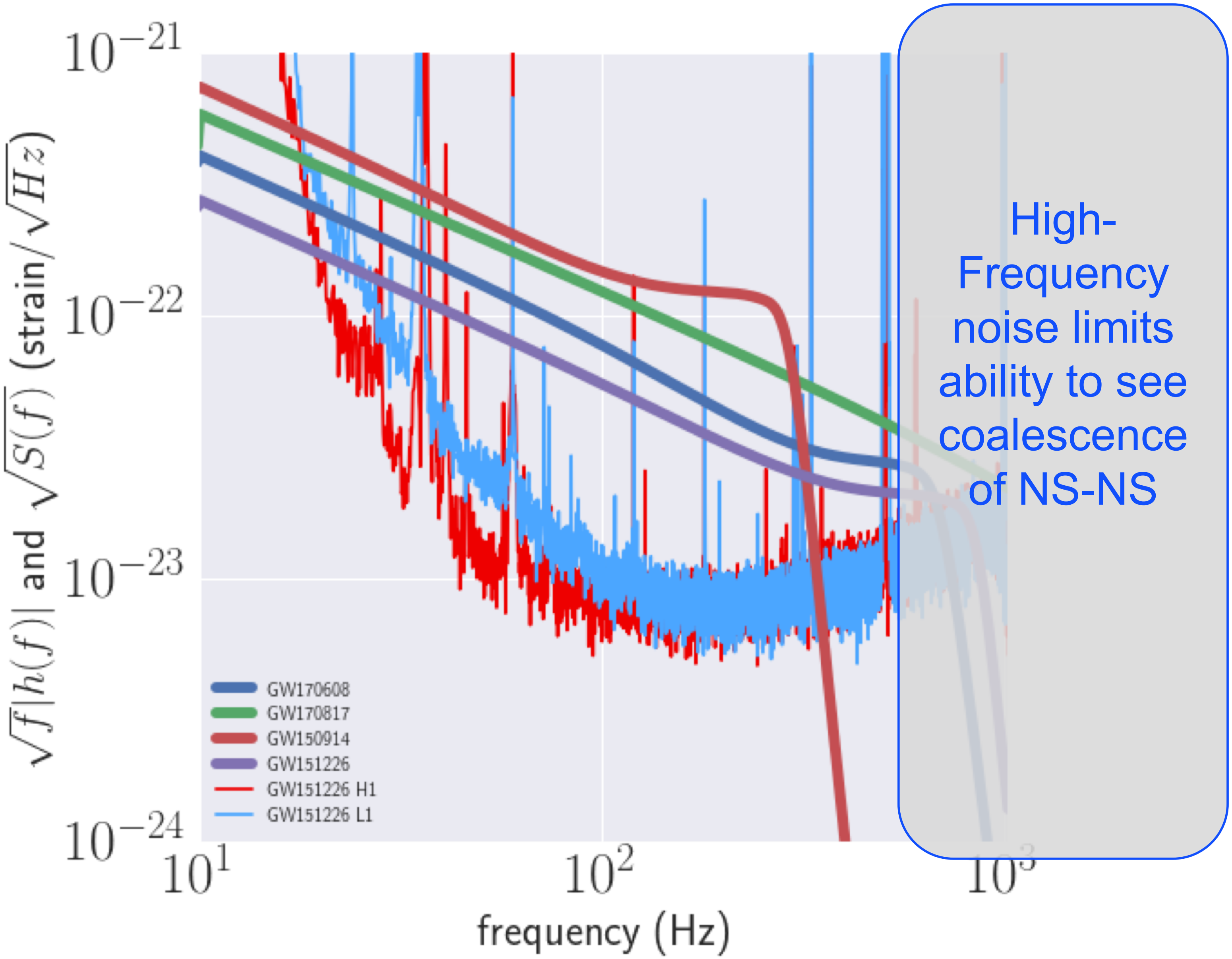


Evolution of detectors



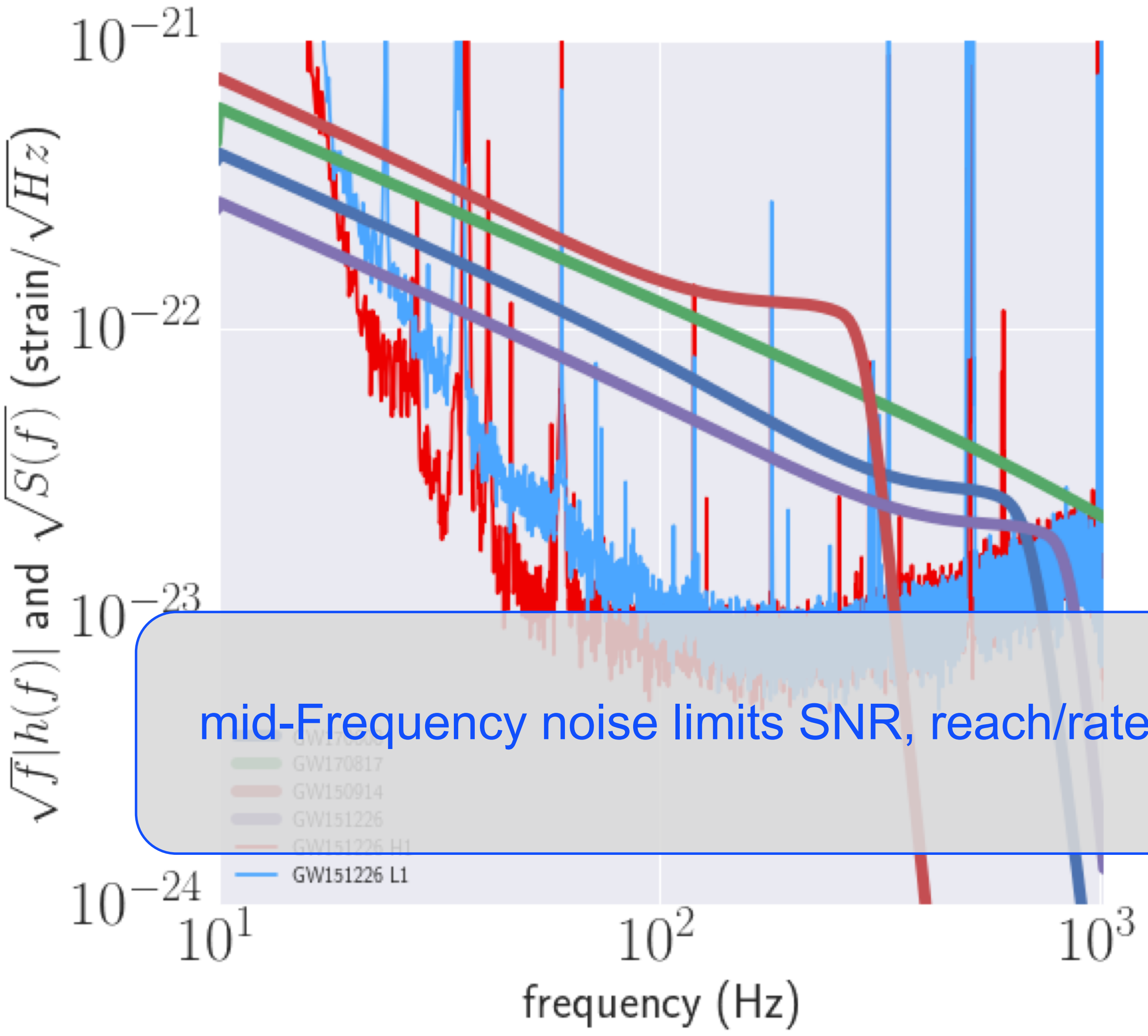
Low-Frequency noise limits largest masses, number of cycles

Evolution of detectors



High-Frequency noise limits ability to see coalescence of NS-NS

Evolution of detectors



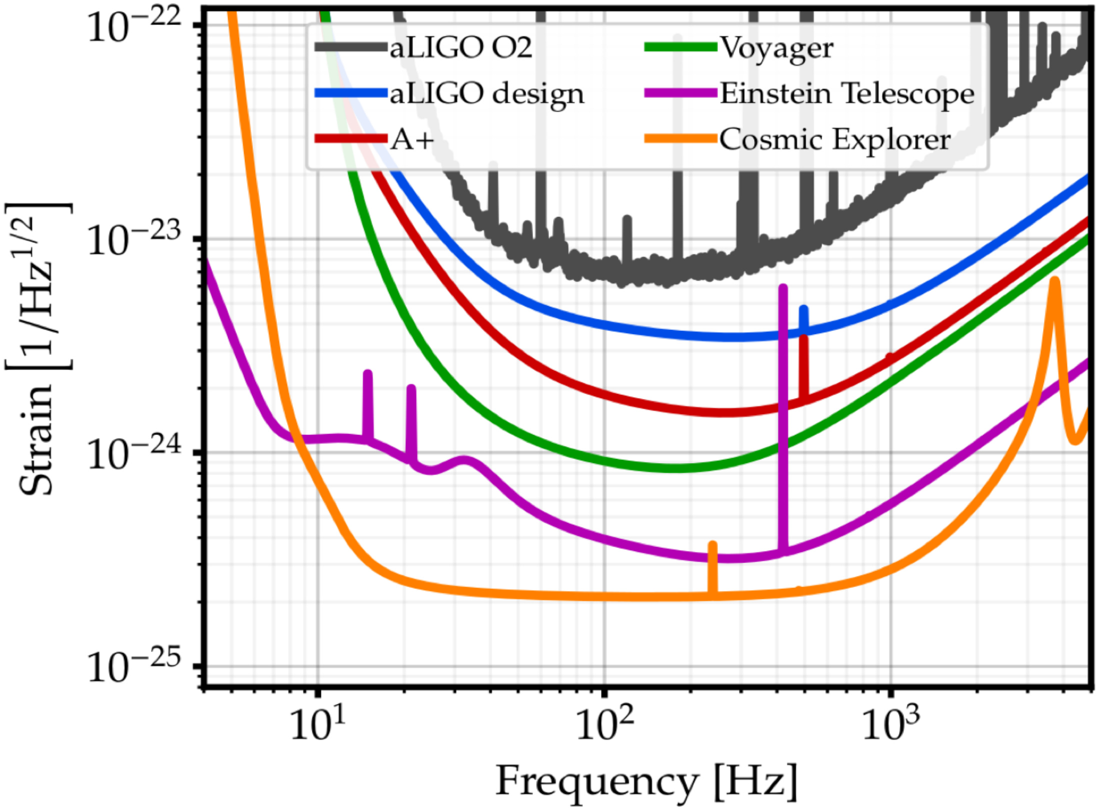
mid-Frequency noise limits SNR, reach/rates

Evolution of detectors

- aLIGO, AdV – commission to full sensitivity by early 2020’s
- A+, AdV+ – add squeezing, lower thermal noise coatings; ~2024
- Voyager – cryogenics to reduce thermal noise; ~2028
- ..at that point there is no choice but to seek longer arms

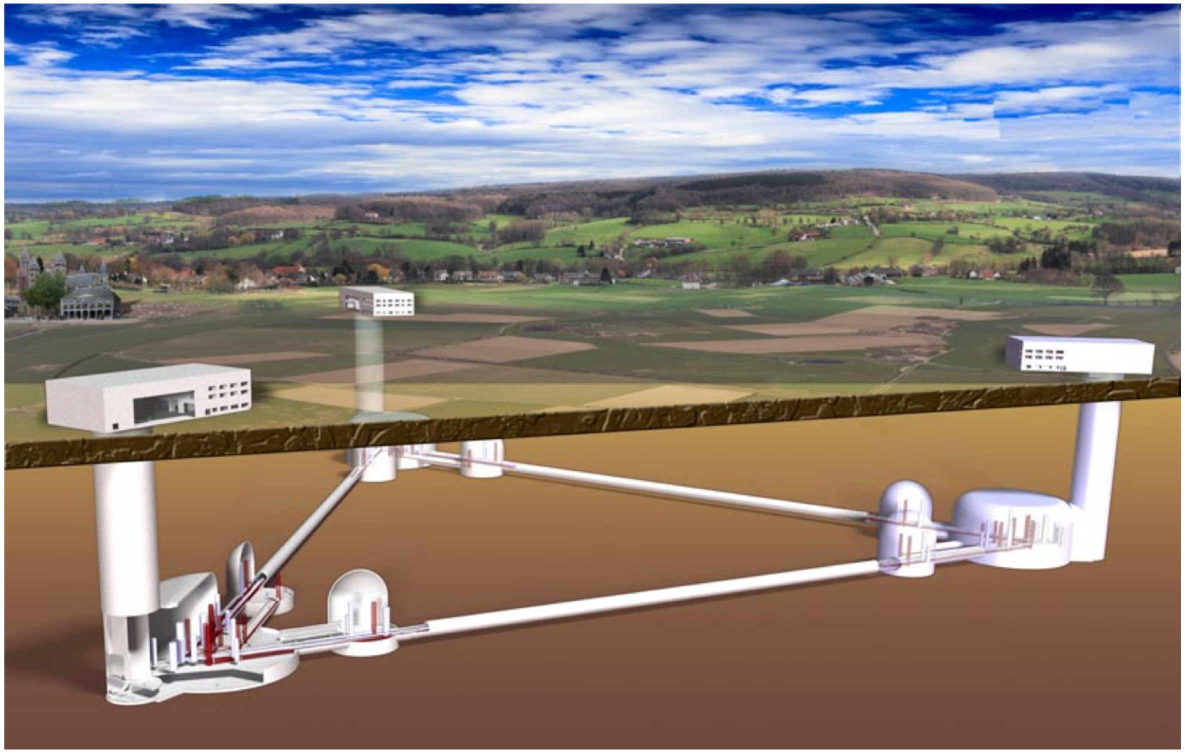
→ Einstein Telescope

→ Cosmic Explorer



Further Future Improvements: The 3rd generation

- One Concept: Einstein Telescope
- Significant design study undertaken for both Facility and Instruments
- Underground construction proposed to reduce Newtonian Background
  - (and be compatible with densely-populated Europe)
- Triangle – LISA-like – with 10km arms
- Multiple instruments in a ‘Xylophone’ configuration
  - Allows technical challenges for low- and high-frequency to be separated
- Designed to accommodate a range of detector topologies and mechanical realizations
  - Including squeezing and cryogenics

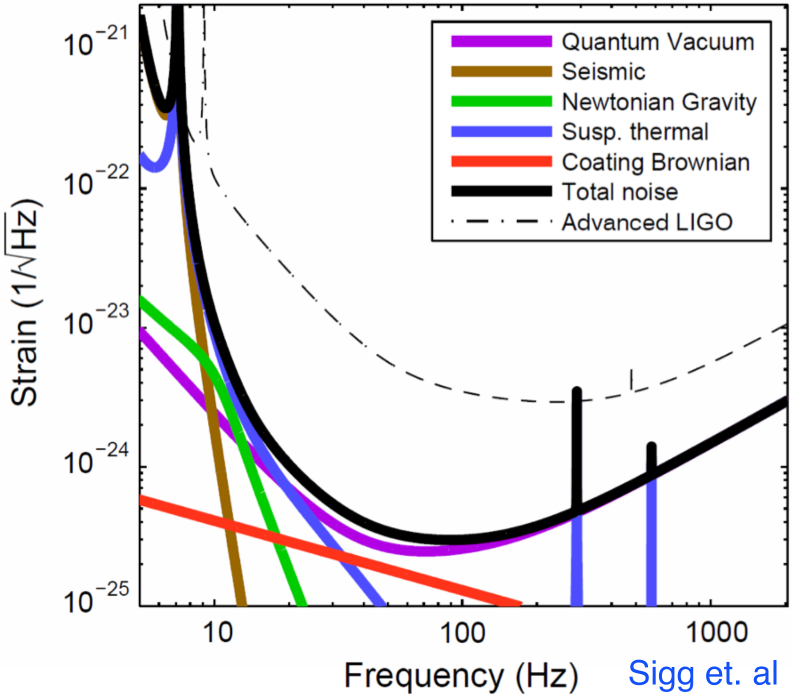


Another Concept:
Make Advanced LIGO 10x longer,
10x more sensitive

Signal grows with length – ***not*** most noise sources

- Thermal noise, radiation pressure, seismic, Newtonian unchanged
- Coating thermal noise improves faster than linearly with length
- 40km surface Observatory 'toy' baseline
  - can still find sites, earthmoving feasible; costs another limit…
- Concept offers sensitivity without new measurement challenges; could start at room temperature, modest laser power, etc.

	Adv. LIGO	**40 km LIGO**
Arm length	4 km	40 km
Beam radius	6.2 cm	11.6 cm
Measured squeezing	none	5 dB
Filter cavity length	none	1 km
Suspension length	0.6 m	1 m
Signal recycling mirror trans.	20%	10%
Arm cavity circulating power	775 kW	
Arm cavity finesse	446	
Total light storage time	200 ms	2 s



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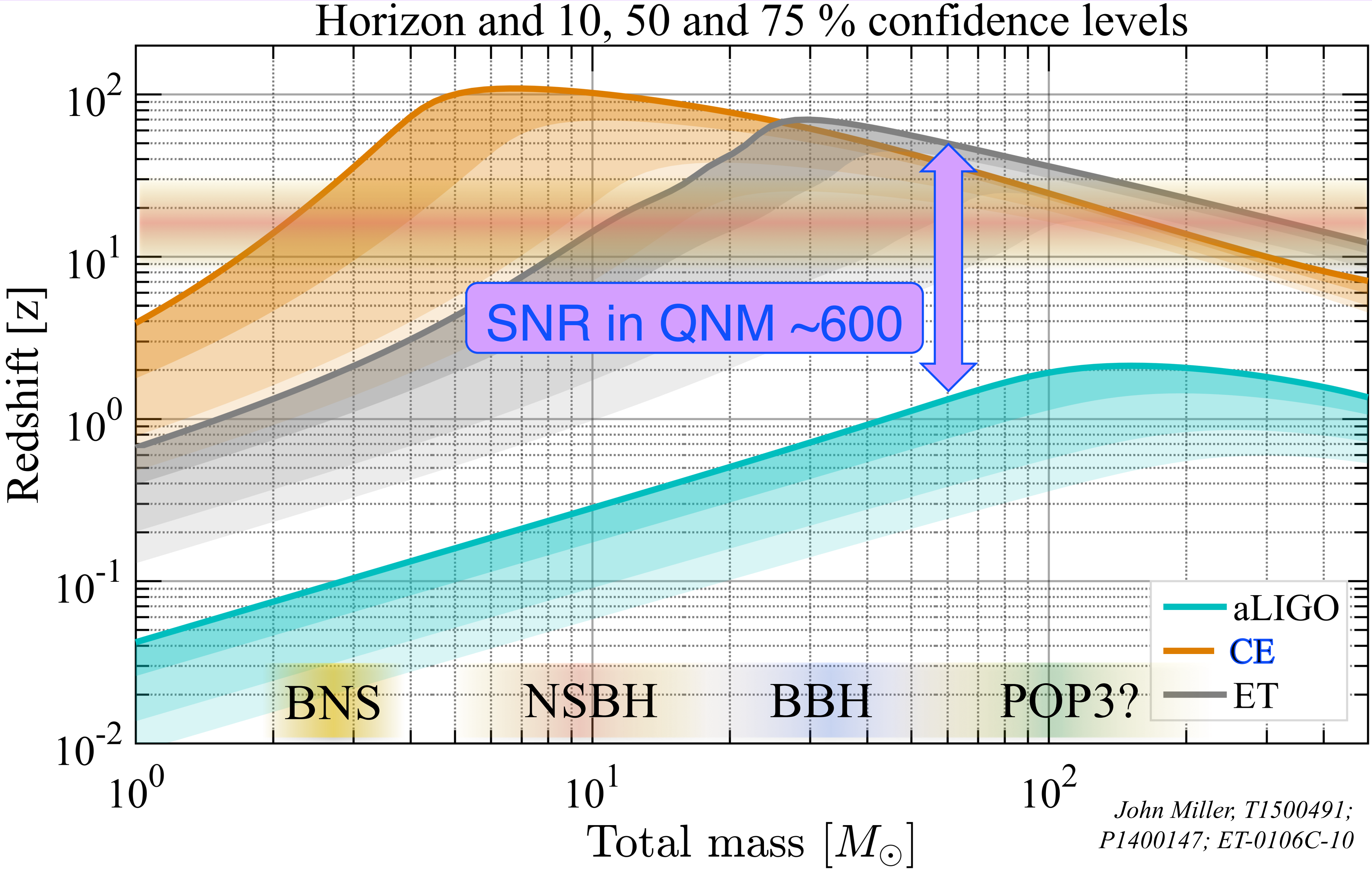
Einstein Telescope, Cosmic Explorer

‘Green field’ multi-generation Observatories ~G$/G€



Horizon and 10, 50 and 75 % confidence levels

SNR in QNM ~600

John Miller, T1500491; P1400147; ET-0106C-10

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3rd Generation

- When could this new wave of ground instruments come into play?
- Appears 15 years from *t*=0 is a feasible baseline
  - Initial LIGO: 1989 proposal, and at design sensitivity 2005
  - Advanced LIGO: 1999 White Paper, GW150914 in 2015
- **Modulo funding, could envision 2030's**
- Should hope – and strive and plan – to have great instruments ready to 'catch' the end phase of binaries seen in LISA (ref. Sesana)
- Worldwide community working together on concepts and the best observatory configuration for the science targets

- **Crucial for all these endeavors: to expand the scientific community planning on exploiting these instruments far beyond the GR/GW enclave**
  - Costs are like TMT/GMT/ELT – needs a comparable audience
  - Events like GW170817 help!

Just the beginning of a new field – new instruments, new discoveries, new synergies

Milliseconds	Minutes to Hours	Years to Decades	Billions of Years
LIGO/Virgo	LISA	Pulsar Timing Array	Cosmology Probes