

Effect of Thermal Lensing on WFS (FFT Study)

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- ASY port:
- ➢ WFS1Q ← DETM (differential ETM)
- Reflected port:
- > WFS3-I ← RM (recycling mirror)
- > WFS4-I ← CETM (common ETM)
- POX port
- ➢ WFS2a (I) ← CITM (common ITM)
- ➢ WFS2b (Q) ← DITM (differential ITM)





2 considerations for WFS Telescopes: [Nergis, T990130]

- ✓ The Gouy phase η_i be optimized for max sensitivity of primary dof
- ✓ The spot size is matched to the size of the quadrant diode (~3mm)

At WFS, the demodulation phase, ϕ_{Di} needs to be adjusted

Representation of Thermal Lensing



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- From FFT IFO (under different thermal states) get 128X128 pixelized data for beams (CR & SBs) at BS, POX, Reflected port at RM.
- Propagate these beams using Matlab "FFTprop" tool (<u>http://www.ligo.caltech.edu/~bbhawal/DOWNLOAD</u>) through telescope, distances and lens to WFS locations on output benches.
- Calculate signal in split-photodiode at WFS location
- Vary WFS location (so Gouy phase) to repeat "FFTprop" calculation



What is that Gouy Phase, η ?

4 dominant Parts of a WFS signal:

 $S_{1} = \overline{\mathrm{CR}}_{00} \overline{\mathrm{SBP}}_{01}^{*}$ $S_{2} = \overline{\mathrm{CR}}_{01} \overline{\mathrm{SBP}}_{00}^{*}$ $S_{3} = \overline{\mathrm{SBM}}_{00} \overline{\mathrm{CR}}_{01}^{*}$ $S_{4} = \overline{\mathrm{SBM}}_{01} \overline{\mathrm{CR}}_{00}^{*}$

Amplitudes \overline{CR}_{00} , \overline{SBP}_{01} , etc are without Gouy phases

Quadrature and Inphase Signals:

$$Q = 2\text{Re}\left[S_1e^{-i\eta} + S_2e^{+i\eta} + S_3e^{-i\eta} + S_4e^{+i\eta}\right]$$
$$I = -2\text{Im}\left[S_1e^{-i\eta} + S_2e^{+i\eta} + S_3e^{-i\eta} + S_4e^{+i\eta}\right]$$

$$\eta$$
 : Gouy phase between TEM00 & TEM01 at output port

All such terms contribute (ex: Pitch mode):

 $CR_{ij}SBP_{ik}^* + \dots$

where
$$abs(j-k) = an \text{ odd integer}$$

Define
$$\begin{split} S_j &= R_j + i * I_j \\ Q &= 2[(R_1 + R_2 + R_3 + R_4)^2 + (I_1 - I_2 + I_3 - I_4)^2]^{1/2} \\ \eta_Q &= tan^{-1} \Bigg[\frac{(I_1 - I_2 + I_3 - I_4)}{(R_1 + R_2 + R_3 + R_4)} \Bigg] \\ I &= -2[(I_1 + I_2 + I_3 + I_4)^2 + (R_2 - R_1 + R_4 - R_3)^2]^{1/2} \\ \eta_I &= tan^{-1} \Bigg[\frac{(R_2 - R_1 + R_4 - R_3)}{(I_1 + I_2 + I_3 + I_4)} \Bigg] \end{split}$$



WFS1Q (C DETM) Signals





- NRSB demodulation: (RM → WFS3I,CETM→WFS4I)
 - 1) CR with \pm NRSB ($\propto \Gamma_{\rm NRSB}$)
- (only for $RM \rightarrow WFS3-I$)
 - 2) Upper RRSB with [RRSB \pm NRSB] $(\propto \Gamma_{\rm RRSB}^2 \Gamma_{\rm NRSB})$
 - 3) Lower RRSB with [-RRSB \pm NRSB] $(\propto \Gamma_{\rm RRSB}^2 \Gamma_{\rm NRSB})$



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WFS 3-I (← RM) & 4-I (← CETM) Signals



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WFS2a (CITM) Signals from POX





WFS2b (CONTM) Signals from POX



Common and Extreme Differential Heating



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Effect of Thermal Lens on Demod. Phase of WFS2



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LIGO Demodulation Phase & Extreme Differential Heating



LIGO Sidebands and Extreme Differential Heating



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(G040442)

LIGO Change of Beam-width in LLO:WFS1 path



LIGO Change of Gouy Phase in LLO: WFS1 path





[Related to this Study only]

- Why is LLO WFS working even though some of its Gouy phase (empirically optimized) locations are not the right places to get strong signals?
- Why NO changes in demodulation phase is observed in WFS2 at LLO?
- 0 ...